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Extended Ionized Fe Objects in the UWIFE Survey

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ABSTRACT

We explore systematically the shocked gas in the first Galactic quadrant of the Milky Way using the United Kingdom Infrared Telescope (UKIRT) Wide-field Infrared Survey for Fe⁺ (UWIFE). The UWIFE survey is the first imaging survey of the Milky Way in the [Fe II] 1.644 μ m emission line and covers the Galactic plane in the first Galactic quadrant (7° < l < 62°; $|b| \leq 1.5$). We identify 204 extended ionized Fe objects (IFOs) using a combination of a manual and automatic search. Most of the IFOs are detected for the first time in the [Fe II] 1.644 μ m line. We present a catalog of the measured sizes and fluxes of the IFOs and searched for their counterparts by performing positional cross-matching with known sources. We found that IFOs are associated with supernova remnants (25), young stellar objects (100), HII regions (33), planetary nebulae (17), and luminous blue variables (4). The statistical and morphological properties are discussed for each of these.

Key words: circumstellar matter - catalogues - surveys - infrared: ISM - ISM: kinematics and dynamics

1 INTRODUCTION

Understanding the role of shocks is essential for comprehensively studying the ecology of the Milky Way, since they transfer mass and kinetic energy into the interstellar medium (ISM), provide heavy elements for future star formation by destroying dust grains, and regulate star formation. Shock waves are one of the principal mechanisms of the interaction between stars and the ISM in galaxies, thereby playing an important role in the evolution of the Galaxy. Among the most powerful shock-driving sources are outflows and jets from young stellar objects (YSOs), stellar winds from massive OB stars, and supernova (SN) explosions. To understand the physics of the interactions as well as the nature of the shock-driving sources, observations of emission lines from the shocks are essential.

The [Fe II] $a^4 D_{7/2} \rightarrow a^4 F_{9/2}$ 1.644 μ m transition results in one of the brightest emission lines in near-infrared (NIR). It originates from one of the 16 levels of Fe⁺ that have a low excitation energy. Therefore they are easily excited in shocked gas, resulting in many lines, particularly in NIR. This emission line is thought to be bright in shock-excited gas; one suggested reason is that due to far-ultraviolet (FUV) radiation from the shock front, the Fe atom is in the form

of Fe⁺ over extended regions (McKee et al. 1984, Hollenbach et al. 1989, Oliva et al. 1989, Koo et al. 2016). In contrast, in photoionized regions. Fe atoms are predominantly at higher ionization states. except when the ionizing radiation is hard enough that it can penetrate further into the interstellar cloud (Koo et al. 2016). Therefore, [Fe II] emission lines from shocked gas are stronger than those from photoionized regions; for example, [Fe II] 1.257 μ m / Pa β is over 0.1 in supernova remnants (SNRs) compared to 0.01-0.03 in Orion (Koo & Lee 2015; Mouri et al. 2000). Furthermore, the Fe abundance can be enhanced by shocks owing to grain destruction, making the [Fe II] lines stronger (Koo 2014; Greenhouse et al. 1991; Mouri et al. 2000 and references therein). These characteristics of [Fe II] make its lines extremely useful for studying interstellar shocks (e.g., Dinerstein 1995; Nisini 2008).

For example, the 1.644 μ m emission line as a tracer of shocked atomic gas enables us to study shocked regions in jets/outflows of YSOs (Nisini et al. 2002; Caratti o Garatti et al. 2006; Takami et al. 2006; Pyo et al. 2006; Pyo et al. 2009; Oh et al. 2016), planetary nebulae (Welch et al. 1999; Smith et al. 2005), supernova remnants (Koo et al. 2007; Lee et al. 2009; Lee et al. 2013), and nebulae of luminous blue variables (Smith 2002). Since NIR [Fe II] lines suffer less extinction than widely used optical emission lines such as $H\alpha$,

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[S II], and [O III], the NIR lines can give us information on deeply embedded regions inaccessible by optical lines.

Lee et al. (2014) conducted an unbiased [Fe II] 1.644 μ m narrowband imaging survey, which is called the United Kingdom Infrared Telescope (UKIRT) Wide-field Infrared Survey for Fe⁺ (UWIFE). The survey area $(7^{\circ} < l < 62^{\circ}; |b| \le 1^{\circ}, 5)$ is located in the first Galactic quadrant. This survey is the first unbiased, high-resolution [Fe II] survey of the Milky Way. It therefore enables us to discover more [Fe II]emitting sources and conduct a statistically meaningful investigation of Galactic [Fe II] line sources. Alongside [Fe II]-emitting Galactic SNR to study similar to [Fe II] line objects in nearby galaxies, the survey is expected to systematically detect low-brightness [Fe II] line sources from other kinds of [Fe II] emitters. Therefore it enables us to assess the level of contribution of each [Fe II]-emitting population. Further spectroscopic studies of new [Fe II] sources found in UWIFE can be used to derive critical densities in the range of $\sim 10^4 - 10^5$ cm^{-3} and temperatures up to 10^4 K (Pesenti et al. 2003), filling the gap in density between [S II] $\lambda 6731 \sim 10^4$ cm⁻³ and [O I] $\lambda 6300 \sim 10^6$ cm⁻³ (Osterbrock 1989). With other [Fe II] lines and emission lines such as [Fe II] 1.533 μ m, density diagnostics of ~10²-10⁵ cm⁻³ can be measured and line ratio diagrams with other [Fe II] lines (Pesenti et al. 2003) can help us understand the new parameter range.

Shinn et al. (2014) systematically searched for outflows from ultra-compact HII regions (UCHIIs), inferred [Fe II] outflow massloss rates, and discussed the travel time of the [Fe II] outflows using the UWIFE data. The statistical [Fe II] line study of Galactic SNRs in UWIFE and the UKIRT Wide-field Infrared Survey for H₂ (UWISH2; Froebrich et al. 2011) survey revealed a detection rate of 24 per cent for both surveys and suggested a relatively higher coincidence with mixed-morphology and/or radio-bright SNRs (Lee et al. 2019).

A comprehensive catalog of UWIFE sources will give an opportunity to compare shocked [Fe II] line objects with other tracers in previous large-scale Galactic plane surveys. Particularly, the UWIFE survey area is fully covered with the complementary survey, UWISH2 (Froebrich et al. 2011), which was carried out using UKIRT and the Wide-Field Camera (WFCAM, Casali et al. 2007). The catalog of extended H₂-emitting sources identified in UWISH2 (Froebrich et al. 2015) will be useful for the comparison of shocked molecular gas with higher excitation atomic gas. Also, the Isaac Newton Telescope (INT) Photometric H α Survey of the Northern Galactic Plane (IPHAS; Drew et al. 2005) and the UWISH2 survey can provide a chance to compare different outflow/shock tracers. Surveys tracing continuum sources in embedded regions such as the UKIRT Infrared Deep Sky Survey (UKIDSS) Galactic Plane Survey (GPS; Lucas et al. 2008) in the near-infrared, the Galactic Legacy Infrared Midplane Survey Extraordinaire (GLIMPSE; Churchwell et al. 2009), the Multiband Imaging Photometer for Spitzer Galactic Plane Survey (MIPSGAL; Carey et al. 2009) in the mid-infrared (MIR), and the Herschel infrared Galactic Plane Survey (Hi-GAL) (Molinari et al. 2010) in the far-infrared (FIR) were published.

Furthermore, the source catalog of various kinds of objects, namely, the catalog of UCHIIs from the Co-Ordinated Radio 'N' Infrared Survey for High-mass star formation (CORNISH, Hoare et al. 2012) and the catalog of Extended Green Objects (EGO, Cyganowski et al. 2008) can be good candidates to compare with [Fe II] sources, as well as emission line source catalogs, viz., catalogs of H α emission-line sources from IPHAS (Witham et al. 2008), and Molecular Hydrogen emission-line Object (MHO, Davis et al. 2010). In accordance with these catalogs and aims, we designate [Fe II] 1.644 μ m emission-line sources as ionized Fe objects (IFOs) and compile the first comprehensive catalog of Galactic extended

IFOs. The catalog includes basic physical properties of IFOs, such as coordinates (l, b), size, position angle, and flux. Information about possible counterparts and their distance is also included.

The organization of the paper is as follows. In Section 2, we describe the data reduction, the source identification, the photometry of the detected sources, and the procedure for searching for counterparts or exciting sources of IFOs. In Section 3, we first present the catalog of IFOs. The catalog contains the sizes and fluxes of IFOs as well as their counterparts. The IFOs are classified by their counterpart types. We then explore the statistics of the physical properties and the distribution of IFOs. In Section 4, we discuss the results of the individual types of IFOs. In Section 5, we summarize our paper.

2 DATA AND SOURCE CATALOG

2.1 UWIFE Survey Data

We have used the UWIFE survey data to search for extended IFOs in the Galactic plane. The UWIFE survey was carried out using WFCAM at UKIRT in 2012 and 2013 (Lee et al. 2014). The [Fe II] narrow-band filter was used, having a central wavelength of 1.644 μ m and an effective bandwidth of 0.026 μ m. The WFCAM hosts four Rockwell Hawaii-II HgCdTe 2 k × 2 k arrays, each covering 13.65 arcmin × 13.65 arcmin in area at a pixel scale of 0.4 arcsec. Four pointings of the telescope covered a contiguous area of 0.75 deg^2 (designated as 'tile', following the WFCAM terminology). Each pointing was composed of a set of dithered and microstepped observations, fully sampling the point spread function in good seeing conditions (<0.8 arcsec). The total integration time per pixel was 720s. The final [Fe II] images have a nominal 5σ detection limit of 18.7 mag for point sources, with a median seeing of 0.83 arcsec. For extended diffuse sources, the corresponding surface brightness limit is 8.1×10^{-20} W m⁻² arcsec⁻².

Lee et al. (2014) also produced continuum-subtracted [Fe II] images (hereafter [Fe II]-H images) by using the *H*-band images from the GPS. The continuum subtraction was carried out in two steps, i.e., point-like continuum sources were first removed in both [Fe II] and *H*-band images, and then the point source removed *H*-band images were subtracted from the point source removed [Fe II] images to remove extended continuum sources. The details of the observation and data processing procedure can be found in Lee et al. (2014).

All [Fe II] and [Fe II]-H images from UWIFE are available at the UWIFE web page ¹. The images consist of 220 tiles, where a single tile is a square of $54' \times 54'$ in equatorial coordinates. The tiles are arranged as 55 stripes of four consecutive tiles at constant declination along the Galactic plane, covering a region within the First Galactic Quadrant of $7^{\circ} < l < 62^{\circ}$; $|b| \lesssim 1^{\circ}5$ (see figure 1 of Lee et al. 2014). On the web page, the UWISH2 H₂ and GPS *JHK*-band images are also available.

2.2 Source Identification

In this study, we first aimed to identify IFOs in the continuumsubtracted images (hereafter, [Fe π]-H). We identified most of the IFOs through visual inspection and added several faint IFOs by mean of an automatic source identification, which uses the same algorithm as UWISH2 (Froebrich et al. 2015).

We focused on the extended sources in this study. Visual inspections were carried out twice for the whole survey area. We searched for all possible emission features and, for each feature, we defined an ellipse approximately surrounding the emitting area by eye. Then the central coordinates, radii, and P.A. of the ellipses are measured and listed in Table 1. All IFO candidates identified in the [Fe II]-H image were double-checked in both un-subtracted [Fe II] and GPS *H*-band images to confirm whether they were a real source or not. As the UWIFE and GPS observations were separated by several years, variable sources were seen as emission or absorption in the [Fe II]-H images. In particular, since artifacts with a negative digital number (DN) in GPS *H*-band resemble real sources in the [Fe II]-H image, we checked the position of all IFO candidates in the corresponding *H*-band data.

In addition, there are various kinds of artifacts in the [Fe II]-H images. Representative artifacts are: the residuals of bright stars, ghosts, cross-talks, cross stripes after star subtraction, and the diffraction pattern of bright stars (see Appendix A.1 for examples). Residuals of high proper-motion stars were also left in the [Fe II]-H images. We also excluded the features hampered by the artifacts from bright stars. The sources that show diffuse structures in both [Fe II] and H which are significantly brighter in [Fe II] compared to the GPS H-band, or the sources with a low probability of being scattered emission from dust seen in the GPS H-band, were selected as real sources.

Using the [Fe II]-H images, we conducted an unbiased automatic detection with the code used for identifying MHOs in UWISH2 (Froebrich et al. 2015) to benefit from its objectiveness. We adjusted the code to fit the specifications of UWIFE data: (1) Remove smallscale features (residual of star subtraction), determine the large-scale background level from a 40 arcsec scale median filter, and calculate its noise value. (2) Draw contours at the 1σ level in ds9² and identify the isolated contours as 'regions'. The level was determined empirically to include faint emission of IFOs. The low (1σ) level produces contours around the remaining point sources and noise peaks, but those 'false' regions are removed by a minimum size limit in the next stage. (3) Remove contours that are too small ($<4 \text{ arcsec}^2$) or near the image borders. (4) To avoid mistakenly identifying star residual as IFOs, remove contours smaller than 35 arcsec^2 if they are located within 3 arcsec to the Two Micron All-Sky Survey (2MASS) H-band stars brighter than 15 mag. This procedure was conducted for all UWIFE tiles except for tile 003, 080, and 196 due to the late release of the H-band data in the GPS survey.

All sources identified by the automatic detection were cross checked by visual inspection. We first examined whether the identified source from the code is an image artifact or not. Appendix A.1 shows some examples of the artifacts, including residuals of detector cross-talk and diffraction patterns from saturated stars. These nonastronomical sources can be easily distinguished by comparing them in the [Fe II] and H-band images and were removed from the catalog. We also rejected point-like sources (e.g., high proper motion stars, variables, [Fe II]-emitting stars, etc.) that are not considered in this paper. Note that the visual identification treats a group of clumpy structures as a single object (e.g., shells of SNR). On the other hand, the code identifies the substructures separately. We fitted each automatically-identified IFO with an ellipse and derived the geometrical parameters of the semi-major axis, semi-minor axis, and position angle. This process added 14 IFOs, and the complete catalog is presented in Table 1, which also provides their coordinates, sizes, fluxes, and counterparts.

2.3 Photometry

We conducted photometry of the IFOs in the [Fe II]-H images. Since our targets have an extended structure, we adopted aperture photometry. In the [Fe II]-H data, artifacts often have higher digital counts than IFOs. Therefore, masking artifacts is a crucial process. We masked the identified artifacts to prevent large uncertainties in the aperture photometry. The residuals of point sources (stars) brighter than 14th magnitude in the *H*-band (based on the 2MASS point source catalog, Cutri et al. 2003) were also masked. The size of the masking area was 6 arcsec in diameter, which is large enough to cover general residual patterns. When instrumental artifacts such as electronic cross-talk or diffraction patterns intruded on the aperture, we manually masked them to prevent any contamination.

In order to derive the total flux (F_{tot}) of the identified IFOs in a scientific unit (W m⁻²), we used the following equation :

$$F_{\text{tot}} = F_0 \cdot \left(\frac{DN}{t_{\text{exp}}}\right) \cdot 10^{-0.4 \cdot m_{\text{zpt}}}$$

 F_0 is the in-band flux of Vega falling in the [Fe II] filter (3.27 × 10^{-11} W m⁻², Lee et al. 2019), whereas t_{exp} and m_{zpt} are the net exposure time (60 s) and the zero-point magnitude of each image, respectively ('EXPTIME' and 'MAGZPT' in the image header). DN is the total digital number falling in the aperture corrected for the sky background. This local background of each source was estimated from a sky annulus with an inner and outer radius of 1.2 and 1.5 times the aperture. We took the mode of the sky values to further avoid the effect of any possible artifacts. The uncertainty of the flux is estimated considering the photometric calibration error from the uncertainty of the zero-point magnitude of ~0.06 mag, which corresponds to ~6 per cent of the total flux (Lee et al. 2019). The contribution of Poisson noise from aperture photometry and sky subtraction is negligible. The former is less than one-fifth, and the latter is less than one-tenth of absolute calibration uncertainty.

2.4 Search for Associated Exciting Sources

We have searched for the possible driving source(s) of IFOs via positional cross-matching with previously known sources: SNRs, HII regions, compact and ultra-compact HII regions, luminous blue variables (LBVs), planetary nebulae (PNe), and YSOs. IFOs associated with these sources are classified as SNR-IFO, HII-IFO, CHII-IFO, LBV-IFO, PN-IFO, and YSO-IFO, respectively. The rest of the IFOs are classified as 'unknown-IFO'. In the following, we describe the processes and catalogs employed for the search for the individual exciting source types.

SNRs have complex and filamentary structures often with a large spatial extent. Thus, a careful identification and the separation of genuine SNR-origin from mere superposition was required. We first selected IFOs located within the boundary of known SNRs, using the central positions and sizes of SNRs in the Galactic SNR catalog of Green (2019). We then referred to the references in the catalog and also SIMBAD³ for the multiwavelength morphology of SNRs for the confirmation of the association. If an IFO shows a coherent structure occupying a similar extent and/or its morphology implies a spatial correlation with the SNRs, we categorised it as an SNR-IFO. We also checked the area in SIMBAD for a possible superposition

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<sup>2</sup> https://sites.google.com/cfa.harvard.edu/saoimageds9
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of unrelated, superposed sources such as PNe along the same line of sight. An IFO without noticeable morphological correlation but positionally coincident with evident SNR emission was categorised as an SNR-IFO (e.g., IFO 38).

Diffuse HII regions also occupy a large spatial area and have complex morphology, so that a SIMBAD/VizieR query by IFO coordinate with an arcmin radius often returns various kinds of incidental sources such as sub-filaments of HII regions, jets/outflows from neighboring YSOs, and merely superposed sources along the line of sight. Therefore, keeping in mind that proximity alone does not necessarily guarantee a true correlation, a morphological correlation was also taken into account for identifying HII regions as a exciting source. If needed, a query with a larger angular scale was conducted to locate the diffuse HII region. We then compared the morphology of IFOs with that of HII regions obtained from high-resolution radio images (GPS; White et al. 2005, New-GPS; Helfand et al. 2006, and the H I, OH, recombination line survey of the Milky Way, THOR continuum; Beuther et al. 2016) and/or datasets from large-scale multi-wavelength studies (Fujita et al. 2021; Povich et al. 2009; Roshi et al. 2017). We also used small-scale surveys and targeted studies (see § 4.3). The IFOs with a positive correlation have been categorized as HII-IFOs. However, since the [Fe II] line emission from an HII region is inherently faint, morphological correlation with radio is occasionally hard to confirm. On the basis of this possibility, a few extended and faint IFOs have also been regarded as HII-IFOs although they do not have a clear morphological relationship with an HII region (see comments in \S 4.3).

We further explored whether IFOs are associated with compact, ultra-compact, or hyper-compact HII regions (CHII, UCHII, and HCHIIs) by querying VizieR within an arcmin-scale radius. Two comprehensive lists of UCHII regions were selected for the VizieR positional matching: the CORNISH UCHII region catalog (Kalcheva et al. 2018), which is appropriate for the comparison with the UWIFE survey data in terms of comparable high-resolution (1.5) and spatial coverage $(10^{\circ} < l < 65^{\circ}, |b| < 1^{\circ})$, and the catalog presented by Bronfman et al. (1996) which is a large-scale compilation of Infrared Astronomical Satellite (IRAS) FIR color-selected UCHII regions with higher Galactic latitude coverage ($|b| < 2^{\circ}$). The IFO positions were subsequently searched in SIMBAD to refer to targeted studies. We compared the [Fe II] line morphologies with available radio continuum images (see \S 4.2). When the IFOs show morphological correlation with radio structures or delineate the boundary of radio structures, they are classified as CHII-IFOs. IFOs having counterparts supposedly earlier or at a lower-mass evolutionary stage of an UCHII region (e.g., hypercompact HII region, UC HII precursor, ultra-compact embedded cluster which was suggested as a lowermass class of UCHII) are also included in this category (see § 4.2).

In order to identify IFOs associated with luminous blue variables (LBV-IFOs), the SIMBAD query was conducted with a radius criterion of 10 arcmin. But we noted that the list of LBVs and LBV candidates (hereafter cLBVs) has not been fully incorporated in SIMBAD, so we also used the catalog of (c)LBVs compiled by Nazé et al. (2012) which lists the coordinates of 68 (c)LBVs. As far as we know, this is the most comprehensive catalog of (c)LBVs. For example, Weis & Bomans (2020), in their review article of LBVs, presented a catalog, but it has a smaller number of (c)LBVs than Nazé et al. (2012), i.e., 47 versus 68. In the catalog of Nazé et al. (2012), twenty-two LBVs (including 19 candidates) are located inside the UWIFE area. There was also a [Fe II] survey of 9 LBVs by Smith (2002). Among the 9 LBVs, only one was located within the UWIFE survey area and it has been identified in our survey, too.

For IFOs associated with PNe (PN-IFOs), the SIMBAD query

was used with a radius criterion of 10 arcmin. We additionally compared the morphology of IFOs with multiwavelength data from references in SIMBAD. In order to incorporate recently discovered PNe and PN candidates that have not been updated in SIMBAD, we made use of the following databases and catalogs. We used the Hong Kong/AAO/Strasbourg H α (HASH) planetary nebula database which lists multi-wavelength data of newly found ~3500 PNe and PN candidates (Parker et al. 2016). The database includes three largescale catalogs of Galactic PNe; the Strasbourg-ESO catalog, the catalog of Galactic Planetary Nebulae version 2000, and the Macquarie/AAO/Strasbourg H α (MASH) catalogs, together with 159 new PNe from the related IPHAS survey and ~400 from the literature. A large number of unpublished, new PN candidates are accessible in this database, which are mostly (1) older, redder, and have lower surface brightness or (2) are more remote and small-scaled, faint PNe (Parker et al. 2016). When the counterpart is not a well-known source and is identified only in the HASH database, we checked the association using the references provided. There are PNe only detectable in NIR, so the recent study of PNe based on the UWISH2 data (Gledhill et al. 2018) was also checked for possible counterparts. This study complements faint or small-scale PNe previously undiscovered.

For the remaining IFOs, we made use of several large-scale catalogs of YSOs alongside catalogs for specific regions or targets. The large-scale survey of YSOs in four evolutionary stages (i.e., quiescent, YSO, protostellar, and massive star-forming stages, Urguhart et al. 2018) was used to find YSO-IFOs, keeping in mind the survey resolution (30 arcsec). The Infrared Array Camera (IRAC) red-source catalog was also used in the same manner (Robitaille et al. 2008) to locate YSO-candidate counterparts. When there was a positive match, we subsequently displayed their positions on the [Fe II]-H images with H₂ contours of UWISH2 data to confirm their association. H₂ images are useful since H2 emission is usually more easily excited, forming a series of knots between an IFO and the YSO that drives a H₂ outflow. When the positional match and morphological information could not pinpoint an obvious YSO counterpart, we listed up to two YSOs. Also used are small-scale survey catalogs to benefit from a deeper searches for YSOs. Kim et al. (2015) conducted a detailed survey of YSO candidates in the infrared dark cloud (IRDC) G53.2 region and investigated their evolutionary stages. Povich & Whitney (2010) investigated the M17 region where we have identified many IFOs, and the study provided the evolutionary stages of the YSOs. Ragan et al. (2009) covered multiple IRDC regions in the UWIFE survey area and a YSO Class with MIR color and distance information was provided. Other small-scale catalogs of YSOs available in Vizier were also used when available (\S 4.1).

Since Herbig-Haro (HH) objects are often bright in [Fe II] emission, we attempted to locate the [Fe II]-emitting HH objects separately from YSO-IFOs. 454 Galactic HH objects have been compiled by Reipurth (2000), who continuously updated the SIMBAD database to include newly found HH objects. We retrieved all HH objects in SIMBAD, up to HH 1213, which includes 3140 sub-structures (e.g., HH 250A and 250B). Firstly, we search for YSO-IFOs and unknown-IFOs within a radius criterion of 10 arcmin for a given HH object. When there was a match, we looked for a possible association of the IFO with HH object structures via multi-wavelength images (mainly $H\alpha$ from IPHAS, Witham et al. 2008). For example, IFO 195 which is associated with the parsec-scale HH 803 has a very compact, small-scale structure. It was originally categorized as an unknown-IFO since we could not find any associated source just based on positional proximity. However, when we plot the IPHAS H α and UWISH2 molecular hydrogen emission contours together, we could associate IFO 195

with the south-western tip of the series of aligned structures of HH 803 in H α and H₂ emission. Figure 1 shows the representative IFOs with respect to each counterpart.

3 RESULTS

3.1 Catalog of IFOs and their Statistical Properties

The full catalog of extended IFOs is presented in Table 1 and the description of each column of the catalog is as follows:

Column 1. IFO identifier : Designations of IFOs by a catalogue number in ascending order. When the IFO is identified only by a source detection algorithm, we marked them with an 'a' after its designation.

Column 2. IFO conventional designation : IFO full-name derived from Right Ascension and Declination (J2000) of the source center. It follows the 2MASS convention for the naming, i.e.: IFO JHH:MM:SS.SSS±DD:MM:SS.SS.

Column 3, 4. Galactic longitude (l) and latitude (b) : The center position of the source, in units of degree, in Galactic coordinates. For automatically identified IFOs, we adopt the geometric centre of the polygon by 2-dimensional Gaussian fitting of an ellipse.

Column 5, 6. Semi-major axis (r_1) and semi-minor axis (r_2) : Maximum semi-major and minor angular radius of the IFO in units of arcseconds.

Column 7. Position angle (P.A.) : The angle of the semi-major axis of an ellipse, in a counter-clockwise direction, from North to East in units of degree.

Column 8. Area : An area of an ellipse determined by the semimajor axis and semi-minor axis (Column 5, 6), in square arcseconds.

Column 9. F_{tot} : Total flux derived from summing up all flux inside an ellipse drawn from columns 5, 6. See the photometry section § 2.3 for details.

Column 10. Counterpart : Classification of the IFO indicating the most probable known object as follows: YSO-IFO - outflows or jets from an YSO or YSO candidate, HII-IFO - any outflows surrounding emission originated from the HII region, subdivided into HCHII, UCHII, CHII, and diffuse HII region, SNR-IFO - emission originates in SNR, PN-IFO - emission associated with PN/PN candidates, further classified into PN, PNc, and post-asymptotic giant branch (AGB), LBV-IFO - nebula structure around an LBV or LBV candidate, Unknown-IFO - multiple corresponding known object candidates or no possible known source in the vicinity,

Our IFO catalog contains 204 sources identified from 219 tiles, which is about 180 deg² in total. This number corresponds to an average surface density of ~1.1 IFOs per deg² in the first quadrant of the Galactic Plane (GP). This number should be regarded as a lower limit since our source identification methods were conservative. In general, the results of the manual and automatic search by the source detection algorithm were in good agreement. The 14 sources found only by the source detection algorithm, are marked with an 'a' after the IFO number in Table 1. They were either very faint or resembled artifacts. The majority of catalog sources are new discoveries of [Fe II] emission, and represent an order of magnitude increase in the number of extended [Fe II] sources in the first Galactic quadrant.

Table 2 presents basic statistics of IFOs for each counterpart type. We identified 100 YSO-IFO (87 YSOs, 13 HHs), 33 HII-IFO (22 CHII, 11 HII), 25 SNR-IFO, 17 PN-IFO, 4 LBV-IFO, and 25 IFOs without counterparts. Note that if a counterpart source has two distinct [Fe II] structures, they are counted as two separated IFOs which share a common counterpart (e.g., IFO 85 and 86 are from SNR G21.5-0.9 and are counted as 2 SNR-IFOs). Also, one SNR-IFO (IFO 7) is included in the number statistics in Table 2 but not used for flux statistics.

In total, 65 per cent of identified IFOs are related to star formation (49 per cent YSO- and 16 per cent HII-IFO), and 22 per cent are associated with evolved objects with 12 per cent of IFOs remaining as of unknown origin. Among them, YSO-IFO is the most frequent population showing [Fe II] emission. However, they account for only 1.6 per cent of the total [Fe II] flux. On the contrary, SNR-IFOs contribute 76 per cent of the total [Fe II] flux, though represent only 12 per cent of the IFOs by number. On average, the SNR-IFOs are 191 times brighter than the YSO-IFOs. The total flux of PN and LBV-type IFOs is similar, contributing 1 per cent of the total [Fe II] flux, albeit the number of PN-IFOs is 4 times larger. In order to understand the surface brightness of each type, the size and structure of the [Fe II] sources should be taken into account. In the next section, we will compare each counterpart's characteristics in more detail.

3.2 Flux and Size Distribution

In Figure 2a, we present the flux distribution of the IFOs. The flux distributions of the individual IFO types are shown in different colors. As mentioned above, some of the IFOs share the same exciting/driving source (e.g., 8 of 25 SNR-IFOs and 6 of 17 PN-IFOs). Bearing this in mind, we see that SNR-IFOs and HII-IFOs are bright with F_{tot} as large as ~10⁻¹⁴ W m⁻², while YSO, PN, and LBV-IFOs are much fainter, with a two-orders of magnitude smaller maximum F_{tot} . YSO and PN-IFOs appear in a similar flux range but the majority of PN-IFOs are brighter than YSO-IFOs. The unknown-IFOs are generally much fainter than the other types of IFOs.

Figure 2b shows the semi-major axis distribution of IFOs. IFOs appear in a wide range of sizes, from very compact, arcsecond-scale knots to large-scale objects up to ~47 arcmin in size. The distribution with respect to types is similar to that of the flux distribution, e.g., SNR-IFOs and HII-IFOs are large and bright while YSO-IFOs and PN-IFOs are small and faint. The radius range (<10 arcsec) of unknown-IFOs is similar to that of YSO-IFOs except for a few outliers. Although there are some exceptions and scatter, the overall fluxes and sizes seem to be proportional to each other. Especially for HII-IFO, the correlation coefficient of flux and size is 0.87. When divided into CHII and HII region sub-types, it is 0.52 and 0.83, respectively. The correlation coefficient of unknown IFOs is 0.99. In contrast the coefficient for SNR-IFOs is only 0.39.

Figure 2c presents the surface brightness distribution of IFOs. Unlike the flux and size distributions, the surface brightness distribution of each type shows slightly stratified distributions. Small IFOs appear to have a higher surface brightness in general, i.e., YSO-IFOs, PN-IFOs, and unknown-IFOs have higher surface brightness than HII-IFOs and SNR-IFOs. The reason for this might be due to the low surface filling factor of [Fe II]-emitting regions in the latter sources. For example, the IFO with the lowest surface brightness is SNR-IFO 117 (Kes 78). This SNR has a large size and the [Fe II] emission is patchy, apparent only around the northern and southern caps with a marginally detectable limb. For such sources, the true surface brightness of the [FeII]-emitting regions could be much greater. In Table 1, we made a note for IFOs with small surface filling factors.

3.3 Spatial Distribution

Figure 3 shows the distribution of IFOs in Galactic longitude and latitude. One can notice the Galactic longitude distribution is clus-

tered albeit the sky coverage is more or less homogeneous. The most outstanding overdensities are seen at $l \sim 16^{\circ}$ and $l \sim 51^{\circ}$. At $l \sim 16^{\circ}$, the dominant populations are YSO- and HII-IFOs, while at $l \sim 51^{\circ}$, they are unknown- and YSO-IFOs. Including other clustered IFOs in longitude, the dominant populations responsible for these peak distributions are YSO-IFOs, followed by HII- and unknown-IFOs. A detailed description of the individual peak regions will be presented later in this section. Note that there are also voids free of IFOs at $40^{\circ} \lesssim l \lesssim 50^{\circ}$.

The distribution of the whole population of IFOs in Galactic latitude shows a Gaussian-like distribution. The distribution yields an average latitude at $b = -0^{\circ}.12$ and standard deviation $\sigma = 0^{\circ}.65$. Some concentrations of YSO-IFOs are found at $b \sim -0.^{\circ}7$, 0° and $0.^{\circ}8$. The average latitude of YSO-IFOs is -0.08 ± 0.67 . The centroids of the HII- and SNR-IFO distributions are also below the Galactic plane with an average latitude of $b = -0.09 \pm 0.63$ and b = -0.27 ± 0.58 , respectively. The average latitude of unknown-IFOs is also less than zero, i.e., $b = -0.25 \pm 0.73$. For comparison, the average latitude of PN-IFOs is $b = 0.05 \pm 0.05 \pm 0.05$. A similar trend has been observed in the UWISH2 survey; the average latitude of the jets and photodissociation regions (PDRs) was -0.018 ± 0.01 and -0.017 ± 0.01 while that of the PN group was -0.01 ± 0.01 toward the Galactic mid-plane (Froebrich et al. 2015). The distribution of IFOs (excluding PN-IFOs) being slightly shifted to the negative latitude might be related to the 'bone' structure in the first Galactic quadrant. The bone structure refers to highly elongated, dense giant molecular filaments (GMF) that are the most probable tracer of spiral arm structure (Zucker et al. 2018). It is also worth noting the scarcity of IFOs at 0.9 < b < 1.5. The number of LBV-IFOs is too small for their distribution to have any statistical meaning.

Figure 4 shows the two-dimensional distribution of IFOs in Galactic longitude and latitude along with their flux distributions. Several IFOs in the same system (e.g., jet and counter-jet of an HH object) are shown as concentric circles, as in many cases they are only a few arcsec away from each other. On average, all populations show clustered distributions with some differences from each other, though the survey coverage is homogeneous. As well as the inhomogeneous distribution of IFOs, all populations except LBV- and unknown-IFOs have more sources toward the Galactic center ($l \leq 30^\circ$). About half of unknown-IFOs are located close to those of YSOs. In addition to the similar physical properties of YSO- and unknown-IFOs might originate from activities involved in YSOs.

The region relatively devoid of IFOs in the one-dimensional longitude and latitude distribution (Fig 3) turned out to form a large-scale 2-dimensional region; IFOs hardly exist toward $l \gtrsim 30^\circ$, $b \gtrsim 0.9^\circ$ and $35^\circ \lesssim l \lesssim 50^\circ$ near the Galactic mid-plane. This might reflect spiral arm structures and the sightline toward them, where we are seeing a shorter sightline toward the Galactic bar at $l \lesssim 30^\circ$. Above this Galactic longitude we are seeing the local arm branching from Perseus Arm and Sagittarius-Scutum Arm (line of sight tangential to $l \sim 45^\circ$) at a greater distance.

We have identified some areas rich in IFOs (upper panel of Fig 4), where in particular YSO and HII-IFO are major causes of overdensity. The respective regions are as follows.

(i) $l \sim 10^{\circ}2$, $b \sim -0^{\circ}3$: This region is coincident with the HII region G10.2–0.3, one of the three HII regions in the giant star-forming region W 31. The HII region is known to be very young (~0.6 Myr). At least four O stars are residing in it, where the brightest star W 31-1 showed permitted Fe II at 1.6878 μ m and brackett lines in the NIR

spectrum. In the *H*- and *K*-band spectra ($\lambda/\Delta\lambda \approx 3000$) the [Fe II] 1.644 μ m emission line was not detected (Blum et al. 2001).

(ii) $l \sim 12^{\circ}8$, $b \sim -0^{\circ}2$: This region matches with [MDF2011b] cl1, which encircles the O4-6 (super-)giant #23 (Messineo et al. 2015). This region is immediately east of the embedded protocluster W 33 Main which is located inside the massive star-forming complex W 33. The $K_{\rm s}$ -band spectroscopy of #23 showed that the extinction of the region is $A_{\rm K} = 1.20 \pm 0.03$ mag and the luminosity class is III-I. The Oe star #22 is located between W 33 Main and #23, with line identifications of Fe II 2.0895 μ m and H₂, an extinction of $A_{\rm K} = 2.87 \pm 0.07$ mag.

(iii) $l \sim 15^{\circ}$ 1, $b \sim -0^{\circ}$ 7: This over-density is coincident with one of the most massive star-forming regions, M 17. About a hundred O- and B-type stars are responsible for the emission and the system is quite young (<1 Myrs, Hanson et al. 1997). Bautista & Pradhan (1998) reported the detection of multiple iron species, including at 1.644 μ m.

(iv) $l \sim 16^{\circ}$ 9, $b \sim +0^{\circ}$ 8: Multiple compact IFOs are connected to the apex of pillars of creation located in M 16, an active star-forming region. At the tip of the apex, there are protostars in the pillar's EGGs ('Evaporating Gaseous Globules'), which are not yet hot enough to emit X-rays. Therefore, the IFOs in M 16 might be tracing some of the youngest protostars.

(v) $l \sim 25^{\circ}4$, $b \sim -0^{\circ}2$: The region corresponds to W 42, an obscured giant HII region. The closest nearby source is [BCD2000] W 42 1, an O5.5 star (Blum et al. 2000). There are several point-like sources that might be true [Fe II] sources or mere variables.

(vi) $l \sim 30.^{\circ}7$, $b \sim -0.^{\circ}0$: The IFO is close to one of the closest starburst regions, W 43. This giant H_{II} region has a central open cluster with massive stars.

(vii) $l \sim 49^\circ$ 1, $b \sim -0^\circ$ 6: Multiple HII- and YSO-IFOs are located in the vicinity of W 51, which is one of the most massive giant molecular clouds that is optically obscured. All large-scale representative structures, namely W 51 Main, IRS 1, and IRS 2, are bright in the [Fe II] 1.644 μ m line. Each structure shows a distinct star-forming phase as follows: W 51 Main - several UCHIIs are located. IRS 1 evolved HII region with a size of ~1 pc. IRS 2 - went through recent star formation, and an ~O3 star and a massive YSO were found (Barbosa et al. 2008). An LBV-IFO is also coincident with the region, which is a high-mass evolved star (P Cygni supergiant) with evidence for chemical enrichment (Clark et al. 2009).

(viii) $l \sim 53^{\circ}2$, $b \sim +0^{\circ}0$: Multiple YSO-IFOs coincide with an IRDC G53.2, which was formerly catalogued as three IRDCs in the *Midcourse Space Experiment (MSX)* dark cloud (MSXDC) catalog (Simon et al. 2006). The three IRDCs, viz. MSXDC G053.11+00.05, MSXDC G053.25+00.04, and MSXDC G053.31+00.00 harbor hundreds of YSO and YSO candidates, some of them in the vicinity of IFOs.

(ix) $l \sim 59^{\circ}4$, $b \sim -0^{\circ}2$: The IFOs are located in the central part of SH 2-87, a complex massive star-forming region. The three submillimeter clumps, SMM 1, SMM 2, and SMM 3 constitute this HII nebula. These three clumps are at separate evolutionary stages (Xue & Wu 2008), and two HII-IFOs were found in the vicinity of the hottest and most massive star-forming clump, SMM 1.

4 DISCUSSION

4.1 Outflows/Jets from Young Stellar Objects

Outflows/jets of YSO are composed of ejected and circumstellar swept-up material, and are recognized as an important signpost of recent star-forming activity. This phenomenon plays a key role in conventional disc accretion-outflow theories, the outflow being responsible for the removal of angular momentum and kinetic energy of accreting material that enables accreting material to overcome the centrifugal force and collapses to form a star (see theories of disk-wind; Pudritz & Norman 1983, X-wind; Shu et al. 1994, and observational studies; Ellerbroek et al. 2013 for reference).

Thanks to the development of IR instruments, previously undiscovered, highly obscured outflows have begun to be found in the near-infrared. The optical HH objects and their IR counterparts basically refer to the same phenomena, and only the conditions of jet and circumstellar matter differ. So far, molecular emission (e.g., MHO) has drawn attention in the NIR, alongside atomic/ionic lines in the optical, yet less attention has been brought to the [Fe II] lines in the NIR. The [Fe II] 1.644 μ m line, the brightest iron line in the H-K band, is reported to unveil a shocked region that is denser and/or more ionized than regions where optical lines are generated (Nisini et al. 2002). In this aspect, previous studies using frequently used molecular tracers, namely SiO, CO, and HCO⁺ in the sub-mm to mm, only revealed secondary outflows, tracing masses of low-density, distant (up to a few pc) outflows. Whereas the [Fe II] 1.644 μ m line from the jet is found to extend a few AUs to parsec-scales in the form of a dense irradiated jet (Reiter et al. 2015).

Most previous [Fe II] outflow studies are confined to certain types of objects or regions: specific star-forming regions (Orion; Takami et al. 2002, Carina; Reiter et al. 2016, Shinn et al. 2013) or a certain mass range of YSOs (Caratti o Garatti et al. 2006; Caratti o Garatti et al. 2015). Recently, outflow studies toward external galaxies, namely the Large Magellanic Cloud (LMC) and the Small Magellanic Cloud (SMC), became feasible (Reiter et al. 2019). These studies showed that the [Fe II] emission tends to be observed at the tip of the bipolar outflow and is rather collimated, compared to H₂ and H α which predominantly show the morphology of a 'wake' enclosing the [Fe II] emission (Reiter et al. 2015).

We have detected 100 YSO-IFOs (Table 3). Our result provides a large and comprehensive sample for the study of [Fe II] emission associated with YSOs. Figure 5 shows the example of identified YSO-IFOs, displaying UKIDSS *KHJ*-band RGB images to show how the YSO-IFOs reveal unique structures in comparison to hot dust continuum structures. YSO-IFOs show diverse morphologies, diverse compared to traditionally observed/expected [Fe II] features that are located at the tip of bipolar outflows and/or are highly collimated toward the driving sources (Caratti o Garatti et al. 2006; Reiter et al. 2016).

We classified YSO-IFOs into four morphological categories; bipolar, cometary, knot-like, and amorphous. A representative case of each category would be IFO 13-14, 125, 122, and 4 in Figure B.1, respectively. Bipolar YSO-IFOs are a textbook case of star formation, consistent with the accretion-jet theory with the aid of a magnetic field (Konigl 1982; Shang et al. 2020; Frank 1999). They typically show two lobes located on opposite sides of a central source, but some show two wakes, tips, and collimated bow-shock shapes, distributed laterally from the apparent YSO jet axis. The prototypical bipolar YSO-IFOs are IFO 13 and 14. The [Fe II] 1.644 μ m emission with bipolar morphology usually represents either the 'cap' of bow shock where an outflow collides with the ambient medium or dense, collimated jets. Cometary YSO-IFOs resemble a comet with a bright head around the driving source and a narrow faint tail-like structure. The prototypical cometary YSO-IFOs are IFO 125 and 131, both having well-defined conical structures. They are located at quite different distances, i.e., 4.7 and 1.1 kpc, and the extent of the associated conical structures has very different linear scales, i.e., ~45000 AU (10 arcsec) and 5000 AU (5 arcsec). For the wide-angle tails of cometary morphology, it is possible that either (1) the jet is bending and/or precessing (Paron et al. 2016) (2) a cavity structure is revealed (Hsieh et al. 2017) (3) a multiple systems presence is implied (Fuente et al. 1998). Knot-like YSO-IFOs appear as knots, sometimes located symmetrically from a driving source along a certain axis. The representative knot-like YSO-IFOs, 122 and 123, are showing well-isolated compact features. These knot-like features might imply that the ejection of accreted material in the system is accompanied by sporadic bursts of accretion (Caratti o Garatti et al. 2015). Amorphous YSO-IFOs represent the remaining YSO-IFOs that are diffuse and do not have a definitive structure. The nature of the amorphous YSO is uncertain. The number of YSO-IFOs classified as bipolar, cometary, knot-like, and amorphous is 16, 18, 19, and 47, respectively.

The morphologies of YSO-IFOs are closely related to the nature of YSOs and their mass-loss histories (Caratti o Garatti et al. 2015; Paron et al. 2016). For example, the collimated and continuous jet morphology indicates a continuous ejection of accreting material from the accretion disc system (Reiter et al. 2016; Reiter et al. 2017). The overabundance of amorphous morphologies might suggest highly varying environments or multiple systems are affecting the outflow structure. But the morphology of YSO-IFOs might depend on environments as well as foreground extinction, so detailed studies are needed of the individual objects to confirm their nature. Thirteen YSO-IFOs are associated with HH objects (Table 4). Figure 6 shows a comparison of their [Fe II] and H α images.

YSO-IFOs constitute half the number of our cataloged sources, making YSO the most common IFO in the inner Galaxy. The number density of YSO-IFOs is 0.55 deg^{-2} . For comparison, the H₂ number density probed by UWISH2, which covered an almost identical area with a comparable surface brightness limit, is 2.15 deg^{-2} (Froebrich et al. 2015). The flux density of YSO-IFOs ranges (2–820) $\times 10^{-18}$ W m⁻² with a mean of 4.3×10^{-17} W m⁻². This range can be compared with the results of other surveys. Caratti o Garatti et al. (2006) targeted H₂-emitting low-intermediate luminosity Class 0/I YSOs and reported that among 23, 74 per cent were also detected in [Fe II]. For the newly observed 9 [Fe II] line jets in the reference, the flux range is $(2.8-27.0) \times 10^{-18}$ W m⁻². Caratti o Garatti et al. (2015) observed 18 intermediate- to massive-YSOs having H₂ and EGO counterparts, and the flux range is $(2.5-61.9) \times 10^{-18}$ W m⁻². Note that these fluxes are obtained from spectroscopic studies using a slit of width 1 arcsec. The majority of YSO-IFOs have flux densities comparable to those of previous studies. But a few sources are exceptionally bright. The number of YSO-IFOs brighter than outflows observed in Caratti o Garatti et al. (2015) is 10 per cent of the YSO-IFOs. Since these bright YSO-IFOs do not share certain morphologies and 40 per cent of them have RMS counterparts, they might be preferentially massive YSO outflows, which have simply not yet been identified due to the limited sky coverage of past [Fe II] observations. One possible speculation is that [Fe II] brightness does not strictly scale with driving source brightness or other outflow tracers, based on the target of previous studies, which tend to be bright IRAS sources accompanying outflows discovered in other tracers. This illustrates the importance of an unbiased study to correct our understanding of the characteristics of [Fe II] emitters.

The YSO-IFOs and jet-group MHOs of the UWISH2 survey can be compared one-to-one since the UWIFE survey area was fully covered by UWISH2. The spatial distribution of YSO-IFOs in Figure 4 shows a highly clustered distribution, accompanied by the high-latitude sources in $l \sim 15-30^{\circ}$ and the absence of YSO-IFOs in the Galactic mid-plane at $l \sim 40-50^{\circ}$. This characteristic distribution is also shared in jet-group MHOs (see Figure 8 in Froebrich et al.

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2015). As seen in Figure 5, about 85 per cent of YSO-IFOs accompany jet/PDR-group MHOs in the vicinity. For example, in the M 16 (Eagle nebula), 6 YSO-IFOs were identified, and a few hundreds of jet/PDR-group MHOs are also present. A detailed comparison of YSO-IFOs with jet/PDR-group MHOs discovered in the subsequent UWISH2 studies will be helpful for the comparison of different shock tracers (Ioannidis & Froebrich 2012a; Ioannidis & Froebrich 2012b; Froebrich & Makin 2016; Makin & Froebrich 2018; Samal et al. 2018).

We can compare our results with the results of the RMS survey where NIR spectra of YSO candidates have been obtained. In the common survey area $(10^{\circ} \leq l \leq 62^{\circ}, |b| \leq 1.5^{\circ})$, there are 182 RMS sources, and among the 72 sources from which spectra have been obtained, 58 have [Fe II] line emission, though some of the detections could be confused with the Br 12 line. For comparison, only 17 of 182 RMS objects have been identified as YSO-IFO in our study (for some RMS sources, 2-3 IFOs correspond to one RMS source.) Among these 17 sources, the NIR spectra have been obtained for 8 sources, and [Fe II] lines were detected in 6 sources, i.e., [Fe II] lines were reported as non-detection for two sources in the RMS survey. We note that the non-detection for the two is based on a comparison of Br 11 and Br $12/[Fe_{II}]$ line strengths (Br 11×0.788 > Br $12/[Fe_{II}]$) and it might be possible that a weak [Fe_{II}] line is in fact present but missed by low spectral resolution, as the authors noted (Cooper et al. 2013). To assess this possibility, we checked the slit configuration (central position, position angle) in Cooper et al. (2013) and compared it with YSO-IFO morphology. For both IFO 72 and 141, the RMS slit intersects the driving source but does not include the bright part of extended YSO-IFO structures. Indeed, the authors tried to include extended structures inside the slit in imaging mode prior to spectroscopy mode, yet even narrow-band [Fe II] images of UWIFE without continuum subtraction turned out to severely hinder extended emission. Therefore, most YSO-IFOs apparently do not have RMS source counterparts, which is claimed to be a 90 per cent complete list of massive protostellar populations (Lumsden et al. 2013). This seems to suggest that most of YSO-IFOs are associated with low-mass star formation. It is also worthwhile to note that the majority of YSO-IFOs (87 per cent) are not associated with HHs, which suggests that the [Fe II] emission is tracing optically hidden star-forming regions.

4.2 Compact HII Regions

Compact and ultra-compact HII regions are the earlier stages of 'classical' HII regions. An UCHII region is a photoionized region with a diameter $\lesssim 0.1$ pc and an electron density $n_e\gtrsim 10^4~cm^{-3}$, embedded in a molecular cloud (Wood & Churchwell 1989). In this evolutionary stage, mass accretion of the central star is thought to be insignificant (Churchwell 2002; Zinnecker & Yorke 2007). A CHII region is a HII region in the intermediate phase between UCHII and classical HII regions, having a radius $\lesssim 0.1$ pc and $n_e\gtrsim 10^3~cm^{-3}$. The lifetime of UCHII and CHII is $\sim 2-4\times 10^5$ yr (Davies et al. 2011, Mottram et al. 2011).

In UCHII and CHII regions, [Fe II] emission can be enhanced by the interaction of stellar wind with the ambient medium. Bloomer et al. (1998) detected enhanced shell-like [Fe II] emission along the periphery of the CHII region NGC 7538 IRS 2. The observed [Fe II] 1.644 μ m / Br γ ratio was 0.15, which is an order of magnitude greater than that of HII regions, and it implies that the [Fe II] line emission emanates from shocked stellar wind material. Shinn et al. (2014) searched for [Fe II] 1.644 μ m emission associated with UCHII regions employing the CORNISH UCHII catalog and the UWIFE survey data. Among the 237 UCHII regions in the survey area, five and one candidate were found to have associated [Fe II] emission features, which were suggested to be shock-excited by outflows from central YSOs. Kim et al. (2017) also reported the detection of [Fe II] emission from UCHII Monoceros R2. Hereafter, we refer to IFOs associated with CHII/UCHII regions or with HII regions in even earlier evolutionary stages as CHII-IFOs.

We have detected 22 IFOs associated with 16 UCHII/CHIIs (Table 5). Six IFOs (IFO 24, 25, 26, 97, 107, and 156) had been previously reported by Shinn et al. (2014). We have discovered IFOs associated with an UCHII precursor (IFO 137) and an ultra-compact embedded cluster (IFO 138), which are thought to be earlier progenitor or less massive populations (Molinari et al. 1998; Alexander & Kobulnicky 2012). Among the 16 UCHII/CHII regions with [Fe II] emission features, 10 are catalogued in CORNISH, which corresponds to 4 per cent of the 237 UCHII regions in the CORNISH catalog in the survey area. The detectability might be partly due to the large extinction in UCHII/CHII regions, which is typically $A_V \sim 30-50$ or $A_K \sim 3-5$ (Hanson et al. 2002). Indeed, the A_V of three UCHIIs with associated IFOs had been found to have relatively low extinction ($A_V \sim 9-20$, Shinn et al. 2014).

Figure 7 shows the 22 CHII-IFOs. CHII-IFOs have diverse morphologies, e.g., jet-like, shell-like, and amorphous morphologies. A representative IFO with jet morphology is IFO 97, which appears as a collimated beam from the center to the boundary of the HII region. The jet appears to extend beyond the radio continuum boundary (see Figure 7), which might reflect a possible correlation with the boundary of the ionization front (Goddi et al. 2020). The representative IFO with a shell-like morphology is IFO 132. An exemplary CHII-IFO of amorphous morphology would be the IFO 138, having a diffuse structure either outside or inside of the HII region in the radio. The properties of a compact HII region have been rarely investigated in [Fe II] emission. Shinn et al. (2014) proposed that some IFOs identified in the vicinity of ultra-compact HII (UCHII) regions (IFOs 24-26, 97, 107, 156) are the 'footprint' outflow features of UCHIIs, i.e., the features produced by outflowing material ejected during an earlier, active accretion phase of massive young stellar objects, based on the morphological relation between the [Fe II] and 5 GHz radio features, the outflow mass-loss rate, the travel time of the [Fe II] features, and the existence of several YSO candidates near the UCHIIs. The newly discovered CHII-IFOs in this study might serve as a chance to investigate the origin of the [Fe II] emission in the vicinity of CHIIs.

4.3 HII Regions

HII regions are not expected to be bright in the [Fe II] lines, since in their photoionized regions, Fe atoms are predominantly in higher ionization states, and Fe atoms are thought to be mostly locked in dust grains (Koo et al. 2016). According to theoretical models of photoionized regions, the [Fe II] emission from an HII region is mainly emitted in the high-density partially ionized zones near ionization fronts, predominantly excited by electron collisions (Oliva et al. 1989; Bautista & Pradhan 1998). In the Orion nebula, for example, [Fe II] images exhibit filamentary structures and diffuse emission that might be associated with ionization fronts, together with some knotty features (Takami et al. 2002). Expanding HII regions can drive shocks, but the shock velocity is low ($\sim 10 \text{ km s}^{-1}$) so [Fe II] line emission is not expected to be enhanced (e.g., Mouri et al. 2000). The [Fe II] 1.644 μ m/Pa α ratio of the Orion is 0.013, which is more than two orders of magnitude smaller than those of SNRs (Oliva et al. 1989; Mouri et al. 2000). So Galactic HII regions have not been a popular target

of deep and high-resolution [Fe II] imaging (Kraus et al. 2006; Bally et al. 2022). The depletion of Fe atoms in the HII region, however, is uncertain. In the Orion nebula, it has been estimated that 90 per cent of Fe is locked onto dust grains (Baldwin et al. 1991; Baldwin et al. 1996; Osterbrock et al. 1992; Rodríguez 2002). But there are studies which showed that, in many HII and star-forming regions, Fe is not depleted as heavily as in the Orion nebula (Osterbrock et al. 1992; Peimbert 1993; Rodríguez 2002; Okada et al. 2008; Peimbert & Peimbert 2010). It has been suggested that some populations of dust grains might be easily destroyed by UV radiation from OB stars and Fe atoms are released into the gas phase (Okada et al. 2008; Peißker et al. 2020). For external galaxies, Alonso-Herrero et al. (2003) did an imaging study of the starburst galaxies M 82 and NGC 253 in [Fe II] 1.644 μ m and Pa α (1.87 μ m) lines, and, by comparing their intensity ratios, concluded that 6-8 per cent of [Fe II] line fluxes are due to HII regions. Mouri et al. (2000), Riffel et al. (2016), Hennig et al. (2018), and Fazeli et al. (2019) suggested that some of the [Fe II] emission from external galaxies could be due to HII regions based on their low [Fe II] $1.257 \mu m$ / Pa β ratios.

We have identified 11 IFOs associated with 4 HII regions (Table 6). All HII-IFOs are located in the well-known star-forming complexes W 31, M 17, and W 51. Figure 8 shows the 11 HII-IFOs. We can see that some IFOs appear as thin filaments elongated along the radio structure (e.g., IFO 55, IFO 62) or as diffuse amorphous emission structures within the radio structure (e.g., IFO 11 and 12), so the association of IFOs with HII regions is very likely. The filamentary structures might correspond to ionization fronts and/or boundaries of PDRs as in the Orion nebula. On the other hand, some IFOs are faint and diffuse, and they extend beyond the radio boundary of the HII regions, e.g., IFO 159 and 161, so their association is uncertain and needs to be confirmed.

4.4 Planetary Nebulae

Planetary nebulae represent a short-lived phase near the endpoint of low- to intermediate-mass star (1-8 M_{\odot}) evolution which is preceded by the AGB, post-AGB, and pre-PN phases. The circumstellar envelope of the AGB carbon star is considered highly Fe-depleted (Mauron & Huggins 2010), though Fe becomes abundant with time (Fe abundance is negatively correlated with the C/O ratio, Delgado-Inglada & Rodríguez 2014). In turn, PNe are not expected to be strong [Fe II] emitters, also having a Fe-deficit nature with <10 per cent existing in gas and the remaining probably enshrouded in dust grains (Delgado-Inglada & Rodríguez 2014). Meanwhile, in the context of environmental factors, PNe could be a [Fe II] emitter since it has a partially ionized zone where Fe⁺ is apt to exist, and at a certain point of its evolution, a low-velocity shock is expected to occur. In short, suitable ionization conditions and energy to excite Fe (Greenhouse et al. 1991) can be established in PNe, and its iron-depleted nature is a key factor to determine the existence of $[\mbox{Fe}\,{\mbox{$\sc n$}}]$ emission.

Besides the theoretical expectation, previous studies reported the detection of [Fe II] emission towards stellar objects in a variety of evolutionary stages: post-AGB (IRAS 16594-4656; Van de Steene & van Hoof 2003), pre-PN (M 1-92; Davis et al. 2005), and PN (Hubble 12; Welch et al. 1999, M 2-9; Smith et al. 2005, NGC 2440; Hora et al. 1999). Some authors suggested a circumstellar origin (e.g., Clark et al. 2014, Smith et al. 2005), especially Baan et al. (2021) reported the detection of [Fe II] emission revealing the interaction of an accretion inflow, which is composed of material ejected in earlier post-AGB and pre-PN circumstellar material, and stellar outflow.

Table 7 shows PN-IFOs. They are IFOs spatially coincident with PNe, PN candidates, and sources in earlier evolutionary stages such as

post-AGBs. Seventeen PN-IFOs are associated with 14 PNe; 5 PNe, 8 PN candidates, and one post-AGB candidate. For comparison, in a previous study, Lee et al. (2014) reported the detection of [Fe II] emission in six PNe among 29 known PNe. In the survey area, there are 296 HASH 'true' (131), likely (40), and possible (125) PNe, so that the detection rate is 4.7 per cent. If we limit the sample to the 'true' PN, the detection rate slightly drops to 3.8 per cent (i.e., 5 out of 131). This very low detection rate of PNe in [Fe II] emission (4.7 and 3.8 per cent) contrasts with the results in H₂, where detection rates are 30 and 21 per cent, respectively (for $10^{\circ} < l < 66^{\circ}$, |b| <1°5, Gledhill et al. 2018). It is interesting that even with an order of magnitude larger sample of PNe in this study, our result is somewhat consistent with the former [Fe II] and H₂ detection rates of 7 and 39 per cent derived from 41 PNe (Hora et al. 1999). The slightly higher detection rate of Hora et al. (1999) could be because their samples are either moderately sized or optically bright.

The number density of PN-IFOs is 0.07 deg⁻² within 180 deg² whereas it is 1.25 deg⁻² within 209 deg² in UWISH2 (Froebrich et al. 2015). However, unlike the previous argument (Kastner et al. 1996), not all [Fe II]-emitting PNe are seen in H₂ emission; we found 3 of our 14 PNe in Table 7 were absent from the list of PNe with H₂ emission. Also, the median flux of PN-IFOs is greater than that of the H₂-emitting PNe, i.e., 6.46×10^{-17} W m⁻² versus 4.53×10^{-17} W m⁻². Therefore, our result shows that the H₂-emitting PNe are not necessarily brighter than the non-H₂-emitting PNe in [Fe II] emission.

Figure 9 presents [Fe II]-H images of PN-IFOs. Note that there are three bipolar PNe, each of which possesses two associated IFOs (IFO 5 and IFO 6, IFO 8 and IFO 9, and IFO 129 and IFO 130). We classified the morphologies of PN-IFOs using the basic 'ERBIAS' classifier following Parker et al. (2006), where 'E'=elliptical, 'R'=round, 'B'=bipolar, 'I'=irregular, 'A'=asymmetric, and 'S'=quasi-stellar. Their sub-classifiers 'amprs' are also adopted to describe detailed morphology; the main object has a one-sided enhancement/asymmetry 'a', has multiple shells or external structure 'm', exhibits point symmetry 'p', has a well-defined ring structure or annulus 'r', or resolved internal structure 's'. An IFO can have several 'amprs' sub-classifications. The results are summarized in Table 7, where their morphologies in $H\alpha$ and H_2 are also listed (Parker et al. 2016; Gledhill et al. 2018). The H α morphologies are from the HASH survey, while the H₂ morphologies are from the UWISH2 survey. For the PN-IFOs without a counterpart in the UWISH2 PN catalog (IFO 50, 95, 188), we inspected the UWISH2 data and classified their morphology in the same format (see Table 7). Some PNe have different morphologies in the [Fe II], H₂, and H α emission, which implies a complex surrounding environment and/or complex mass-loss history.

The physical sizes of PN-IFOs have been derived for 10 PNe that have previously estimated distances (Table 7). The sizes of 4 IFOs associated with 'true' PNe range from 0.13 pc to 0.92 pc, and three of the PN-IFOs are larger than 0.9 pc. This contrasts with the majority of PNe in H α being \leq 0.2 pc (González-Santamaría et al. 2020). This seems to suggest that the [Fe II] emission preferentially traces large, bright PNe.

For example, in PNG 050.4+00.7, the size of the associated IFO (IFO 157) substantially exceeds the previously known size of the counterparts (2 arcmin and 19 arcsec, respectively).

The IFO has a partial 'S' shape elongated along the east-west direction, with IRAS 19194+1548 superposed at the western part. The structure becomes gradually fainter toward the west, therefore the angular size of the partial 'S' shape should be considered as lower limit. The implied physical scale of 3.1 pc largely surpasses the generally accepted size of PNe (one of the oldest and largest PNe,

the Helix nebula has an outermost size of 1.76 pc). The driving source is suspected to be in a symbiotic star system (Akras et al. 2019) and the updated size is compatible with the sizes of large shells/nebulae around symbiotic stars (McCollum et al. 2008).

4.5 Nebulae of Luminous Blue Variables

Infrared [Fe II] 1.644 μ m emission around prominent nebulosity of LBVs is thought to be ubiquitous. Smith (2002) (henceforth, S02) searched for [Fe II] 1.644 μ m emission in nine well-known LBVs and found the emission in 7 of them, resulting in a detection rate of 77 per cent. This high detection rate surpasses that of SNRs (i.e., 24 per cent, Lee et al. 2019), the population that is thought to provide the most adequate environment for the existence of [Fe II] 1.644 μ m emission. S02 could not pinpoint the essential condition needed for strong [Fe II] emission to exist. Shock heating and radiative heating as possible excitation mechanism of [Fe II] emission were suggested by the author.

Shock-excited [Fe II] emission can arise when the LBV's environment meets requirements such as (i) a large difference in the outflow speed between the stellar wind and pre-existing LBV nebula and (ii) a difference of velocities between the stellar wind and ejected shell created during S Doradus outbursts or giant eruption phases. This velocity difference of 50–150 km s⁻¹ (S02) is ascribed to weaker gravity in an active phase. When LBV evolves toward a cooler temperature (to a local temperature lower than 30,000 K), Hydrogen atoms and opacity-enhancing ions start to emerge on the surface, which is known as the 'modified' Eddington limit (Humphreys & Davidson 1994). The elevated opacity makes the outward radiation pressure stronger and overpowers the inward gravity force. The resultant lower effective gravity helps LBVs easily induce the aforementioned mass loss. In these S Doradus outbursts and giant eruption phases, the weaker gravity results in an ejected shell having a lower expansion velocity than normal stellar winds. The following post-eruption wind has a velocity higher than that of the aforementioned high mass-loss phase and eventually overtakes the ejected shell. Meanwhile, photoionized [Fe II] emission was reported from two hot (30,000 K) LBVs (AG Car and R 127, S02) which was attributed to their stronger UV flux.

We detected [Fe II] emission features associated with 3 LBVs (Table 8). So the [Fe II] detection rate of LBV nebula in our study is 14 per cent. If we include the 9 LBV samples of S02, the detection rate would be 29 per cent, i.e., 9 out of 30 LBV nebulae (HD 168625 duplicated in both studies). This new detection rate with a three-fold sample is lower than the previous study, making the general physical conditions of LBVs not particularly suitable for the [Fe II] 1.644 μ m line to arise but comparable to those of SNRs. The discrepancy in detection rates between this study and S02 might be due to the biased sample S02 used, which includes confirmed LBVs and candidate LBVs showing nebulosity in the Galaxy and the two most famous LBVs in LMC.

The [Fe II]-H images of identified LBV-IFOs are shown in Figure 10. Brief information about them is listed in Table 8. In the [Fe II] emission, all identified LBV-IFOs share an elliptical/circular morphology. This is similar to their morphologies at 8 μ m but the extent appears smaller. We note that for G26.47+0.02 (IFO 102, 103) the south-eastern diffuse structure was noticed in the [Fe II]-H image. But the possibility of it being an artifact prevented us from assigning it as an IFO. The morphological coincidence of this South-East structure, IFO 102 and 103 with respect to the prominent part of the 8 μ m nebula (Paron et al. 2012) implies the possibility of more extended, diffuse [Fe II] emission than seemingly identified. There are some new features revealed by [Fe II] emission: (1) IFO 65 - HD 168625 is

located at the center of optical/IR elliptical structures that are broken toward the north-east. In the [Fe II] emission, we see a complete circular structure, the center of which is offset toward the northeast. (2) IFO 162 - [KW97] 37-17 shows multiple shells in [Fe II] emission, forming together a much brighter elliptical structure than those in 8 μ m or optical. This possibly indicates that the LBV had several active erupting phases that manufactured bright [Fe II]-emitting shells one by one.

We found that all [Fe II]-detected LBVs in the UWIFE survey also accompany nebulosities at 8 μ m, but not vice versa. For example, we could not detect [Fe II] emission in three LBVs with 8 μ m nebulosity (HD 168607, AFGL 2298, and GAL 024.73+00.69). Thus, the question of whether the LBV nebula, on account of the preceding giant eruption, is a prerequisite for the [Fe II] emission remains unanswered (S02). More comprehensive LBV samples and constrained physical properties of LBVs are needed to understand the possible relationship between the existence of the [Fe II] 1.644 μ m line in LBV nebula and their past eruption histories.

4.6 Supernova Remnants

SNRs are the brightest objects in [Fe II] emission. In SNRs this line is mostly emitted from cooling gas behind radiative shocks. [Fe II] lines are strong in shocked gas because Fe abundance could be enhanced by shocks owing to grain destruction (Dinerstein 1995; Nisini 2008; Koo et al. 2016 and references therein). Before the UWIFE survey, a dozen Galactic and LMC SNRs had been observed in the NIR [Fe II] lines. The SNRs that are bright in [Fe II] emission lines may be divided into two groups: (1) middle-aged SNRs interacting with dense molecular (or atomic) clouds such as W 44 (Reach et al. 2005), 3C 391 (Reach et al. 2002), and (2) young SNRs interacting with the dense circumstellar medium (CSM) such as Cas A (Koo et al. 2018), G11.2-0.3 (Moon et al. 2009), RCW 103 (Burton & Spyromilio 1993), and W49B (Lee et al. 2019). Then Lee et al. (2019) (hereafter, L19) searched for [Fe II] emission at the positions of the SNRs in the catalog of Green (2014) using the UWIFE survey data and detected [Fe II] emission features toward 19 SNRs, more than half of which were new detections. In external galaxies, [Fe II] emission is used as a tracer of SNRs (Blair et al. 2014; Bruursema et al. 2014; Long et al. 2020), although strong [Fe II] lines may originate from sources ionized by X-rays, e.g., in active galactic nuclei (Mouri et al. 2000; Morel et al. 2002).

We detected 25 IFOs associated with SNRs. All these SNR-IFOs belong to the 19 SNRs in Lee et al. (2019) except one (Table 9). It is worthwhile to point out that Lee et al. (2019) searched [Fe II] emission at 79 SNRs of the Green's catalog that are fully covered by the UWIFE survey. Four SNRs partially observed in the survey (i.e., G7.0-0.1, G13.3-1.3, G28.8+1.5, G38.7-1.3) were not investigated, and our unbiased search resulted in the identification of a small [Fe II]-emitting patch inside the region of G28.8+1.5. Meanwhile, the Green's catalog of Galactic SNRs has been updated (Green 2019), adding a new SNR G53.4+0.0 (partially covered in the UWIFE) and rejecting four (G20.4+0.1, G21.5-0.1, G23.6+0.3, G59.8+1.2 that were reclassified as HII regions) in the survey area. None of the new or rejected SNRs showed [Fe II] emission features. So the new detection rate for fully covered SNRs is 25 per cent (19/75).

We note that Lee et al. (2019) compensated the [Fe II] line flux for the flux subtracted from the *H*-band by multiplying with 1.15, whereas the fluxes in Table 9 are observed fluxes. As presented in Lee et al. (2019), IFO 147 that matches W49B is the brightest SNR-IFO. The detailed results for the 19 SNRs can be found in Lee et al. (2019).

5 SUMMARY

We have presented the first comprehensive catalog of Galactic IFOs discovered in the UKIRT Widefield Infrared Survey for [Fe II] (UWIFE). It is the first Galactic catalog of extended [Fe II] line emission sources using an unbiased, large-scale survey. We have discovered many previously unreported [Fe II] 1.644 μ m line sources. Therefore, this catalog provides an opportunity to broaden the horizons of the study of the shocked regions of our Galaxy, especially with the synergy of the UWISH2 survey.

We have searched for extended IFOs in the inner Galactic plane $(7^{\circ} < l < 62^{\circ}; |b| \le 1.5)$. In order for the search to be efficient, we removed point-like continuum sources from the [Fe II] 1.644 μ m images using H-band images taken as part of the UKIDSS GPS survey. We identified most of the IFOs by visual inspection and added several faint IFOs with an automatic source identification which uses the same source detection algorithm as in UWISH2 (Freebrich et al. 2015). In total, 204 IFOs were identified. We measured the sizes and fluxes of these 204 IFOs and presented their properties. We have searched for the counterparts of the IFOs via positional cross-matching with previously known sources and found that the majority of IFOs are associated with supernova remnants, young stellar objects, HII regions, planetary nebulae, and luminous blue variables. We group IFOs by their counterpart types and discuss their statistical and morphological properties. The main results are summarized as follows.

(1) In the 180 deg² Galactic plane area of the 1st Galactic quadrant covered by the UWIFE survey (7° < l < 62°; $|b| \leq 1^{\circ}5$), we identified 204 IFOs. The identified IFOs are classified according to their counterparts: IFOs associated with young stellar objects (YSO-IFOs), HII regions (HII-IFOs), compact HII regions (CHII-IFOs), planetary nebulae (PN-IFOs), luminous blue variables (LBV-IFOs), and supernova remnants (SNR-IFOs). There are 100 YSO-IFOs, 11 HII-IFOs, 22 CHII-IFOs, 17 PN-IFOs, 4 LBV-IFOs, and 25 SNR-IFOs. We could not identify counterparts for 25 IFOs, and they are classified as 'unknown-IFOs'. The majority of IFOs are new discoveries that have never been revealed in previous [Fe II] line studies.

(2) The SNR-IFOs and HII-IFOs are the brightest IFOs, and they dominate the [Fe II] 1.644 μ m line flux in the Galactic plane. They contribute 96 per cent of the total [Fe II] 1.644 μ m line flux of the IFOs (2.6 × 10⁻¹³ W m⁻²); 76 per cent by SNR-IFOs and 20 per cent by HII/CHII-IFOs. The YSO-IFOs, PN-IFOs, and LBV-IFOs are generally orders of magnitude fainter, while the unknown-IFOs are the faintest.

(3) The average number density of IFOs is ~1.1 deg⁻². The number density is highly variable spatially, especially for the IFOs associated with objects in the early-evolutionary phase, e.g., IFOs associated with HII regions and YSOs. In Galactic longitude, there are prominent peaks at $l \sim 16^{\circ}$ and 51°, while there is a 'void' at $l \sim 40^{\circ}-50^{\circ}$ where the number of IFOs is very small. The spatial distribution in Galactic latitude is centered at $b=-0^{\circ}$ 12 with a standard deviation of 0.65.

(4) The results on the individual types of IFOs are summarized below.

(i) YSO-IFOs

We detected 100 YSO-IFOs, which constitutes half of the IFOs in our catalog. Only seventeen of those are associated with the RMS sources, which represent massive YSOs. The YSO-IFOs might be preferentially tracing low-mass YSOs. On the other hand, the majority (87 per cent) of YSO-IFOs are not associated with HH objects, suggesting that the YSO-IFOs are revealing previously hidden, optically obscured outflows in star-forming regions. YSO-IFOs have diverse morphologies, and we have classified them into four categories; bipolar, cometary, knot-like, and amorphous.

(ii) HII-IFOs and CHII-IFOs

We have identified 11 IFOs associated with 4 HII regions (Table 6). Almost all HII-IFOs are located in the well-known starforming complexes, W 31, M 17, and W 51. Some HII-IFOs appear as either thin filaments or diffuse amorphous emission structures within the radio structure, so their association with the HII regions is very likely. But some are faint and diffuse and extend beyond the radio boundary of the HII regions, so their association is uncertain and needs to be confirmed. We also detected 22 IFOs associated with 16 CHIIs, including 6 previously reported (Table 5). Among the 16 CHII regions, ten are catalogued in CORNISH, which corresponds to 4 per cent of the 237 CHII regions in the CORNISH catalog in the survey area. CHII-IFOs have diverse morphologies: jet-like, shell-like, and amorphous.

(iii) PN-IFOs

We detected 17 PN-IFOs. They are associated with 14 PNe (i.e., 5 PNe, 8 PN candidates, and one post-AGB candidate; Table 7), which correspond to about 4.7 per cent of the PNe in the survey area. We have classified the morphologies of PN-IFOs following Parker et al. (2006) and compared them with those in H α and H₂. Some PNe have [Fe II] morphologies different from the H α and H₂ morphologies, which implies that the [Fe II] line reveals new substructures, possibly probing additional mass-loss histories. The physical sizes of some PN-IFOs are larger than 0.9 pc.

(iv) LBV-IFOs

We detected 4 LBV-IFOs, They are associated with 3 LBVs out of 22 LBVs and their candidates in the survey area (Table 8), so the detection rate of [Fe II] emission associated with LBVs in this study is 14 per cent. All LBV-IFOs share an elliptical or circular morphology. Some show multiple shells. We found that all [Fe II] detected LBVs in the UWIFE survey also accompany nebulosity at 8 μ m, but not vice versa.

(v) SNR-IFOs

We detected 25 SNR-IFOs. They are associated with 20 SNRs, which corresponds to 25 per cent of the 75 known SNRs in the survey area. The SNR-IFOs occupy 76 per cent of the total [Fe II] flux of IFOs, and the four brightest IFOs are SNR-IFOs. On the other hand, the lowest surface brightness IFOs are also SNR-IFOs, showing the patchy [Fe II] emission in SNRs. All SNRs with [Fe II] emission features except one (G28.8+1.5) have been previously reported by Lee et al. (2019). The detailed results on the [Fe II] emission on the 19 SNRs can be found in Lee et al. (2019).

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Figure 1. Continuum-subtracted [Fe II] and NIR three-color images of IFOs with various origins: (a) YSO outflow: IRAS 18177-1405; (b) compact HII region: G35.2S; (c) diffuse HII region: GAL 10.2-0.3; (d) LBV nebula: [KW 97] 37-17; (e) PN: IRAS 19234+1627; and (f) SNR: W49B. Grey-scale images in the upper rows are UWIFE [Fe II]-H images; Colour-composite images in the lower rows are R/G/B = *KHJ*-band images from the UKIDSS GPS survey. The units of the UWIFE [Fe II]-H images are DNs, with the darker colour denoting a higher DN. The UWIFE images of the panels (a) IRAS 18177-1405 and (d) [KW97] 37-17 are smoothed with a two-pixel Gaussian. In all images, North is at the top, and east to the left side. Note the following artifacts: panel (a) IRAS 18177-1405: diffraction spike from southwest to northeast; (b) G35.2S: crosstalk on the northwest edge of the source; diffraction spikes and an airy disk at the south; (e) IRAS 19234+1627: dead pixels on the north and southwestern part at the boundary of the source; (f) W49B: masked bright stars.

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Figure 2. (a) F_{tot} distribution of IFOs. Note that the flux of a large-scale IFO 7 is excluded in this figure. (b) Semi-major axis r_1 distribution of IFOs. The semi-major axis of automatically-identified IFOs is the best estimate of the coordinate, semi-major, and minor axes from the best-fitting ellipse from IDL procedure 2dgaussfit. (c) Surface brightness distribution of IFOs. The counterparts of IFOs are presented in different colors.

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Figure 3. The spatial distribution of IFOs in Galactic longitude and latitude. The top panel shows the distribution of IFOs in Galactic longitude. The bottom panel shows the distribution of IFOs in Galactic latitude. IFOs are shown in accordance with their counterparts: YSO, HII region, PN, SNR, LBV, and unknown-IFOs.



Figure 4. Two-dimensional distribution of IFOs. The top panel shows the spatial distribution of IFOs having counterparts in the HII region (orange), YSO (blue), and unknown (gray) categories. Each circle represents an IFO, and the size of each circle is proportional to its logarithmic F_{tot} (in order of 10^{-17} , 10^{-16} , 10^{-15} W m⁻²). Star-forming regions whose positions match those of IFOs in the distribution are labelled. Due to clustered IFOs, many circles overlap. The bottom panel shows the spatial distribution of IFOs with counterparts of PNe (gray), SNRs (green), and LBVs (red). Note that the flux of IFO 7 (i.e., SNR G8.7-0.1) is not provided, therefore excluded here.

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Figure 5. IFOs with YSO counterpart candidates in continuum-subtracted [Fe II] images as in Fig. 1. Only six representative IFOs are shown. The yellow crosses denote adjacent YSOs in the field of view, while the red contours are H₂ 2.12 μ m emission contours adopted from UWISH2. The right frames are three-color *KHJ* UKIDSS images of the same field of view. This figure is available in its entirety in the Appendix B.1.

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Figure 6. IFOs with HH counterpart candidates. The left panels show continuum-subtracted [Fe II] images as in Fig. 1. The right panels show H α in the same field of view. IPHAS images were used except for IFO 41–44 where the *Hubble Space Telescope (HST)* F657N image was used. Inset on the IFO 119–121 is a magnified [Fe II]-H image of the saturated star (West of IFO 120, North in the inset) and bright part of IFO 119 (South in the inset). Red contours are H₂ 2.12 μ m emission adopted from UWISH2. The black arrow points to the driving source of the HH object. When the driving source is out of the image field of view, the arrow points from the driving source to the IFO.

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Figure 7. IFOs with CHII counterpart candidates. The left panels show continuum-subtracted [Fe II] images as in Fig. 1. The right panels show a radio continuum in the same field of view. Cyan contours on both images are the boundaries of CHIIs in the radio. The contours of IFO 24–26, 54, 71, 86–87, 97, 132, and 137 are from New-GPS 20 cm, IFO 89–90, 107, 127, 138, 148–149, 156, and 202–203 are from CORNISH 5 GHz, and IFO 82–83 are from the National Radio Astronomy Observatory (NRAO) Very Large Array (VLA) Sky Survey (NVSS). Only the radio image and contour of IFO 204 are from the old-GPS 20 cm. IFO 203 is a YSO-IFO inside the field of view. The yellow crosses in both panels are the same as in Fig. 5. The red cross shows the central position of the UWISH2 H₂ emission.



Figure 8. IFOs with Hu region counterpart candidates in continuum-subtracted [Fe II] images as in Fig. 1. Cyan contours are boundaries of HII regions in the radio continuum: IFO 11, 12, 55, 62 with New-GPS 20 cm data; IFO 58, 64, 66, 67 with GPS 90 cm data; 158, 159, 161 with THOR 1420 MHz continuum + VLA Galactic Plane Survey (VGPS) H I data. Black arrows point to the boundaries of IFO structures. The format for these images is the same as that of Fig. 1.



Figure 9. Continuum-subtracted [Fe II] images of IFOs with PN counterpart candidates. Only six representative IFOs are shown. The units on the UWIFE [Fe II]-H are DNs, with the darker colour denoting a higher DN. The right frames are three-color *KHJ* UKIDSS images of the same field of view. The corresponding source names for each IFO are shown. The yellow cross marks the central position of the counterpart. The images of IFO 73, 95, 112, 129, 157, and 164 are smoothed with a two-pixel Gaussian. In all images, North is at the top and east is on the left side. This figure is available in its entirety in Appendix B.2.

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Figure 10. Continuum-subtracted [Fe II] images of IFOs having LBV counterparts. The right frames are three-color *KHJ* UKIDSS images of the same field of view. Note that there is a spike pattern around a bright star, coincident with IFO 65. The format for these images is the same as that of Fig. 1.

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Figure 11. IFOs with counterpart candidates unknown in continuum-subtracted [Fe II] images as in Fig. 1. Only six representative IFOs are presented. The right frames are three-color *KHJ* UKIDSS images of the same field of view. The format for these images is the same as that of Fig. 1. This figure is available in its entirety in Appendix B.3.

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Table 1. Catalog of identified IFOs

IFO UWIFE designation I D b t t t test test test test IFO 001 18013 6927-224228.03 7.19190 40.062 2.3 2.0 0 1.44 0.89 - IFO 002 18021 555-24223.00 7.19190 40.062 2.3 2.0 0 1.44 0.89 - 5.88 - 5.88 - 5.88 - 5.88 - 5.88 - 5.88 - 5.88 - 5.81 1.00 1.00 1.0011 1.0012 555-2423.05 5.93 5.001 1.001										
FO 0017 180136.927-224228.03 7.1910 40.04007 9.3 2.9 40 8.47 5.88 - FF0 002 1180210.555-114236.69 8.11119 40.337 2.3 2.0 0 1.44 0.89 5. FF0 004 118021.9380-113230.55 8.2721 40.4830 5.4 4.2 120 71.2 4.27 YS FF0 005 118051.6644-19920.77 10.0386 60.73511 160.0 7.3 60 366.9 6.04 PN FF0 006 118051.6644-19937.73 10.0383 2.022 61 4.9 2.9 9.3 3.30 PN FF0 000 118064.84-0373.78.0 7.01503 -0.222 61.4 9.1 9.0 150 3.31.5 7.1 11.0 FF0 010 118064.84-0373.78.0 7.02267 -0.9524 2.34 4.00 4.8 1.9 YS FF0 010 11810.2015-00274.83 9.7066 -0.9529 3.8 170 70.4 2.81 XS <	IFO #	UWIFE designation	l [de	b eg]	r ₁ [arcs	r ₂ ec]	PA [deg]	Area [arcsec ²]	F_{tot} [10 ⁻¹⁷ W m ⁻²]	Counterpart
H+0100 1801692/-122303 1/1910 0.00023 9.3 2.9 40 84.7 5.88 - HF0002 18021238-2270.49 7.3338 0.0164 2.6 1.8 157 1.47 2.25 20 0 1.44 0.89 - HF0005 180511.688-1630-53 6.8721 0.4849 1.60 7.3 0 6.64 9.64 21 7.12 4.27 Y SS HF0005 18051.648-1980277 1.00356 0.7311 1.60 7.3 0 3.63.0 PN HF0005 18064.635-2012.050 S.2594 -0.6232 6.1 4.9 25 9.3 3.8 0 1.63 0.118 1.59 3.31.5 7.210 HI 1.60 1.30 4.33 1.60 1.118		7	-		-	-				
IP 002 IB 002 <thib 002<="" th=""> <thib 002<="" th=""> <thib 002<="" td="" th<=""><td>IFO 0019</td><td>J180136.927-224228.03</td><td>7.19190</td><td>+0.06028</td><td>9.3</td><td>2.9</td><td>40</td><td>84.7</td><td>5.88</td><td>-</td></thib></thib></thib>	IFO 0019	J180136.927-224228.03	7.19190	+0.06028	9.3	2.9	40	84.7	5.88	-
Prod Hill Hill <th< td=""><td>IFO 002 IEO 003</td><td>J180210.505-214520.09 ¹ J180212 308-223720 40</td><td>7 33358</td><td>+0.43376 -0.01604</td><td>2.5</td><td>2.0</td><td>157</td><td>14.4</td><td>0.89</td><td>- VSO</td></th<>	IFO 002 IEO 003	J180210.505-214520.09 ¹ J180212 308-223720 40	7 33358	+0.43376 -0.01604	2.5	2.0	157	14.4	0.89	- VSO
PEOD 05 INSO1 (1984-19902)77 OLOSE 0 OLOSE 0 <td>IFO 005</td> <td>J180219.380-213350.56</td> <td>8.26721</td> <td>+0.48309</td> <td>5.4</td> <td>4.2</td> <td>120</td> <td>71.2</td> <td>4.27</td> <td>YSO</td>	IFO 005	J180219.380-213350.56	8.26721	+0.48309	5.4	4.2	120	71.2	4.27	YSO
PEODOG 188054.464-19927.77 10.0135 0.0731 0.00 366.9 6.04 PS PEODOF 188060.397-220136.6 8.3594 -06230 4.5 4.2 20 59.3 13.80 PS PEOD09 188060.397-220136.6 8.3594 -06230 4.4 4.2 20 59.3 13.80 PS PEOD01 188075.220136.6 9.8489 -00238 3.4 3.4 0 5.6 3.0 10.91 11.81 10.91 11.81 10.91 10.833 -0.327 6.0 3.0 11.81 10.91 2.1 11.81 10.90 10.81 11.81 10.90 8.11 10.91 2.1 10.91 2.1 10.91 2.1 10.91 2.1 10.91 10.91 10.91 11.91 10.91 10.91 10.91 11.91 10.91 11.91 10.90 8.1 10.90 8.1 10.90 8.1 10.90 11.91 11.91 10.90 10.90 10.90 11.9	IFO 005	J180511.698-195040.56	10.09481	+0.74344	18.0	7.2	60	407.1	3.15	PN
IPO 07 H80627378-2122720 8.7590.3 0.3493 H600.0 0.6157521.8 - SNA IF0 080 H80640385-20126.95 8.3994.0 0.5232 6.1 4.9 2.5 9.3 1.380 PN IF0 000 H807533-2026.06 9.8489.0 0.0238 1.4 0.35.3 0.917 9.7 IF0 011 H807533-2026.06 9.8489.0 0.0258 1.4 0.83.15 1.214.3 1.66.00 HII IF0 011 H81050.44-205718.76 9.76257 -0.9542 4.2 3.4 1.60 4.4.8 1.59 YS IF0 015 H811129.621-48551.0 9.76656 -0.9529 3.8 170 7.04 4.53 0.53 0.35 - 1.7 1.5 0.5 0.044 YS IF0 018 H81407.762-H8510.73 H.9882 -0.0264 1.7 1.5 0.5 0.44 YS 1.8 1.8 1.8 1.2 1.7 1.6 2.5 1.9 1.5 1.5 1.6	IFO 006	J180514.644-195027.77	10.10356	+0.73511	16.0	7.3	60	366.9	6.04	PN
PFO 068 J1806-0.357-220126.63 8.3293 -0.6239 4.5 4.2 20 99.3 [13.80] PR PFO 001 J180712.033-20266.06 9.8449 -0.02588 3.4 3.4 0 3.6.3 0.91 YS PFO 010 J180716.73-201852.06 10.1533 -0.3217 66.0 30.150 8.315.5 7.21.0 HII PFO 01 J18016.05-20125155.7 76227 -0.05474 2.2 1.60 4.44 1.59 YS PFO 015 J1813.02-021-6153.2 J18.11 0.511676 1.31.4 1.511.0 5.11.75 1.11.0 5.11.73 D.6005 S.11.7 0.64 YS PFO 015 J1814.07.52-1851.0.73 J1.98433 -0.62012 J3.3 2.11 1.10 2.17.7 0.64 YS PFO 021 J18144.023-1561.05 J2.80944 -0.2602 J18.444.2 J2.91.75 J1.95 YS YS J1.57 J1.57 YS J1.10 2.11.7 J1.44 J1.57 J1.57 <	IFO 007	, J180627.378-213227.20	8.75903	-0.34293	1400.0	1400.0	0	6157521.8	-	SNR
FO 000 J 18060-085-20126.05 8.3620 -0.6222 6.1 4.9 2.5 9.9.9 6.10 PM FPO 010 J 18070-6373-20206.06 9.8489 -0.238 3.4 3.4 0 3.5 T2.10 HI FPO 01 J 180075.93+-20194.10 1015813 -0.3594 7.3 2.5 0 3.14.1 3.16 0.4 4.8 1.59 YS FPO 015 J 18112.03 0.7 0.944 1.8101 1.57577 0.97666 0.52 9 3.07 0.944 YS FPO 015 J 18142.124-164111.15 1.17737 1.0664 3.5 2.4 0.937 0.944 YS FPO 017 J 18143.124-17555.00 1.280436 -0.0464 1.07 8.3 7.4 90 1.92 4.79 YS YS FPO 023 J 181443.145.121-15355.01 1.280436 -0.0464 3.0 2.0 0 0.577 0 2.24 1.295 YS YS FPO 023 J 181446.683-164	IFO 008	J180640.397-220136.63	8.35934	-0.62390	4.5	4.2	20	59.3	13.80	PN
IFO 010 J18072.433-202060.60 9.8489 -0.02588 3.4 3.4 0 3.6.3 0.91 YS IFO 011 J180925.793-201524.10 10.1581 -0.32594 67.3 62.5 10 1214.3 166.00 HII IFO 014 J18105.015-20745.39 9.76056 -0.95529 5.9 3.8 170 70.4 2.81 YS IFO 016 J18131.02 54.116.01 -0.07676 134.1 13.1 0.54.118.7 D00000 SNI IFO 016 J18142.02-16.117.5 1.1581.10 -0.376.6 13.4 13.1 0.53. 0.53 0.53 0.55 1.60.018 J181443.148-17552.80 12.80784 -0.1664 8.3 7.4 90 19.29 4.79 YS S0.0 0.64.4 YS 19.29 4.79 YS S0.0 8.0 10.0 8.3 10.0 5.3 0.53 8.0 10.0 1.7 1.55 1.50 1.0 8.0 10.0 1.50 1.50 1.50	IFO 009	J180640.585-220126.95	8.36205	-0.62322	6.1	4.9	25	93.9	6.10	PN
HP 011 118091-6.3/±-20182-06 10.1503 -0.221/9 6.80 9.90 150 8331.5 12.10 HI IFO 012 118025593-2019541 0.15813 -0.5526 4 2 3.4 160 44.8 1.59 YSS IFO 015 1181129 652-05783.76 9.7665 -0.552 5.9 3.8 170 7.04 2.81 YSS IFO 015 1181224-164111.54 13.7797 -0.05664 3.5 2.8 90 30.7 0.94 YSS IFO 015 11814342-14-164111.54 11.95535.61 12.89454 -0.0464 3.3 1.10 2.17 0.64 YSS IFO 020 11814145.121-17557.61 12.89456 -0.20644 1.7 8.3 1.02 2.79.0 1.57 YSS IFO 023 118143422.22 12.89085 12.89139 -0.46471 3.0 2.7 1.60 0.56.9 YSS IFO 023 118143422.22 12.89085 1.27 1.60 1.57 1.66 </td <td>IFO 010</td> <td>J180732.633-202606.06</td> <td>9.84849</td> <td>-0.02588</td> <td>3.4</td> <td>3.4</td> <td>0</td> <td>36.3</td> <td>0.91</td> <td>YSO</td>	IFO 010	J180732.633-202606.06	9.84849	-0.02588	3.4	3.4	0	36.3	0.91	YSO
Inton.2 Inton.2 <t< td=""><td>IFO 011</td><td>J180916.373-201852.96</td><td>10.15033</td><td>-0.32179</td><td>68.0</td><td>39.0</td><td>150</td><td>8331.5</td><td>72.10</td><td>HII</td></t<>	IFO 011	J180916.373-201852.96	10.15033	-0.32179	68.0	39.0	150	8331.5	72.10	HII
In 00.14 Jislo1101=20374.83 9.7605 -0.98529 5.9 3.8 170 70.4 2.21 Yiss IFO 015 JISL129.62-129515.52 11.1811 -0.94766 JI.14 JI.1 0.51118 0.000 SNI IFO 015 JISL122.42-16111.15 JISL22.096-174758.22 12.2064 +0.04152 JISL22.096 JISL22.096-174758.22 12.2084 +0.04152 JISL22.096-174758.22 JISL22.092-1756.01 JISL22.092-1756.01 JISL22.092-1757.01 JISL22.092-1757.01 JISL22.092-0175.01 JISL22.092-0175.01 JISL22.092-0175.01 JISL22.092-0175.01 JISL22.092-0175.01 JISL22.092-0175.01 JISL22.092-0175.01 JISL22.01 JISL22.01 JISL22.092-0175.01 JISL22.01 JISL22.01 JISL22.01 JISL22.01 JISL22.01 JISL22.02 JIS	IFO 012 IFO 013	J180925.793-201934.10 J181050.844-205738.76	0.15813	-0.35954	67.3	62.5 3.4	160	13214.3	166.00	HII VSO
FO 015 JISI 29 632-192515.52 JISI 1 -0.34766 JSI 4 JSI 4 JSI 7 1090.00 SNI IFO 016 JISI 2424-16411.54 JI 377007 +0.60864 JS 2.8 90 30.7 0.94 YSG IFO 018 JISI 407762-18201.73 119843 -0.62071 JT 15 60 8.0 0.44 YSG IFO 020 JISI 40877-18505.80 J2.2084 -0.19644 JS 3.2 279.0 J.77 YSG IFO 020 JISI 41521-17557.10 J2.2084 -0.06647 J.0 7.6 0.524 1.95 YSG IFO 022 JISI 437.371-16425.07 J3.8824 +0.28200 2.0 0 1507.9 6.60 UCH IFO 025 JISI 437.371-16425.07 J3.8824 +0.28200 2.40 2.00 0 1507.9 6.60 UCH IFO 026 JISI 475.660-12710.24 H41440 -0.07172 3.8 3.5 90 41.7 4.66 YSG IFO 030	IFO 015	J181051.015-205748.39	9.76056	-0.95829	5.9	3.8	170	70.4	2.81	YSO
FPO 016 1181312-424-164111.54 13.7007 -0.06084 3.5 2.8 90 3.07 0.94 YSS FPO 017 1181407.702-18501.73 11.98852 -0.62071 1.7 1.5 60 8.0 0.44 YSS FPO 018 1181407.762-18501.73 11.98852 -0.62071 1.7 1.5 60 8.0 0.44 YSS FPO 020 1181418-1752.01 12.80784 -0.206464 7.4 40 192.9 1.75 YSS FPO 021 118143422-1641.18.3 13.7718 -0.22602 76.6 0.25.4 1.95 YSS FPO 023 118143422-1641.18.3 13.7718 -0.22700 13.0 8.0 60 326.7 9.40 UCH FPO 024 118143422-1641.18.313718 14.5829 +0.46128 4.5 3.9 160 326.7 9.40 UCH FPO 027 1181521.196-16025.94 14.5829 +0.46128 4.5 3.9 160 326.7 9.40 14.7 4.665	IFO 015	J181129.632-192515.52	11.18511	-0.34766	131.4	131.1	0	54118.7	1090.00	SNR
IFO 017 JISI 322 096-174758: 22 IZ2604 -0.04152 J.3 1.3 0 5.3 0.25 - IFO 018 JISI 408 073-18508:00 I198882 -0.0201 J.7 1.5 60 8.0 0.44 YSS IFO 020 JISI 413.148-175528:03 J280784 -0.30644 8.3 7.4 90 192.9 4.79 YSS IFO 021 JISI 413.122-17557.61 I280436 -0.20644 10.7 7.8 122.9 J.77 YSS IFO 022 JISI 43683-164507.0 J38834 J23020 2.40 0 1057.9 6.60 UCR IFO 025 JISI 43663-164507.0 J38707 -0.27800 J3.0 8.0 60 326.7 9.40 UCR IFO 027 JISI 51.0 J1.64525.67 J3.8797 -0.27800 J3.0 8.0 60 326.7 9.40 UCR J1.6528 J0.0 J1.7 4.50 J1.7 4.66 YSS J1.7 J1.80 J1.7 J1.80	IFO 016	J181312.424-164111.54	13.77907	+0.60864	3.5	2.8	90	30.7	0.94	YSO
FIFO 018 J181407.762-185101.73 11.98833 -0.62071 1.7 1.5 60 8.0 0.044 YSS FFO 020 J181408.073-185058.01 11.28803-02071 1.7 1.5 60 8.0 0.044 YSS FFO 020 J181419.212-17557.61 12.8038-0.22602 67.6 4.27 160 00682.2 65.6 4.27 160 00682.2 65.6 YSS FFO 023 J1814322-1651.41527.05 13.83771.8 0.027580 13.0 0.6 00 150.79 6.60 UCH FFO 025 J18143521.196-160255.94 14.5829 +0.46128 4.5 3.9 160 35.1 7.16 JSS 00 0.3 36.6 0.3 2.667 0.3 0.417 4.667 YSS 1.0 0.3 1.8 0.7 1.4 90 7.4 0.3 3.5 0 1.4 0.445 0.08 YSS 1.0 1.8 1.0 1.8 1.0 1.0 1.1 1.0 1.1	IFO 017	J181322.096-174758.22	12.82064	+0.04152	1.3	1.3	0	5.3	0.35	-
IFO 019 J181408 073-188058.01 11.98822 -0.62071 1.7 1.5 60 8.0 0.44 YSS IFO 020 J181413-L81-755230.01 12.8084 -0.2064 7.3 7.4 90 90.29 1.57 YSS IFO 021 J181413-182-175257.61 12.8084 -0.2064 7.6 0.058.2 6.95.0 YSS IFO 022 J181432 023-18503.61 12.3093 -0.46647 3.0 2.7 6.0 0.057.9 6.60 UCH IFO 022 J181436 683-161507.0 13.8234 -0.22500 2.40 0.0 0 1507.9 6.60 UCH IFO 022 J181436 683-161507.0 13.8234 -0.2250 1.3 8.3 90 41.7 1.466 YS IFO 021 J181726.551-172216.03 13.65638 -0.600717 1.8 3.5 90 1.45 0.88 YS IFO 031 J181728.451-172216.03 13.65638 -0.60179 2.0 1.3 0 8.1 0.35 YS	IFO 018	J181407.762-185101.73	11.98433	-0.62012	3.3	2.1	110	21.7	0.64	YSO
HO 020 JIRI 13, JIA=175528.03 12.20784 -0.19604 8.3 7.4 90 192.9 4.7 YS HFO 021 JIRI 145, JIL=175575.61 12.80485 -0.20684 10.7 8.3 12.0 90 92.9 4.7 160 9068.2 69.50 YSS HFO 022 JIRI 42223-H2820851 12.39193 -0.46647 30. 160 100.3 3.18 UCH HFO 025 JIRI 43.683-104507.70 13.88234 40.28200 24.0 20.0 0 1507.9 60 UCH HFO 025 JIRI 43.6683-104507.70 13.88234 40.28200 24.0 20.0 0 55.1 7.16 PS HFO 027 JIRI 52.1096-160255 91 41.58239 44.440 -0.0717 3.8 35 90 41.7 4.643 YS HFO 033 JIRI 750.953-1128192 13.4563 1.40071 3.1 40 38.9 2.0 90 14.45 98 YS 176.03 31.8175445-1027444 84.10674 50.8 1.50 1.0 1.5 1.60 1.5 1.5 1.6	IFO 019	J181408.073-185058.01	11.98582	-0.62071	1.7	1.5	60	8.0	0.44	YSO
Info U21 Ish Info. Li - Li 2537.01 L2 Marka L2 U L2 U L2 U L3 U L2 U L3 U	IFO 020	J181413.148-175528.03	12.80784	-0.19604	8.3	7.4	90	192.9	4.79	YSO
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 021	J181415.121-175557.61	12.80436	-0.20684	10.7	8.3	120	279.0	1.57	YSO
$ \begin{array}{c} 1.0 \\ 1.0 $	IFO 022 IFO 023	J101419.929-1/3010.05 J181422 928-182508 51	12.00898	-0.22002	3.0	42.7 27	100	9008.2 25.4	1 05	150 YSO
IFO 025 ^b II81436.683-164507.70 II83224 40.2200 24.0 20.0 1507.79 6.60 UCH IFO 026 ^b II81437.731-16452.678 II837970 40.2750 II.3 8.0 60 326.7 9.40 UCH IFO 027 II81521.196-1602554 IIS87970 40.2750 II.3 8.3 5 100 5.0 6.0 0.7 1.66 VI.17 1.66 YI.17 1.466 YSC IFO 028 II81638.050-162710.24 14.41440 -001717 3.8 3.5 90 14.17 4.66 YSC IFO 031 II81750.609-120805.80 IS.31683<+1.78972	IFO 023	J181434.822-164514 38	13.87718	+0.28764	8.0	4.0	160	100.3	3.18	UCHII
IFO 026 ^b 1181437.731-164526.78 13.87970 +0.27580 13.0 8.0 60 326.7 9.40 UCH IFO 027 1181521.196-16025.594 14.58529 +0.46128 4.5 3.9 160 55.1 7.16 JK IFO 029 1181658.050-162710.24 14.41440 -0.09117 3.8 3.5 90 14.17 4.065 YS IFO 030 1181750.059-120805.80 18.31683 +1.78972 1.7 1.4 90 7.4 0.43 YS IFO 033 1181758.455-12072.44 18.3405 +1.78952 2.7 2.5 0 21.2 2.11 YS IFO 035 1181828.50-16552.05 14.17066 -0.6193 3.0 8.1 0.54 YS IFO 036 118183.506-1552.721 15.83584 1.20527.7 1.58 1.51 YS 1.51 YS IFO 037 118183.6167-1502.71 15.83584 1.26 1.24 120 2.3 1.50 18.5 7.34 SNR IFO 037 1181837.058-13424.828 17.01527 +0.87673 1.24 <	IFO 025	, J181436.683-164507.70	13.88234	+0.28200	24.0	20.0	0	1507.9	6.60	UCHII
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 026	, J181437.731-164526.78	13.87970	+0.27580	13.0	8.0	60	326.7	9.40	UCHI
IF0 028 II81627.693-I8363.67 I2.45471 -0.99317 7.6 4.3 50 102.6 0.30 IF0 029 II81658.050-I62710.24 II8144440 -0.01712 3.8 3.5 90 41.7 4.66 YSC IF0 030 II81724.551-I72216.03 13.6658.8 -0.60071 4.0 3.1 4.0 3.8 2.01 YSC IF0 031 II81758.445-120802.39 I8.31683 +1.78792 1.7 1.4 90 7.4 0.43 YSC IF0 031 II8178.445-120724.48 I8.34208 +1.76705 2.3 2.0 90 1.445 0.98 YSC IF0 035 II81828.506-165523.72 1.41706 -0.61295 3.1 2.4 1.00 83.7 1.51 YSC IF0 037 II8183.6167-15027.12 1.58.88 1.052.7 1.58 8.021557 6.4 5.1 100 83.7 1.51 YSC IF0 039 II81845.167-15027.12 IS8.88 8.021557 6.4 5.1 1.00 8.1 0.37 1.51 9.8 9.4 0.49 YSC 1.660	IFO 027	J181521.196-160255.94	14.58529	+0.46128	4.5	3.9	160	55.1	7.16	PN
$ \begin{split} & FO 029 \\ FO 030 \\ FO 121 \\ FO 031 \\ FO 031 \\ FO 032 \\ I81750 953 \\ I81858 835 \\ I50 33 \\ I81750 953 \\ I81828 \\ I50 33 \\ I81828 \\ I50 148 \\ I8182 \\ I8183 \\ I8182 \\ I8183 \\ $	IFO 028	J181627.693-183653.67	12.45471	-0.99317	7.6	4.3	50	102.6	0.30	
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 029	J181658.050-162710.24	14.41440	-0.07172	3.8	3.5	90	41.7	4.66	YSO
$ \begin{array}{ c c c c c c c c c c c c c c c c c c c$	IFO 030	J181724.551-172216.03	13.65638	-0.60071	4.0	3.1	40	38.9	2.01	YSO
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 031	J181750.609-120805.80	18.31683	+1.78972	1.7	1.4	90	7.4	0.43	YSO
IFO 034 IIS 128-15-16/024-48 16.342/08 11.3013 2.0 13 0 8.1 0.54 YS IFO 034 IIS 1828.25-16-16523.72 14.17162 -0.61295 3.1 2.4 120 2.33 4.89 YS IFO 035 IIS 1832.55-01-16327.21 14.07162 -0.61295 3.1 2.4 120 2.33 4.89 YSS IFO 037 IIS 1835.53-134237.23 17.022771 +0.86910 2.8 2.5 120 2.13 0.83 YSS IFO 038 JIS 1835.530-14327.23 16.02393 +0.77973 1.2 1.2 0 4.5 0.23 YSS IFO 040 JIS 1845.8.01-13526.3.9 16.91039 +0.7467 2.0 1.5 120 9.4 0.49 YSS IFO 044 JIS 1958.8428-135145.51 16.91939 +0.770150 150 0.0 140 37.6 0.23 HH IFO 045 JIS 195.71-13452.291 17.0256 +0.75343 251 19.0 90 1498.2 4.14 YSS IFO 0469 JIS 1917.16435.78 14.30404	IFO 032	J181750.953-120802.39	18.31835	+1.78895	2.7	2.5	0	21.2	2.11	-YSO
IN 00.50 1181828.590-165523.72 14.17162 -0.631125 2.3 1.4 120 23.3 1.48 YSC IFO 035 J181828.590-165523.72 14.17162 -0.6125 3.1 2.4 120 23.3 1.51 YSC IFO 036 J181837.058-134248.28 17.01271 +0.86610 2.5 120 21.9 0.83 YSC IFO 040 J181847.167-15022.70 16.92393 +0.77973 1.2 1.2 0 4.5 0.23 YSC IFO 040 J181855.428-15514.551 16.91393 +0.77973 1.2 1.2 0 4.5 0.23 YSC IFO 041 J181855.428-15514.551 16.91393 +0.77973 1.2 1.2 0 4.5 0.23 HH IFO 042 J181855.428-15514.551 16.91393 +0.77913 4.0 30 140 37.6 2.24 HH IFO 044 J18190.85871-13452.51 16.91397 +0.7124 4.0 30 140 37.6 2.24 YSC IFO 045 J18190.85871-13452.79 1.10325 +0.75343 <	IFO 033	J181/38.443-120/24.48 J181828 251-165525.05	18.34208	+1./0/05	2.5	2.0	90	14.45	0.98	150 VSO
IFO 036 J181837.058-134248.28 17.01527 +0.87651 6.5 4.1 100 837 1.51 YSC IFO 037 J181839.539-134237.23 J7.02271 +0.86910 2.8 2.5 120 21.3 0.83 YSC IFO 038 J181845.167-150227.0 I6.92393 +0.77973 1.2 1.2 0 4.5 0.23 YSC IFO 040 J181849.365-13452.55 I6.91893 +0.77973 1.2 1.2 0 4.5 0.23 YSC IFO 041 J181855.428-135145.51 I6.91893 +0.74041 10.2 5.6 30 179.4 16.80 HH IFO 044 J18158.335-13522.81 16.90877 +0.71943 4.0 30 140 37.6 0.23 HS IFO 045 J181905.871-13452.291 17.03256 +0.73343 2.51 190 90 1498.2 4.14 YSC IFO 047 J18192.591-J3414.45 17.1257 40.7624 7.1 3.0 30 66.9 0.70 YSC IFO 047 J18192.591-J3454.271 17.06326 -0.	IFO 034	J181828.590-165523.72	14.17162	-0.61295	3.1	2.4	120	23.3	4.89	YSO
IFO 037 J181839.539-134237.23 17.02271 +0.86910 2.8 2.5 120 21.9 0.83 YSC IFO 038 J181845.167-150257.21 15.85388 +0.21557 6.4 5.1 130 102.5 7.34 SNR IFO 040 J181845.467-150257.21 15.9339 +0.77973 1.2 1.2 0 4.5 0.23 YSC IFO 041 J181855.301-135236.39 16.91199 +0.72350 10.0 6.0 150 188.4 2.65 HH IFO 044 J181805.8371-13452.91 17.03256 +0.75343 25.1 19.0 90 1498.2 4.14 YSC IFO 044 J18190.871-13452.91 17.03256 +0.75343 25.1 19.0 91 449.2 4.14 YSC IFO 047 J18191.708-164355.78 14.43321 -0.69611 4.0 3.0 90 37.6 2.24 YSC IFO 048 J18192.529-13414.45 17.12557 +0.7524 7.1 3.0 30 66.5 0.33 YSC IFO 047 J181917.38-161013.05 15.04831	IFO 036	J181837.058-134248.28	17.01527	+0.87651	6.5	4.1	100	83.7	1.51	YSO
FO 038 JI81845.167-150257.21 IS 85388 0.21557 6.4 5.1 130 O25. 7.34 SNR FO 039 JI81847.349-135022.55 16.93498 +0.77687 1.2 1.2 0 4.5 0.23 YSC IFO 040 JI81855.326-134952.55 16.91893 +0.77687 2.0 1.5 120 9.4 0.49 YSC IFO 041 JI81858.830-135226.30 16.90193 +0.77687 2.0 1.50 1884 2.65 HH IFO 044 JI81905.871-13452.91 17.03256 +0.73343 2.61 100 140 471.2 2.52 HH IFO 046 JI81905.871-13452.71 17.03256 +0.73343 2.51 19.0 90 1498.2 4.14 YSC IFO 047 JI81917.98-16499.13 14.34321 -0.66611 4.0 3.0 0.0 67.9 0.70 YSC IFO 048 JI81927.18-1542.71 17.06480 +0.68168 5.4 4.0 0 67.9 1.70 YSC IFO 051 ^b JI8202.0767-161018.45 5.0466 -0.65252	IFO 037	J181839.539-134237.23	17.02271	+0.86910	2.8	2.5	120	21.9	0.83	YSO
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 038	J181845.167-150257.21	15.85388	+0.21557	6.4	5.1	130	102.5	7.34	SNR*
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 039	J181847.449-135022.70	16.92393	+0.77973	1.2	1.2	0	4.5	0.23	YSO
IFO 041 J181855 428-135145.51 16.91893 +0.74041 10.2 5.6 30 179.4 16.80 HH IFO 042 J181858.301-135236.39 16.91199 +0.72350 10.0 6.0 150 188.4 2.65 HH IFO 044 J181905.871-134522.91 16.90173 +0.70150 150 100 140 471.2 2.52 HH IFO 045 J181905.871-13452.91 17.03256 +0.75343 25.1 190.90 1498.2 4.14 YSG IFO 0467 J181914.708-164949.13 14.3049 -0.73101 1.6 1.3 30 66.9 0.70 YSG IFO 048 J181927.591-13411.45 17.12557 40.72624 7.1 3.0 30 66.9 0.70 YSG IFO 051 ^b J18207.178-15121.16 15.79925 >0.00624 3.5 2.3 30 25.2 8.84 PN IFO 052 ^b J18202.767-161018.45 15.04565 -0.65124 2.2 2.0 0 13.8 1.33 YSG IFO 055 ^b J18202.767-161018.45 15.04566 -	IFO 040	J181849.365-134952.55	16.93498	+0.77687	2.0	1.5	120	9.4	0.49	YSO
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 041	J181855.428-135145.51	16.91893	+0.74041	10.2	5.6	30	179.4	16.80	HH
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 042 IFO 043	J181858.301-135250.39 J181858 835-135252 81	16.91199	+0.72330 ± 0.71943	10.0	0 .0	130	188.4	2.05	нн
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 045	J181901 895-135346 30	16 90173	+0.70150	15.0	10.0	140	471.2	2.52	НН
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 045	J181905.871-134522.91	17.03256	+0.75343	25.1	19.0	90	1498.2	4.14	YSO
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 046	¹ J181914.708-164949.13	14.34049	-0.73101	1.6	1.3	30	6.5	0.33	YSO
$ \begin{array}{llllllllllllllllllllllllllllllllllll$	IFO 047	J181917.916-164355.78	14.43321	-0.69611	4.0	3.0	90	37.6	2.24	YSO
$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	IFO 048	J181922.591-134114.45	17.12557	+0.72624	7.1	3.0	30	66.9	0.70	YSO
IFO 050 JI81927.118-151211.16 15.79925 -0.0624 3.5 2.3 30 25.2 8.84 PN IFO 051 ^b JI82019.871-161031.33 J5.04966 -0.65134 2.2 2.0 0 13.8 1.33 YSC IFO 052 ^b JI82020.767-161018.45 J5.04566 -0.65282 2.8 2.0 0 17.5 4.05 YSC IFO 054 ^b JI82021.725-161015.05 J5.04831 -0.6577 2.3 1.7 30 12.2 1.05 YSC IFO 055 ^b JI82021.736-161024.98 J5.09363 -0.67472 18.7 9.4 140 552.2 5.11 HCH IFO 055 JI82032.784-160124.98 J5.19955 -0.62538 7.6 3.9 0 93.1 0.63 YSC IFO 055 JI82035.196-161942.63 J6.19955 -0.62538 7.6 3.9 0 93.1 0.63 YSC IFO 056 JI82035.224-14043.64 16.92057 +0.28355 1.9 1.8 90 10.7 0.85 YSC IFO 0601 J182035.224-14043.84 16.92057 <td>IFO 049</td> <td>J181925.259-134542.71</td> <td>17.06480</td> <td>+0.68168</td> <td>5.4</td> <td>4.0</td> <td>0</td> <td>67.8</td> <td>1.01</td> <td>YSO</td>	IFO 049	J181925.259-134542.71	17.06480	+0.68168	5.4	4.0	0	67.8	1.01	YSO
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 050	J181927.118-151211.16	15.79925	-0.00624	3.5	2.3	30	25.2	8.84	PN
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 051	J182019.871-161031.33	15.04081	-0.65134	2.2	2.0	0	13.8	1.33	YSO
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 052'	J182020.767-161018.45	15.04566	-0.65282	2.8	2.0	0	17.5	4.05	YSO
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 053	J182021.725-161015.05	15.04831	-0.65575	2.3	1.7	30	12.2 552.2	1.05	YSO
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 054	J182024.450-101126.80	15.03583	-0.60915	18.7	9.4	140	552.2 78520.9	5.11	HCHII
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	TEO 055	J102020.170-101245.10 J182032 784_160124 00	15 19905	-0.09813	230.0 7.6	3.0	130	10239.8	1257.00	
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 050 IFO 057	J182034.306-160158 97	15,19359	-0.63521	7.0 8.0	2.0	120	95.1 50.2	0.03	YSO
$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	IFO 058	J182035.196-161942.63	14.93464	-0.77759	330.0	120.0	160	124407.0	473.00	HII
$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	IFO 059	J182035.224-140436.84	16.92057	+0.28355	1.9	1.8	90	10.7	0.85	YSO
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 060	J182035.656-140409.72	16.92803	+0.28556	5.8	5.1	30	92.9	7.82	YSO
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 061	J182036.014-140344.82	16.93481	+0.28754	3.3	3.3	0	34.2	0.54	YSO
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 062	J182037.224-160828.36	15.10369	-0.69649	341.8	197.3	115	211860.0	875.00	HII
$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	IFO 063	J182049.312-140353.86	16.95793	+0.23896	2.6	2.4	30	19.6	0.60	YSO
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	IFO 064	J182056.997–161934.88	14.97762	-0.85347	100.0	50.0	160	15707.9	58.27	HII
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 065	J182119.587-162224.78	14.97848	-0.95541	15.0	13.5	140	636.2	81.16	LBV
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 066	J182121./01-160424.16	15.24/3/	-0.82163	18.2	16.2	125	926.2	8.96	
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	1FO 067	1182228 548-171548 32	14 32188	-1.61616	23.3 3.5	0.1 2 3	30	048.8 25.2	11.40	-
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	IFO 068	J182432.827-130950.81	18,17823	-0.13740	5.5 101 0	2.3 56.0	50 70	23.2 17768 8	21.50	SNR
IFO 071 J1825459.449-131552.08 18.14002 -0.27977 130.0 90.0 60 36756.6 110.10 UCH IFO 072 ^a J182548.520-130629.85 18.37107 -0.38288 7.0 4.5 120 98.9 82.00 YSC	IFO 070	J182448.037-131345.15	18.14955	-0.22238	6.0	5.0	90	94.2	1.38	SNR
IFO 072 ^a J182548.520-130629.85 18.37107 -0.38288 7.0 4.5 120 98.9 82.00 YSC	IFO 071	J182459.449-131552.08	18.14002	-0.27977	130.0	90.0	60	36756.6	110.10	UCHII
	IFO 0726	¹ J182548.520-130629.85	18.37107	-0.38288	7.0	4.5	120	98.9	82.00	YSO

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Table 1 (cont'd)

IFC) #	UWIFE designation	l [d	b	r ₁	r ₂	PA [deg]	Area	F_{tot} [10 ⁻¹⁷ W m ⁻²]	Counterpart
	22	1192610 105 101219 67	20.092(1	.0.95202	10.0	7.4	120	1(2)(DN
IFO 0 IFO 0)74	J182656.992-113210.92	20.98261 19.89167	+0.85302	4.2	7.4 3.4	60	462.6 44.8	0.64 4.51	PN YSO
IFO 0	075	J182851.018-124415.55	19.04409	-0.86640	67.4	27.4	15	5801.7	14.10	SNR
IFO 0	076	J182852.671-124311.10	19.06306	-0.86404	16.4	5.5	0	283.3	2.55	SNR
IFO 0)77 ^a	J182859.486-115026.04	19.85474	-0.48056	6.1	2.3	80	44.0	3.45	YSO
IFO 0)78	J182919.659-124153.90	19.13295	-0.95127	132.0	89.4	140	37073.3	192.00	SNR
IFO)/9 >>>0b	J182930.563-131350.84	18.68160	-1.23/35	8./	3.0	/0	98.3	1.89	SNR
IFO	180 181 ^b	J183314.281-100831.20 J183328 075-110726 68	21.84285	-0.01852	11.8	200.0	40 25	314159.2 237.2	49.30	DN
IFO	382^{a}	I183328.975=110720.08	26.339907	+1 69464	6.8	3.5	165	237.2	49.30	UCHII
IFO 0)83 ^a	J183330.673-050110.94	26.42214	+1.68483	15.4	12.0	114	580.5	150.00	UCHII
IFO 0	084	J183331.327-103257.93	21.51352	-0.86841	35.0	17.0	50	1869.2	6.59	SNR
IFO 0)85	J183333.430-103402.86	21.50149	-0.88437	44.8	42.1	90	5925.2	59.70	SNR
IFO 0)86	J183404.384-071820.28	24.45479	+0.50637	60.0	30.0	160	5654.8	9.42	UCHII
IFO 0)87	J183408.045-071801.82	24.46632	+0.49530	12.7	7.6	30	303.2	1.72	UCHII
IFO 0)88	J183420.390-084722.27	23.16829	-0.23606	600.0	400.0	10	753982.2	584.00	SNR
IFO	189	J183425.284-075448.33	23.95514	+0.14969	11.0	9.5	20	321.3	0.32	UCHII
IFO)90)91	J183541 856-072203 61	23.90283	+0.14077	1.1	8.0	140	376 9	3.06	YSO
IFO 0)92	J183648.912-071850.94	24.76014	-0.10123	2.3	1.9	20	13.7	2.13	YSO
IFO 0)93	J183716.440-032958.39	28.20160	+1.54899	4.0	2.3	10	28.9	0.39	SNR
IFO 0)94	J183720.713-064200.69	25.36590	+0.06398	2.0	1.7	30	10.6	1.18	YSO
IFO 0)95	J183730.398-061412.49	25.79595	+0.24115	8.0	8.0	0	201.0	3.93	PN/UCHII
IFO 0)96 ^b	J183740.829-061452.41	25.80594	+0.19768	180.0	180.0	0	101787.6	221.00	-
IFO 0)97	J183813.600-064815.32	25.37390	-0.17819	25.2	8.0	45	633.3	25.60	UCHII
IFO 0	198 ⁰	J183907.168-043230.84	27.48618	+0.66204	500.0	400.0	20	628318.5	167.00	SNR
IFO 0	199 100	J183909.562-071927.89	25.01779	-0.62238	4.0 2 1	3.0 2 0	0	37.6	2.05	YSO
IFO I IFO 1	101	J183913.302-072019.31	25.00933	-0.64748	3.4 1.8	∠.ŏ 15	90 130	29.9	2.56	YSO
IFO 1	102	J183931.338-054409.74	26.47082	+0.02555	3.3	2.3	80	23.8	3.07	LBV
IFO 1	103	J183931.437-054414.64	26.46980	+0.02456	2.4	1.4	0	10.5	1.27	LBV
IFO 1	104	J183950.426-043037.60	27.59648	+0.51675	4.5	2.0	0	28.2	-0.20	YSO
IFO 1	105	J184120.333-045606.47	27.38989	-0.00960	150.0	120.0	70	56548.6	274.00	SNR
IFO 1	106	J184358.299-035306.11	28.62385	-0.11304	216.2	170.1	0	115534.0	197.00	SNR
IFO 1	107	J184414.391-041754.32	28.28667	-0.36139	7.0	3.6	165	79.1	2.70	UCHII
IFO I	108	J184422.810-041/34./8	28.30/48	-0.39003	3.9	2.8	90	34.3 17.1	0.57	480
IFO 1	109	I184515 462-031604 01	29 31952	-0.11656	43	3.5	120	47.2	0.37	-
IFO 1	111	J184559.282-024502.58	29.86281	-0.04271	7.3	3.7	135	84.8	5.47	YSO
IFO 1	112 ^a	J184829.526-021003.18	30.81541	-0.25716	2.9	1.5	150	13.6	0.52	PN
IFO 1	113	J184927.068-005638.37	31.86530	+0.01156	240.0	210.0	120	158336.2	3423.55	SNR
IFO 1	14 ^{<i>ab</i>}	J184933.121-003810.21	32.59821	+0.35877	8.9	6.3	156	176.1	6.46	PN/HII
IFO 1	115	J184955.670-010153.39	31.84176	-0.13438	1.9	0.9	130	5.3	0.25	YSO
IFO 1	116^{ab}	J185026.138+012739.08	33.81379	+0.73328	5,5	2.1	68	36.2	0.83	-
IFO 1	1170	J185125.777-000930.42	32.78995	-0.07040	740.0	530.0	170	1232132.7	114.00	SNR
IFO 1	118	J185128.102+002840.15 J185140.610±002850.90	33.36065 33.38709	+0.21110	8.6	8.0	170	216.1	1.02	HH UU
IFO I IFO 1	20	J185141.136+002850.89	33,39049	+0.16537	1.0	12	90	70.0 5.2	0.24	НН
IFO 1	121	J185144.114+002911.43	33.39880	+0.15568	7.0	4.0	90	87.9	1.64	НН
IFO 1	122	J185249.992+022802.16	35.28688	+0.81455	1.4	0.9	90	3.9	0.56	YSO
IFO 1	123	J185251.980+022804.91	35.29134	+0.80753	1.2	1.0	90	3.7	0.74	YSO
IFO 1	124	J185353.538+015714.24	34.95065	+0.34503	3.9	2.8	60	34.3	1.87	YSO
IFO 1	125	J185516.571+030512.10	36.11639	+0.55408	8.0	4.8	160	120.6	9.03	YSO
IFO 1	126	J185521.357+030154.61	36.07665	+0.51134	1.5	1.4	175	6.5	0.42	YSO
IFO 1	127"	J185534.205+021908.12 J185602 316±012120 70	33.51833	+0.06246	6.3	4.2	175	83.1	1.46	UCHII SNP
IFO 1	129 ^b	1185737 231±020350 02	35 47351	-0.43365	11.0	11.0	150	380.1	5742.00	PN
IFO 1	130 ^b	J185738.014+020330.02	35.47060	-0.43876	11.0	11.0	0	380.1	7 11	PN
IFO 1	131	J185808.531+010048.26	34.59832	-1.02921	4.5	4.5	0	63.6	3.22	YSO
IFO 1	132 /	J185810.531+013656.88	35.13815	-0.76162	17.5	14.7	0	808.1	24.60	CHII
IFO 1	133	J185905.114+004833.51	34.52432	-1.33211	2.1	1.4	0	9.2	1.60	-
IFO 1	134	J185910.539+014013.31	35.30084	-0.95905	5.0	3.1	130	48.7	1.56	YSO
IFO 1	135	J185923.204+010413.33	34.79106	-1.27998	2.8	2.8	0	24.6	0.39	-
IFO 1	136	J190003.046+055926.63	39.24352	+0.82130	26.6	13.0	170	1086.3	3.29	YSO
IFO 1	137	J190347.070+050946.79	38.93288	-0.38355	35.0	18.0	100	1979.2	8.31	UCHII
IFO I	130	J190403.033+050/53.38	30 22125	-0.45909	5.2 107.2	4.4	30	/1.8	2.74	SND
1FO 1	40	J190540.231+074634.40	41.46973	+0.39930	550.0	450.0	100	93037.4 777544 2	1421 45	SNR
IFO 1	141	J190659.919+052253.12	39.49429	-0.99412	5.4	5.2	90	88.2	12.20	YSO
IFO 1	142	J190731.328+052333.18	39.56407	-1.10471	5.5	4.4	130	76.0	1.17	-
IFO 1	143	J190734.278+070829.06	41.12259	-0.31090	170.0	120.0	115	64088.4	1691.00	SNR
IFO 1	144	J190813.552+052757.00	39.70977	-1.22650	1.6	1.4	90	7.0	0.23	YSO

Table 1 (cont'd)

ШО #		1	1			DA	A		G
IFU #	U WIFE designation	í [di	<i>b</i> egl	r ₁ fare	r ₂ secl	PA [deg]	Area [arcsec ²]	$[10^{-17} \text{ W m}^{-2}]$	Counterpart
		- Lu		Lare		(61			
FO 145	J190816.446+052726.79	39.70784	-1.24102	3.7	1.6	125	18.5	0.40	YSO
FO 146	J190816.782+052506.10	39.67378	-1.26024	3.9	3.6	120	44.1	1.69	YSO
50.147 50.148^{a}	1191100.840+090004.55 1191327 650+105334 62	45.20393	+0.13279	11.2	8.5	120	299.0	4739.00	UCHII
0.140^{a}	J191327.754+105413.66	45.13129	+0.13745	3.2	2.4	105	255.0	0.48	UCHII
O 150	J191530.963+132747.36	47.63089	+0.88215	1.7	1.4	120	7.4	0.29	-
O 151	J192026.005+111955.24	46.30789	-1.17527	5.6	3.9	120	68.6	2.25	YSO
O 152	J192029.411+111942.04	46.31118	-1.18928	1.3	1.1	90	4.5	0.27	YSO
O 153	J192029.485+110159.44	46.05061	-1.32806	7.5	5.3	90	124.8	8.71	HH
O 154	J192054.201+143031.29	49.16624	+0.21581	1.6	1.6	0	8.0	0.75	YSO
0 155	J192113.714+105232.92	45.99659	-1.56168	4.2	2.8	110	36.9	2.34	HH
0 156	J192127.938+154426.63	50.31742	+0.67543	4.5	2.6	30	36.7	1.58	UCHII
0 15/	J192142.900+155351.18	50.48401	+0.69629	60.0	20.0	90	3/69.9	1.55	PN
2 158	J192309.853+142912.05	49.40473	-0.27709	40.0	1000.0	160	2760011.2	1647.00	
0 159	J192233.023+140743.93	49.00144	-0.59509	580.0	430.0	30	783513.2	582.37	SNP
D 161	J192348.822+143137 35	49.51449	-0.39670	350.0	210.0	115	230907.0	182.47	HII
D 162	J192348.169+143641.50	49.58771	-0.35441	50.0	30.0	25	4712.3	45.50	LBV
D 163	J192354.032+143548.00	49.58825	-0.38096	45.0	35.0	0	4948.0	5.76	YSO
0 164	J192451.838+155729.06	50.89493	+0.05763	1.5	1.5	0	7.0	0.24	PN
O 165	J192516.759+144625.72	49.89974	-0.59204	2.3	1.9	0	13.7	0.72	YSO
O 166	J192529.675+151646.36	50.36959	-0.39785	1.1	0.9	0	3.1	0.20	-
O 167	J192531.202+151603.90	50.36214	-0.40884	2.1	1.3	0	8.5	0.79	-
0 168	J192531.399+151556.79	50.36075	-0.41049	0.8	0.8	0	2.0	0.13	-
169	J192532.882+151538.18	50.35903	-0.41819	1.5	1.3	40	6.1	0.50	-
171	J192533.41/+151616.62	50.36947	-0.41501	1.1	1.1	0	3.8	0.26	(
) 172	I192534 538+151632 38	50 37544	-0.41690	1.2	1.4	90	5.0 7.0	0.23	
2173^{a}	J192540.546+163305.18	51.50973	+0.16761	7.0	5.8	16	127.5	11.50	PN
0 174	J192547.157+145145.84	50.03612	-0.65762	4.0	2.5	140	31.4	4.22	YSO
0 175	J192557.625+150231.65	50.21401	-0.60951	1.8	1.6	90	9.0	0.36	YSO
0 176	J192557.848+150243.23	50.21727	-0.60877	10.0	5.5	0	172.2	1.98	YSO
0 177	J192852.403+171458.61	52.48804	-0.17205	3.3	2.8	90	29.0	0.92	YSO
O 178	J192918.342+175615.42	53.14142	+0.06679	15.0	10.0	140	471.2	4.95	YSO YSO
D 179	J192918.796+175723.68	53.15891	+0.07429	3.7	1.9	130	22.0	0.31	YSO
O 180	J192920.127+175716.54	53.15971	+0.06871	25.0	8.0	80	628.3	4.41	YSO
0 181	J192920.506+175458.14	53.12668	+0.04898	7.0	5.0	110	109.9	1.11	480
> 162 > 183	J172922.491+1/4442.54 J192931 617+175951 30	53 21927	+0.03983	4.0 4.1	3.3 3.0	0	20.5 41.4	1.08	- V\$0
D 184	J192931.871+180058 11	53,23604	+0.05734	1.7	1.4	90	7.4	0.25	YSO
D 185	J192932.874+180106.35	53.23994	+0.05495	7.0	4.0	45	87.9	1.06	YSO
O 186 ^b	J193001.921+175455.44	53.20473	-0.09547	5.4	3.3	20	55.9	3.15	YSO
O 187	J193120.744+192014.92	54.60141	+0.31496	1.5	1.3	90	6.1	0.35	-
O 188	J193323.546+195647.07	55.36730	+0.18676	20.0	18.0	40	1130.9	5.22	post-AGBc/YSO
O 189	J193831.665+202519.19	56.36978	-0.63373	7.0	4.3	150	94.5	10.50	YSO
O 190	J193914.355+224021.52	58.41141	+0.32789	7.0	6.0	¥ 40	131.9	3.75	YSO
D 191	J194014.058+232652.51	59.19889	+0.51050	5.5	4.2	80	72.5	1.57	-
) 192	J194103.922+220340.80	58.08778	-0.34083	3.2	1.7	0	17.1	0.62	YSO
) 193 > 104	J194127.149+222739.58	50.00574	-0.22095	14.2	6.8	130	303.3	14.70	¥80
) 194) 105	J174241.010+225417.72	50 12550	-0.24738	0.5	3.0	00	01.2	2.70	130
) 195	J194256.665+232435.17	59.47362	-0.02522	29.0	5.0 18 5	40	20.2 1685 4	5.20	НН
D 197	J194306.295+231810.63	59.39926	-0.13356	1.9	1.3	90	7.7	0.52	-
O 198	J194310.286+234358.03	59.77970	+0.06707	4.9	2.9	60	44.6	2.83	YSO
D 199	J194310.930+234402.64	59.78203	+0.06557	5.3	2.5	90	41.6	2.23	YSO
O 200	J194320.930+232952.89	59.59633	-0.08502	10.5	8.0	150	263.9	1.34	HH
O 201	J194610.902+221559.08	58.85575	-1.26581	7.1	2.5	120	55.7	0.96	-
O 202	J194620.335+243520.73	60.88253	-0.13043	14.0	12.9	80	567.3	17.20	UCHII
O 203	J194621.675+243516.78	60.88413	-0.13538	1.3	1.3	0	5.3	0.45	YSO
O 204	J194646.921+251241.33	61.47104	+0.09568	40.5	36.5	140	4644.0	111.00	UCHII / HII

^aIFOs marked with 'a' are identified only by an automatic detection method.

^bNote on the individual sources. IFO 7: Due to the complexity of the region, flux is not provided. IFO 25–26: The flux of the superposed part is allocated only to IFO 26. IFO 51–55: IFO 51–54 are located inside IFO 55. IFO 62, 98: The flux is derived for a partial region free from severe artifacts. IFO 65, 81, 114, 116, 186: missing flux due to 2MASS-bright star mask. IFO 80, 96, 117: Contaminated by an instrumental artifact. The pixels with DN > $\pm 3\sigma$ are masked for the flux measurement. IFO 129–130: Contaminated by an instrumental artifact. The pixels with DN > $\pm 3\sigma$ are masked for the flux measurement. IFO 139: Contaminated by an instrumental artifact. The pixels with DN < -2σ are masked for the flux measurement. IFO 139: Contaminated by an instrumental artifact. The pixels with DN < -2σ are masked for the flux measurement. IFO 159–160: The flux of the superposed part is allocated only to IFO 160. Note that there is an astrometry problem with certain continuum-subtracted [Fe I] images, where IFO 4, 8, 9, 73, 114, 155, 165, and 186 are located. Therefore, we determined the central positions of the IFOs based on the UKIDSS NIR image.

*IFO 38 is located within the SNR G15.9+0.2 domain but highly confined to a southwestern region (see fig. 1 of Sasaki et al. 2018 for an X-ray image of the SNR). Since there is no other possible counterpart in the SIMBAD query and the X-ray emission is coincident, we concluded that the SNR origin cannot be ruled out.

					-	
	N	Flux _{total}	Flux _{min}	Flux _{max}	Flux _{mean}	F ^{sb} mean
YSO	100	4.3 (-15)	2.0 (-18)	8.2 (-16)	4.3 (-17)	6.6 (-19)
CHII	22	5.2 (-15)	3.2 (-18)	1.5 (-15)	2.4 (-16)	4.0 (-19)
ΗII	11	4.8 (-14)	8.9 (-17)	1.6 (-14)	4.3 (-15)	0.7 (-19)
PN	17	1.4 (-15)	2.4 (-18)	4.9 (-16)	8.5 (-17)	7.6 (-19)
SNR	25	2.0 (-13)	3.9 (-18)	4.7 (-14)	8.2 (-15)	1.3 (-19)
LBV	4	1.3 (-15)	1.3 (-17)	8.1 (-16)	3.3 (-16)	9.7 (-19)
Unknown	25	2.4 (-15)	1.3 (-18)	2.2 (-15)	9.7 (-17)	5.0 (-19)
total/mean	204	2.6 (-13)	1.6 (-17)	4.7 (-14)	1.3 (-15)	5.3 (-19)

Table 2. Statistics of IFOs

Note. — N: Number of IFOs in each type. Flux units are in W m^{-2} . F^{sb}_{mean}: Mean surface brightness of each type (flux divided by area) in W $m^{-2} \ arcsec^{-2}$. Note that one SNR-type (IFO 7) was not used for statistics of fluxes.

A

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candidates
YSO
or
YSO
with
Associated
IFOs
Table 3.

	Table 3. IFOs Associated with YSO or YSO car	ndidates			
Ŝ.	IFO # YSO / YSOc Name N	Aorphology	d [kpc]	Reference YSO Counterpart	Dist
	IBO 003 YSO AGAL G007.333-00.016	k	2.96	ro8/u18	u18
	IFO 004 YSO candidate ALLWISE J180219.38-213351.9	а	1	ro8	ı
	IFO 010 Class I YSO [RBG2009] G009.86-0.04 4	k	$2.36^{+0.78}_{-0.88}$	ro8/r09	r09
	IFO 013 Northern lobe of YSO candidate SSTGLMC G009.7612-00.9575	p	-	ro8	ı
	IFO 014 Southern lobe of YSO candidate SSTGLMC G009.7612-00.9575	p	ı	ro8	ı
	IFO 016 YSO AGAL G013.779+00.609	а	2.90	ro8/u18	u18
	IFO 018 YSO candidate 2MASS J18140816-1850560	a		ro8	ı
	IFO 019 YSO candidate 2MASS J18140816-1850560	а	1	ro8	ı
	IFO 020 W 33, IRS 3 having an O6.5 star	a	$2.40^{+0.17}_{-0.15}$	969	i13
	IFO 021 W 33, IRS 1 having an 06.5 + an 07.5 or 08 star	а	$2.40^{+0.17}_{-0.15}$	b98	i13
	IFO 022 W 33, [MDF2011b] cl1 which encircles O6-7 star #23	a	$2.40^{+0.17}_{-0.15}$	m15	i13
	IFO 023 1. IRAS 18114-1825: Class I. 2. J181421.71-182459.0: Class I/IIc	c	2.41	ro8/yu12/m16	yu12
	IFO 029 YSO AGAL G014.414-00.069	c	3.1	u18	u18
	IFO 030 IRAS 18144-1723, Class I/II binary	k	4.33	c13/v18	v18
	IFO 031 1. YSO IRAS 18151-1208 2. 2. 2MASS J18175094-1208028, Class I/II YSO	а	3.00	m16	m13
	IFO 032 1. 2MASS J18175094-1208028: Class JU, 2. ALLWISE J181749.45-120751.1	а	3.00	m16	m13
	IFO 033 YSO IRAS 18151-1208	а	3.00	v10	m13
	IFO 034 [PW2010] 236, Class 0/I	а	2.10	ro8/c13/p10	p10
	IFO 035 [PW2010] 236, Class 0/I	а	2.10	ro8/c13/p10	p10
	IFO 036 In the middle of multiple YSOs in M 16	с	2.14	c13	66q
	IFO 037 In the middle of multiple YSOs in M 16	k	2.14	c13	66q
	IFO 039 Tip of column 3 of M 16, either T-Tauri star [TSH2002] S-1 or S-2	¥	2.14	t02	66q
	IFO 040 Edge of column 2 of M 16	¥	2.14	t02	66q
	IFO 045 Near the edge of M 16 Pillar V, RMS massive YSO G017.0332+00.7476A	в	2.14	ro8/c13	66q
	IFO 046 Proximity of Class 0/I YSO [PW2010] 378	в	2.10	ro8/p10	p10
	IFO 047 Feature connected to Class 0/I YSO [PW2010] 411	c	2.10	c13/p10	p10
	IFO 048 In the vicinity of YSOs in M 16	а	2.14	g07	66d
	IFO 049 Spatially connected to massive YSO G017.0666+0.6826	U I	2.14	c13	66q
	IFO 051 Compact feature in the crowded region of YSOs in M 17	¥	1.60 ± 0.30	si	n01
	IFO 052 Ditto	k	1.60 ± 0.30	si	n01
	IFO 053 Ditto	Č	1.60 ± 0.30	si	n01
	IFO 056 Multiple YSO candidates in the northern region of M17, EB (extended bubble)	J	1.98	ro8/p09	c16
	IFO 057 Ditto	а	1.98	ro8/p09	c16
	IFO 059 Southern jet of IFO 060	q	1.85 ± 0.2	ro8/c13	x19
	IFO 060 IRAS 18177-1405 aligned with IFO 059, 061, in M 16	p	1.85 ± 0.2	ro8/c13	x19
	IFO 061 Northern jet of IFO 060	p	1.85 ± 0.2	ro8/c13	x19
	IFO 063 Located at the edge of IRDC HEC G016.93+00.24	a a	2.40	si	r10

Table 3 (cont'd)

 IFO 072 Class I/T YSO IRAS 18229-1308 IFO 077 Class I/T YSO IRAS 18229-1308 IFO 077 Class I/YSO Statistic MSO RAGIABLE 1154. In the elected IFO 077 Class I/YSO Statistic MSO RAGIABLE 1151-1154. In the elected IFO 077 Class I/YSO Class I/YSO candidate I/YSO Statistic MSO RAGIABLE 1152/001 IFO 079 Class I/YSO Class I/YSO candidate I/YSO Statistic MSO RAGIABLE 1157/081 40641584. Multiple nearby YSOS IFO 090 Constdy Ray Rue Class I/SO candidate I/YSO Statistic MSO RAGIABLE 1157/081 40641584. Multiple nearby YSOS IFO 090 Constdy Rue I/YSO I/YSE 11872081 406411570 40641584. Multiple nearby YSOS IFO 000 Constdy Rue I/YSO Radia Rue I/YSO RAGIA RUE 100 (A) 100 (IFO 072 Class J/I YSO IRAS 18229-1308 IFO 077 Class I/I YSO IRAS 18239-1304 IFO 079 Massive YSO IRAS 18234-1134. (Fe11) L64 µm detected IFO 070 Class I YSO candidate 1182389:33:115006 IFO 091 Biporheal structure concident with IFR clumps new PN in Froebrich et al. (2015) IFO 091 Class I YSO candidate structures concident with IFR 100. & 101 IFO 092 Concident with UKIDSS source UGFS 81.064158.4. Multiple nearby YSOs IFO 093 Constit Area and the revision is the vicinity IFO 004 Constit Area and the row ALI WSO candidate IFO 100 Sunt-wester injet. HIFO 100. & 101 IFO 101 Sourt-wester injet. HIFO 100. & 101 IFO 101 Sourt-wester injet. Aligned with IFO 102.8.1 IFO 100 Class I YSO candidate variable ASSSN-V 118394.8.07 043015.9 IFO 100 Sourt-wester injet. Aligned with IFO 103.8.350.00.1284, a 2.01 IFO 101 Sourt-wester injet. Aligned with IFO 102.8.1 IFO 101 Spatially coincident with H₂ which is connected to MSXCG G031.8360.00.1284, a 2.01 IFO 113 Spatially coincident with H₂ which is connected to MSXCG G031.8360.00.1284, a 2.02 IFO 113 Spatially coincident with H₂ which is connected to MSXCG G031.8360.00.1284, a 2.02 IFO 113 Spatially coincident with H₂ which is connected to MSXCG G031.8360.00.1284, a 2.03 IFO 123 Coincident with H₄ which is connected to MSXCG G031.9340.00.1284, a 2.03 IFO 124 Matche is protostellar clump at 9.349.40.353 IFO 125 Coincident with H₂ which is connected to MSXCG G031.8360.00.1284, a 2.06 IFO 124 Matche is protostellar clump at 9.349.35500.00.1284, a 2.06 IFO 124 Matche is PASY ISAN. IFY PASY ISAN.00.00.1284, a 2.06 IFO 124 Matche is protostellar clump at 9.349.35500.00.1284, a 2.06 IFO 124 Matche is protostellar clump at 9.349.35500.00.1284, a 2.06 IFO 124 Matche is protostellar clump at 9.340.35500.00.1284 IFO 124 Matche is pro	 1.1.64 μm detected 1.609.6 IR clumps, new PN in Froebrich et al. (2015) stellar clumps in the vicinity PS J183720.81-064158.4. Multiple nearby YS (00 & 101 ported as AGB candidate 9 & 100 16-043113.8 Class III or more evolved YSO (J183948.07-043015.9 candidate GaiaDR2 4258232818679065216 3] G029.8623-0.0437, [Fe II] detection reporte s connected to MSX6C G031.8380-00.1284, fib IFO 123, aligned East to West and/date SSTGLMC G035.2913+00.8076 40.38 I ne vicinity tih IFO 123, aligned East to West and/date SSTGLMC G035.2913+00.8076 40.38 I solverted by p16 6035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518 	 (0) 072 Class I/II YSO IRAS 18229-1308 (0) 074 Massive YSO IRAS 18241-1134. [Fe (0) 077 Class I YSO candidate J182859.53-11 (0) 091 Biconical structure coincident with FI (0) 092 Located in IRDC 24.764-0.12. Proto- (0) 093 North-western jet, aligned with IFO 19 (0) 100 Class I YSO candidate, previously rep (0) 101 South-eastern jet, aligned with IFO 99 (0) 101 Beroximity of pre-main sequence star contant semi-regular variable ASASSN-V (0) 111 Massive protostellar object (VEN201. (0) 112 Spatially coincident with H₂, which is YSO candidate SSTGLMC G031.836 (0) 123 Coincident with flat-spectrum YSO candidate SSTGLMC G031.836 (0) 123 Coincident with flat-spectrum YSO candidate SSTGLMC G031.836 (0) 123 Coincident with flat-spectrum YSO candidate SSTGLMC G031.836 (0) 124 Matches to proto-stellar clump 34.934 (0) 125 Cometary structure coincident with more evolved YSO ALLW
10007 Class I YSO candidate I R32835.3.1 1500.06 cl 22.00 cl 10007 Class I YSO candidate I R32835.3.1 1500.05 a 12.60 cl 17008 Diopheal structure considiate with FR temps, new PN in Froebricht et al. (2015) b 3.37 si 17008 Constal in RNUC andidate. Frontossellar clumps, in the vicinity b 3.35 m68/21 17008 Fordidate myth FP (00 k 10) constal in RNO candidate. Fordidate proviously reported a ACIB standidate b 3.30 m68/21 17000 South-eastin jet, algoed with FP (00 k 10) constal in RNO candidate. Fordidate proviously reported a ACIB standidate b 3.30 m68/21 170010 South-eastin jet, algoed with FP (00 k 10) cold 13850-00.1384, a 2.00 w58/21 170010 Route surface variable ACIB SSIN V1 R3880.70-43015.50 a 2.00 w58/21 m64/21 170010 Route surface surface surface variable ACID R380.70-43015.50 a 2.00 w58/21 170010 Route surface surf	 UP0 077 Massive YSO IRAS 18241-1134. [Fe.it] 1.64 µm detected FO 077 Class 1 NSO candidate 1120006 FO 077 Class 1 NSO candidate 1182389.53-1150005 FO 091 Storbard structure conclored my FIR clumps in the vicinity FO 091 Storbard structure conclored my FIR clumps in the vicinity FO 093 Correct Jin IRDC 24.764-012. Proto-stellar clumps in the vicinity FO 093 Correct Jin IRDC 24.764-012. Proto-stellar clumps in the vicinity FO 093 Correct Jin IRDC 24.764-012. Proto-stellar clumps in the vicinity FO 093 Nuch vestor right with IFD 100. & 101 FO 100 Class 1 NSO candidate, previously reported as AGB candidate FO 101 South-stepsine variable AXSN-V 1183948.07-043015.9 FO 103 South-stepsine variable AXSN-V 1183948.07-043015.9 FO 104 Inthe middle of ALLHMSE 118320.2311.604311.56 FO 104 Inthe middle of ALLHMSE 118320.2311.604311.56 FO 105 Provinity of the randidate clamp 24.94.06 FO 111 Massive protostellar clump 34.94.07-043015.9 FO 105 Provinity of the randidate clamp 24.33.45 FO 115 Spatially coincident with He visitiy FO 112 Stormetary structure compact structure with IFO 12.3. aligned East to West FO 123 Storas a stimilar compact structure with FO 123. aligned East to West FO 124 Matches to proto-stellar clump 34.94.34.3 FO 125 Context y structure concident with IFO 124.3. aligned East to West FO 125 Context y structure concident with IFO 125.2011.603.235.291340.0765 FO 125 Context y structure concident with FO 125.20134.03.06.01253. FO 125 Context y structure concident with FO 124.3. and the sterst vicinity FO 125 Context y structure concident with FO 125.2014.003.01.03.1.03.4. FO 126 Context y structure concident with FO 125.2014.01.03.1.04.4. FO 127 Structure the number of the structure the structure transite transite structure transite transite structure transite tra	rr] 1.64 μ m detected 15009.6 IR clumps, new PN in Froebrich et al. (2015) stellar clumps in the vicinity PS J183720.81-064158.4. Multiple nearby YS 00 & 101 ported as AGB candidate 9 & 100 16-043113.8 Class III or more evolved YSO $\sqrt{183948.07-043015.9}$ candidate GaiaDR2 4258232818679065216 3] G029.8623-0.0437, [Fe II] detection reporte sconnected to MSX6C G031.8380-00.1284, 5]-00.1408 in the vicinity ith IFO 123, aligned East to West andidate SSTGLMC G035.2913+00.8076 $\pm 0.338.4$ is the elosest, VISE J185522.49+030130.3 in 30" distance ± 0.56 , [Fe II] veported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518	 0 074 Massive YSO IRAS 18241-1134. [Fe 0 077 Class I YSO candidate J182859.53-11 0 091 Biconical structure coincident with FI 0 092 Located in IRDC 24.764-0.12. Proto- 0 099 North-western jet, aligned with IFO 1 0 100 Class I YSO candidate, previously rep 0 101 In the middle of ALLWISE J183951.1 and semi-regular variable ASASSN-v 0 108 Proximity of pre-main sequence star c 0 101 Massive protostellar object [VEN201. 0 115 Spatially coincident with H₂. which is YSO candidate SSTGLMC G031.836 0 123 Coincident with flat-spectrum YSO c 0 123 Coincident with flat-spectrum YSO c 0 124 Matches to proto-stellar clump 34.934. 0 125 Cometary structure coincident with m 0 124 Matches to proto-stellar clump 34.934.
Fig007 Totals: 1/SSO standides (1828), 55:115006 a c) k21 Fi0091 Big/mest transmer coincident with FR clumps, new PN in Frechrich et al. (2015) b 3.4 c) Fi0092 Constraint with FR clumps, new PN in Frechrich et al. (2015) b 3.50 res/21 Fi0092 Constraint with FFD 00.8 (10) Fi0003 Constraint with FFD 00.8 (10) b 3.50 res/21 Fi00103 Constraint with FFD 10.8 (10) B 3.50 res/21 3.50 res/21 Fi00103 Constraint with FFD 10.8 (10) B 3.50 res/21 3.50 res/21 Fi01013 Constraint with FFD 10.8 (10) B 3.50 res/21 8.60 3.50 res/21 8.60 Fi01013 Finally constraint with FFD 1138 Class III on more evolved YSO B 3.50 res/21 8.60 9.50 1.50	 FOOT Class I YSO candidate J 182859.53-115009.6 FOO Bi Ecoheral arteures coincident with FRO lumps, new PN in Froebrich et al. (2015) FOO 092 Constellar Source UCFPS J 183720.81-064158.4. Multiple nearby YSOs FOO 092 Constellar Source UCFPS J 183720.81-064158.4. Multiple nearby YSOs FOO 093 Constellar Charge with IFO 100 & 101 FOO 1003 Sourth search pict, aligned with IFO 100 & 101 FOO 1003 Sourth search pict, aligned with IFO 100 & 101 FOO 1015 Source UCFPS J 183721.16-043113.8 Class III or more evolved YSOs FOO 103 Sourth-search pict, aligned with IFO 100 & 101 FOO 104 In the middler of ALLWNES J 183951.16-043013.8 S00.0037516 FOO 104 In the middler of ALLWNES J 183951.16-043013.8 Class III or more evolved YSO FOO 104 In the middler of ALLWNES J 183951.16-04371.7 Fer il detection reported FOO 104 In the middler SYTGLIC CO01 g and the vicinity FOO 105 Sourth-search pict aligned with IFO 123, aligned East to West FOO 115 Spatialy coincident with H₂ which is connected to MSX CG G031.8380-00.1284, a FOO 115 Spatialy coincident with H₂ which is connected to MSX CG G031.8380-00.1284, a FOO 125 Shares a similar compact structure with FO 123, aligned East to West FOO 125 Shares a similar compact structure with FO 123, aligned East to West FOO 125 Shares a similar compact structure with H₂ MS 5005.0 Fe 10.10805.0 He 10.2 MSX 18.5 MS 5005.0 Fe 10.2 MSX 18.5 MS 5005.0 He 10.2 MS 18.5 MS 5005.0 He 10.2 MS	 15009.6 IR clumps, new PN in Froebrich et al. (2015) stellar clumps in the vicinity PS J183720.81-064158.4. Multiple nearby YS 100 & 101 ported as AGB candidate ported as AGB candidate 9 & 100 15.043113.8 Class III or more evolved YSO / J183948.07-043015.9 candidate GaiaDR2 4258232818679065216 3] G029.8623-0.0437, [Fe II] detection reporte sconnected to MSX6C G031.8380-00.1284, 51-00.1408 in the vicinity fib IFO 123, aligned East to West andidate SSTGLMC G035.2913+00.8076 +0.338 J aassive YSO IRAS 18527+0301 aassive YSO IRAS 18527+0301 aassive YSO IRAS 19045+0518 current of the stoch of the s	 0 077 Class I YSO candidate J182859.53-11 0 091 Biconical structure coincident with FI 0 092 Located in IRDC 24.764-0.12. Proto- 0 094 Coincident with UKIDSS source UGH 0 099 North-western jet, aligned with IFO 1 0 100 Class I YSO candidate, previously rep 0 101 South-eastern jet, aligned with IFO 95 0 104 In the middle of ALLWISE J183951.1 and semi-regular variable ASASSN-V 0 108 Proximity of pre-main sequence star c 0 101 Massive protostellar object (VEN201: 0 115 Spatially coincident with H2, which is YSO candidate SSTGLMC G031.836 0 123 Coincident with flat-spectrum YSO c 0 124 Matches to proto-stellar clump 34.934 0 125 Cometary structure coincident with more evolved YSO ALLW
$ \begin{array}{c} \label{eq:constraint} (2015) \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \$	 IFO 091 Biophical structure coincident with FIR clumps, new PN in Frechrich et al. (2015) 160 092 Coeach in RIDC 24, 764-01.2. Proto-stellar clumps in the vicinity 160 003 Confident with UKDC 24, 764-01.2. Proto-stellar clumps in the vicinity 160 001 Const 1YSO candidate, previously reported as AGB candidate 160 101 South-asstem jet, aligned with IFO 100 & 101 160 101 South-asstem jet, aligned with IFO 98 1183720811-004158.4. Multiple nearby YSOS 160 101 South-asstem jet, aligned with IFO 99 & 100 160 101 South-asstem jet, aligned with IFO 99 & 100 160 101 South-asstem jet, aligned with IFO 99 & 100 160 101 South-asstem jet, aligned with IFO 99 & 100 160 101 South-asstem jet, aligned with IFO 99 & 100 160 101 Massive protostellar object (VBR2013) GO39 S070-001384, 170 115 Spatialty coincident with I₄, which is connected to MSXcG G031 S380-00.1284, 170 115 Spatialty coincident with I₄, which is connected to MSXcG G031 S380-00.1284, 170 115 Shares a similar compact structure with IFO 123, aligned East to West 170 123 Coincident with Iar-spaces. 170 123 Coincident with Iar-spaces. 170 123 Coincident with massive SYSO RAS 1803-5013-603 S035-03451 S6 170 123 Coincident with massive SYSO RAS 1803-5013 03.3 170 124 Coincident with massive SYSO RAS 1803-5013 0.3 170 124 Coincident with massive SYSO RAS 1803-5090-5038 in West 170 125 Coincident with transpectum YSO SCIICLAC G032-5060-90538 in West 170 126 Coincident with transpectum SSSO 2003-5060-90538 in West 170 128 Coincident with transpectum YSO RAS 19045-40518 in West 170 128 Coincident with transpectum SSSO RAS 1903-5090-888 in West 170 128 Coincident with transpect RAS 18525-40903780-90358 in 9045 40518 in West 170 128 Coincident with transpect RAS 18555-4003780 rol 438 in West 170 128 Coincident with transpect	IR clumps, new PN in Froebrich et al. (2015) stellar clumps in the vicinity PS J183720.81-064158.4. Multiple nearby YS (00 & 101 ported as AGB candidate 9 & 100 16-043113.8 Class III or more evolved YSO v J183948.07-043015.9 candidate GaiaDR2 4258232818679065216 3] G029.8623-0.0437, [Fe II] detection reporte s connected to MSX6C G031.8380-00.1284, 5] 00.1408 in the vicinity ith IFO 123, aligned East to West andidate SSTGLMC G035.2913+00.8076 +0.3384 is the vicinity ith IFO 123, aligned East to West andidate SSTGLMC G035.2913+00.8076 +0.3384 is the closest, VISE J185522.49+030130.3 in 30" distance +0056, [Fe II] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518	 (0) 091 Biconical structure coincident with FI (0) 092 Located in IRDC 24.764-0.12. Proto- (0) 094 Coincident with UKIDSS source UGF (0) 099 North-western jet, aligned with IFO 1 (0) 100 Class I YSO candidate, previously rep (0) 101 South-eastern jet, aligned with IFO 99 (0) 101 South-eastern jet, aligned with IFO 99 (0) 101 In the middle of ALLWISE J183951.1 and semi-regular variable ASASSN-V (0) 108 Proximity of pre-main sequence star coincident with H₂, which is YSO candidate SSTGLMC G031.836 (0) 115 Spatially coincident with H₂, which is YSO candidate SSTGLMC G031.836 (0) 123 Coincident with flat-spectrum YSO candidate binary Gaia2 42807567266 (0) 125 Cometary structure coincident with more evolved YSO ALLW
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Fer ul detection reported FO 1118 Masive protostent with Asykin Score concectal to MSX6G G031.8.380-00.1284, west FG 0125 Spatially coincident with Rat-secturm YSO standidate SSTGLMC G035.2913+00.8076 FG 0123 Coincident with Rat-secture with FG 123, aligned East to West FG 123 Solve option-selar option structure with Rat-secture with FG 123, aligned East to West FG 123 Coincident with Rat-secturm YSO standidate SSTGLMC G035.2913+00.8076 FG 123 Coincident with Rat-secture with Rat-secture YSO RAS 18577+0301 FG 123 Coincident with Rat-secture YSO RAS 18577+0301 FG 124 Matchs to proto-selar coincident with massive YSO RAS 18577+0301 FG 125 Context structure coincident with massive YSO RAS 18577+0301 FG 125 Context structure coincident with massive YSO RAS 19045-0515 FG 124 Matchs to proto-selar coincident with massive YSO RAS 19045-0513 FG 124 Matchs to proto-selar coincident with massive YSO RAS 19045-0513 FG 125 Context structure with massive YSO RAS 19045-0513 FG 126 Context structure with the sector optic coincident with massive YSO RAS 19045-0515 <l< td=""><td>stellar clumps in the vicinity PS J183720.81-064158.4. 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$ \begin{tabular}{ c c c c c c c c c c c c c c c c c c c$	 IFO 104 In the middle of ALLWNSE J183951.16-043113.8 Class III or more evolved YSO and semi-regular warfable ASASSN v J183948.07-043015.9 IFO 108 Proximity of pre-marfable ASASSN v J183948.07-043015.9 IFO 108 Proximity of pre-marfable ASASSN v J183948.07-043015.9 IFO 110 Massive protostellar object/VEBX013] G029.862.3.0437. [Fe 1] detection reported c YSO candidate SSTGLMC G031, 3561.00.1408 in the vicinity IFO 112 Shares a similar comports structure with FO 123, aligned East to West IFO 122 Shares a similar comports structure with FO 123, aligned East to West IFO 123 Coincident with flat-spectrum YSO candidate SSTGLMC G035.2913400.8076 IFO 124 Matches to proto-stellar clung 34.9.3.0384 is the elosest, IFO 125 Connetary structure concident with massive YSO IRAS 18527.40301 IFO 126 Ultrawide binary Gaia2 428075675668653384 is the elosest, IFO 126 Ultrawide binary Gaia2 428075675668653384 is the elosest, IFO 126 Ultrawide binary Gaia2 42807567566661330.30.3 in 30° distance IFO 126 Ultrawide binary Gaia2 428075675666613 0.30.3 in 30° distance IFO 126 Ultrawide binary Gaia2 428075675668695388 to 905328, proto-stellar clump is coincident a 2.48 IFO 126 Ultrawide binary Gais2 2199400.8638. G039.2106400.8818 in West IFO 131 Class III massive YSO RAS 1904549518 IFO 134 Aligned with YSO AGLL G039.7189-0103.838. f004549518 IFO 134 Aligned with YSO AGLL G039.708-01.237 and IFO 144 IFO 144 Aligned with YSO AGLL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.2109-106.5851-0.3814 IFO 145 Longate component matches to massive YSO IRAS 1904549518 IFO 145 Longated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 IFO 145 Longated and pointing toward YSO AGAL G039.2090-00.5851-0.3814 IFO 145 Longated and pointing toward YSO AGAL G039.708-01.237 and IFO 145	 16-043113.8 Class III or more evolved YSO V J183948.07-043015.9 candidate GaiaDR2 4258232818679065216 3] G029.8623-0.0437, [Fe II] detection reporte s connected to MSX6C G031.8380-00.1284, 51-00.1408 in the vicinity 51-00.1408 in the vicinity th FO 123, aligned East to West and/date SSTGLMC G035.2913+00.8076 40.338 J and/date SSTGLMC G035.2913+00.8076 40.338 J and/date SSTGLMC G035.2913+00.8076 40.338 J and/date SSTGLMC G035.2913+00.8076 40.358 J 6035.2984 is the closest, VISE J185522.49+030130.3 in 30" distance e0056, [Fe II] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518 	 (0) 104 In the middle of ALLWISE J183951.1 and semi-regular variable ASASSN-V (0) 108 Proximity of pre-main sequence star c (0) 111 Massive protostellar object (VEN201. (0) 115 Spatially coincident with H₂, which is YSO candidate SSTGLMC G031.836 (0) 125 Shares a similar compact structure with flat-spectrum YSO c (0) 123 Coincident with flat-spectrum YSO c (0) 124 Matches to proto-stellar clump 34.934 (0) 125 Cometary structure coincident with m (0) 125 Cometary structure coincident with m (0) 126 Ultrawide binary Gaia2 42807567266 (1) 128 M of the secoled of the second structure with m
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FC 2049.9010-00.5922, Sectal index of flat (FS) according to k21 FO 175 Roundered to Class 1 YSO SSTDERC G049.9010-00.5922, spectral index of flat (FS) according to k21 FO 174 Emerges from Class 1 YSO RML2017] MC2 M105 FO 174 Emerges from Class 1 YSO RML2017] MC2 M105 FO 174 Emerges from Class 1 YSO RML2017] MC2 M105	 V J183948.07-043015.9 candidate GaiaDR2 4258232818679065216 3] G029.8623-0.0437, [Fe II] detection reporte s connected to MSX6C G031.8380-00.1284, 51-00.1408 in the vicinity j1-00.1408 in the vicinity ith IFO 123, aligned East to West and/date SSTGLMC G035.2913+00.8076 +0.338 J and/date SSTGLMC G035.2913+00.8076 +0.338 J and/date SSTGLMC G035.2913+00.8076 40.338 J and/date SSTGLMC G035.2913+00.8076 40.358 J and/date SSTGLMC G035.2913+00.8076 40.358 J 40.358 J 40.352 J 49+030130.3 in 30" distance 40056, [Fe II] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. 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FO 108 Proximity of pre-mair sequence star candidate GaiaDR2 4258232818679065216a2.01v.20FO 111 Massive provoselian object //PSQ 031 6029 8653-00437, [Fe 1] detection reportedc6.21c13/a20FO 115 Sparially concident with H ₂ withen is compact structure with H ₂ with FO 123 Concident with flat-spectrum YSO standidate STGLMC G031.8861.001.1844,a2.01v.20FO 123 Concident with flat-spectrum YSO standidate STGLMC G035.2913+00.8076k4.80ros8k21FO 123 Concident with massive YSO flat S157-0103n30" distancea2.90r15FO 124 Matches to proco-stellar clump 34 93-03.53from S2a2.90r15FO 125 Connetary structure coincident with massive YSO flat S1852.294403103a1.00r13/p16FO 126 Current with massive YSO RAS 18552.294503103a2.30r13/p16FO 124 Matches to proco-stellar clump 34 93-03.53a2.00r15FO 125 Connetary structure coincident with massive YSO RAS 190450401881 in Westa2.470m96FO 134 Close to YSO candidate SSTGLMC G035.2868-00.9578, proto-stellar clump is coincidenta2.48m66FO 134 Close to YSO candidate SSTGLMC G039.23600400.8818 in Westa2.470m16FO 134 Close to YSO candidate SSTGLMC G039.2366-0.9381 19.06from starterc1.10c13/p16FO 134 Close to YSO starterfrom starterfrom starterc1.10c13/p16FO 134 Close to Y	 IFO 108 Proximity of pre-mair sequence star candidate GaiaDR2 4258232818679065216 a 2.01 FO 111 Massive protostellar objest/VEN20131 6029 8623-0.0437, [Fe 11] detection reported c 6.21 FO 115 Spatially coincident with P₄ which is connected to MSXGC G031.8380-00.1284, a YSO candidate SSTGLMC G031, 8361/00.1408 in the vicinity FO 122 Shares a similar compact structure with FO 123, aligned East to West FO 123 Concident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 FO 123 Connectator structure concident with massive YSO IRAS 1852-40301 FO 123 Connectator structure concident with massive YSO IRAS 18522.4030130.3 FO 123 Connectator structure concident with massive YSO IRAS 18522.4030130.3 FO 124 Matches to proto-stellar clump 34.93-4.3387 FO 125 Connectator structure concident with massive YSO IRAS 18522.4030130.3 FO 124 Matches to proto-stellar clump 34.93-4.03.5381 FO 125 Connectator structure concident with massive YSO IRAS 18522.4030130.3 FO 134 Close to YSO candidate SSTGLMC G035.2868-00.9528, proto-stellar clump is coincident FO 134 Close to YSO candidate SSTGLMC G035.2868-00.9528, proto-stellar clump is coincident FO 134 Close to YSO candidate SSTGLMC G035.2868-00.9528, proto-stellar clump is coincident FO 134 Close to YSO candidate SSTGLMC G039.2199-00.8818 in West FO 134 Close to YSO candidate SSTGLMC G039.2199-00.8518 in West FO 134 Close to HI 186, 42 arcsec away from [FO 144, 145, H₂ knots FO 134 Close to HI 186, 42 arcsec away from [FD 120101 L673 10 and 13 FO 135 Close to HI 180, 42 arcsec away from [TBP20101 L673 10 and 13 FO 135 Close to HI 180, 42 arcsec away from [TBP20101 L673 10 and 13 FO 135 Close to HI 180, 42 arcsec away from [TBP20101 L673 10 and 13 FO 135 Close to HI 180, 42 arcsec away from [TBP20101 L673 10 and 13 FO 135 Close to HI 18	candidate GaiaDR2 4258232818679065216 3] G029.8623-0.0437, [Fe II] detection reporte s connected to MSX6C G031.8380-00.1284, 51-00.1408 in the vicinity ith JFO 123, aligned East to West andidate SSTGLMC G035.2913+00.8076 +0.338 1 ansive YSO IRAS 18527+0301 assive YSO IRAS 18527+0301 assive YSO IRAS 18527+0301 566953984 is the closest, VISE J185522,49+030130.3 in 30" distance +0056, [Fe II] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518	 (0) 108 Proximity of pre-mair sequence star c (0) 111 Massive protostellar object [VEN201] (0) 115 Spatially coincident with H₂, which is YSO candidate SSTGLMC G031,836 (0) 122 Shares a similar compact structure wii (0) 122 Shares a similar compact structure wii (0) 123 Coincident with flat-spectrum YSO ca (0) 124 Matches to proto-stellar clump 34.934 (0) 125 Cometary structure coincident with m (0) 125 Cometary structure coincident with m (0) 126 Ultrawide binary Gaia2 42807567266 (1) 128 III or more evolved YSO ALLW
FO 111 Massive protostellar object/VEN2013J G029.8623-0.0437, [Fe II] detection reported c 6.21 c13A20 FO 115 Spatialy concident with H ₂ , which is connected to MSX6G G031.8380-00.1284, a - ro8603 FO 125 Shares a similar compact structure with H ₂ , which is connected to MSX6G G031.8380-00.1284, a - ro86X21 FO 123 Coincident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 k 4.80 ro86X21 FO 123 Coincident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 k 4.80 ro86X21 FO 123 Coincident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 k 4.80 ro86X21 FO 123 Concident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 k 4.80 ro86X21 FO 124 Contravide binary Gain2 428075675686939346 pite bioaxt, a 2.90 t15 FO 134 Class VII massive YSO IRAS 18552-40430130.3 in 30° distance c 1.20/m16 FO 134 Class VII massive YSO RASI 18552.294-9430130.3 in 30° distance c 1.20/m16 FO 134 Class VII massive YSO RASI 18552-406-90.888 in West a 2.48 ro86/m18 FO 135 Class VII massive YSO RASI 18552.406-90.888 in West a 2.48 <	 IFO 111 Massive protostellar object (VEN2013] G029.8623-0.0437, [Fe II] detection reported IFO 115 Spatially coincident with H₂, which is connected to MSX6C G031.8380-00.1284, YSO candidate SSTGLMC G031.861 00.1408 in the vicinity IFO 122 Shares a similar compact structure with JFO 123, aligned East to West IFO 123 Coincident with flat-spectrum YSO Eandidate SSTGLMC G035.2913+00.8076 IFO 124 Matches to proto-stellar clump 34.934.0.338 J IFO 125 Connetary structure coincident with massive YSO RAS 18527.40301 IFO 125 Connetary structure coincident with massive YSO RAS 18527.40301 IFO 125 Connetary structure coincident with massive YSO RAS 18527.4030130.3 in 30" distance IFO 136 Class III on more evolved YSO ALLWISE J18552.2.494030130.3 in 30" distance IFO 131 Class III massive YSO RAS 18552.4056, IFE u) reported by p16 IFO 134 Close to YSO candidate SSTGLMC G035.2816.40538.6039.2060.400.8818 in West IFO 134 Close to YSO candidate SSTGLMC G039.2199+00.8638. G039.2060.400.8818 in West IFO 134 Close to YSO candidate STGLMC G039.708.01.237 and IFO 144. IFO 144 Compact component matches to massive YSO IRAS 19045.40513 IFO 144 Aligned with YSO AGAL G039.708.01.237 and IFO 144. IFO 145 Elongated and pointing toward YSO AGAL G039.708.01.237 and IFO 144. IFO 146 To PN candidate, yet aligned with FD 144, 145, H₂ hous IFO 147 Coincident with class 1 and flat.SED YSOs, [TBP2010] L673 10 and 13 IFO 146 IIP NC candidate, yet aligned with FD 144, 145, H₂ hous IFO 156 Coincident with class 1 and flat.SED YSOs, [TBP2010] L673 10 and 13 IFO 156 Coincident with class 1 and flat.SED YSOs, [TBP2010] L673 10 and 13 IFO 156 Close to HH 186, 42 acces eavay from [TBP2010] L673 10 and 13 IFO 156 Close to HH 186, 42 acces eavay from [TBP2010] L673 10 and 13 IFO 156 Close to HH 186, 42 acces eavay from [TBP2010] L	3] G029.8623-0.0437, [Fe II] detection reporte s connected to MSX6C G031.8380-00.1284, 51-00.1408 in the vicinity ith IFO 123, aligned East to West andidate SSTGLMC G035.2913+00.8076 +0.338 1 ansive YSO IRAS 18527+0301 assive YSO IRAS 18527+0301 assive YSO IRAS 18527+0301 586953984 is the closest, VISE J185522-49+030130.3 in 30" distance +0056, [Fe II] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518	 (0 111 Massive protostellar object [VEN201: (0 115 Spatially coincident with H₂, which is YSO candidate SSTGLMC G031,836 (0 122 Shares a similar compact structure with fiat-spectrum YSO ca 0 123 Coincident with fiat-spectrum 34.934 (0 124 Matches to proto-stellar clump 34.934 (0 125 Cometary structure coincident with m (0 126 Ultrawide binary Gaia2 42807567266 (0 126 MI or more evolved YSO ALLW
$ \begin{array}{rcl} \mbox{FO} 115 \ Spatially coincident with H_3, which is connected to MSX6C G031 8380-00.1284, a - 008603 \\ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ $	IFO 115 Spatially coincident with H ₂ which is connected to MSX6C G031.8380-00.1284, YSO candidate SSTGLMC G031.8661*00.1408 in the vicinity TFO 122 bares a similar compact structure with IFO 123 aligned East to West TFO 122 bares a similar compact structure with IFO 123 aligned East to West TFO 123 Conncident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 k 4.80 FFO 123 Conncident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 k 4.80 FFO 123 Conncident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 k 4.80 FFO 124 Matches to proto-stellar clump 34.93+0.318.318.522.49+0.30130.5 in 30° distance = 2.470 FFO 124 Class III or more evolved YSO ALLWISE J18552.2.49+0.30130.3 in 30° distance = 2.48 FO 130 Class III or more evolved YSO ALLWISE J18552.2.49+0.30130.3 in 30° distance = 2.48 FO 134 Class to YSO SSTGLMC G039.2199+0.08.638 G039.2060+0.0818 in West = 2.48 FO 134 Class to YSO SSTGLMC G039.2199+0.05.635 G039.2060+0.0818 in West = 2.48 FO 136 Flat spectrum YSO SSTGLMC G039.2199+0.05.635 G039.2060+0.0818 in West = 2.48 FO 136 Flat spectrum YSO SSTGLMC G039.2199+0.05.635 G039.2060+0.0818 in West = 2.48 FO 136 Flat spectrum YSO SSTGLMC G039.2199+0.05.635 G039.2060+0.0818 in West = 2.48 FO 136 Flat spectrum YSO SSTGLMC G039.2108-01.237 and IFO 144 Highed with YSO AGAL G039.708-01.237 and IFO 144 Highed with YSO AGAL G039.708-01.237 and IFO 145 Flot gated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 Flot gated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 Flot gated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 Flot gated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 Flot gated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 Flot gated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 Flot gated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 Flot gated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 Flot gated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 Flot gated and pointing toward YSO AGA	s connected to MSX6C G031.8380-00.1284, 51-00.1408 in the vicinity ith IFO 123, aligned East to West andidate SSTGLMC G035.2913+00.8076 +0.338 J anssive YSO IRAS 18527+0301 anssive YSO IRAS 19045+0518 in West sive YSO IRAS 19045+0518	 (0) 115 Spatially coincident with H₂, which is YSO candidate SSTGLMC 0031,836 (0) 122 Shares a similar compact structure wi (0) 123 Coincident with flat-spectrum YSO ca (0) 124 Matches to proto-stellar clump 34.934 (0) 125 Cometary structure coincident with mi (0) 125 Cometary structure coincident with mi (0) 126 Ultrawide binary Gaia2 42807567266 (0) 120 Class III or more evolved YSO ALLW
$ \begin{array}{rcl} YSO candidate SSTGLMC G031,3561-00.1408 in the vicinity \\ FO 122 Shares a similar compact structure with FFO 123, aligned East to West \\ FO 123 Coincident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 k k 4.80 ro88/21 \\ FO 123 Coincident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 k k 4.80 ro88/21 \\ FO 124 Matches to proto-stellar clump 34.93+0.389 \\ FO 125 Uhrawide binary Gaia2 42807567.7660 Fish be plosest, \\ C 0.121 Matches to proto-stellar clump 34.93+0.389 \\ FO 125 Uhrawide binary Gaia2 42807567.7660 Fish be plosest, \\ C 0.121 Class III massive YSO IRAS I8555+0056, Fish be plosest, \\ C 0.131 Class III massive YSO RAS I8555+0056, Fish in Foot 120 distance \\ FO 134 Close to YSO candidate SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West \\ FO 134 Close to YSO candidate SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West \\ FO 134 Close to YSO candidate SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West \\ FO 134 Close to YSO candidate SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West \\ FO 134 Close to FSO candidate SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West \\ FO 135 FIO 135 FI$	YSO candidate SSTGLMC G031, 3561-00. 1408 in the vicinity FPO 123 Coincident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 k 4.80 FPO 123 Coincident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 k 4.80 FPO 123 Connetary structure coincident with masker YSO FRAS 18527+0301 FPO 125 Connetary structure coincident with masker YSO FRAS 18527+0301 FPO 126 Ultrawide binary Gaia2 4280756726686953984 is the closest, Class III or more evolved YSO ALLWISE J18552.249+030130.3 in 30° distance FFO 126 Ultrawide binary Gaia2 4280756726686953984 is the closest, FFO 131 Class III on more evolved YSO ALLWISE J18552.249+030130.3 in 30° distance FFO 131 Class III massive YSO FRAS 18555+0056, [Fe n] reported by p16 FFO 131 Class III massive YSO RAAS 18555+0056, [Fe n] reported by p16 FFO 131 Class III massive YSO RAAS 18555+0056, [Fe n] reported by p16 FFO 131 Class III massive YSO RAAS 18555+0056, [Fe n] reported by p16 FFO 134 Close to YSO candidate SSTGLMC G039.219945+0518 FFO 144 Compact component matches to massive YSO RAS 19045+0518 FFO 144 Aligned with YSO AGAL G039.708-01.237 and FFO 144 FFO 144 Aligned with YSO AGAL G039.708-01.237 and FFO 144 FFO 144 Aligned with YSO AGAL G039.708-01.237 and FFO 144 FFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and FFO 144 FFO 146 FPN candidate, yet aligned with FFO 145, hous FFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and FFO 144 FFO 154 Class III/photosphere YSO SSTOERC G049.1657, no for foreground, see A09, k FFO 154 Class III/photosphere YSO SSTOERC G049.9160-00.5922, spectral index of flat (FS) according to k21 FFO 175 A souther compact to Class I YSO IRML2017] MC2 M105 FFO 176 A souther compact to Class I YSO IRML2017] MC2 M105 FFO 176 A souther compact to Class I YSO IRML2017] MC2 M105 FFO 176 A souther compact index of flat (FS) according to k21 FFO 176 A souther compact index of flat (FS) according to k21 FFO 176 A souther compact index of flat (FS) according to k21 FFO 176 A souther compact index of flat (FS) accordin	51-00.1408 in the vicinity ith IFO 123, aligned East to West andidate SSTGLMC G035.2913+00.8076 +0.338 J anssive YSO IRAS 18527+0301 anssive YSO IRAS 18527+0301 anssive YSO IRAS 18527+0301 anssive YSO IRAS 18522, 49+030130.3 in 30° distance 4056, [Fe II] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518	YSO candidate SSTGLMC G031.836 (0 122 Shares a similar compact structure wi (0 123 Coincident with flat-spectrum YSO ca (0 124 Matches to proto-stellar clump 34.93+ (0 125 Cometary structure coincident with m (0 125 Cometary structure coincident with m (0 126 Ultrawide binary Gaia2 42807567266 Class III or more evolved YSO ALLM
$ \begin{array}{rcrcr} \mbox{FO} 122 \ Shares a similar compact structure with IFO 123 \ Sindler extra with IFO 123 \ Coincident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 k k 4.80 \ ro88/221 \ FO 124 \ Matches to proto-stellar clump 34.93.81 \ FO 124 \ Matches to proto-stellar clump 34.93.81 \ FO 125 \ Constary structure coincident with massive YSO RAS 18527+0301 \ coincident with massive YSO RAS 18527+0301 \ coincident with massive YSO RAS 18527+0301 \ coincident with massive YSO RAS 18527-403013 \ mbox{mode} \ FO 125 \ Constary structure coincident with massive YSO RAS 18522.49+03013.03 \ mbox{mode} \ rot 126 \ Ultrawide binary Gaita.248075675666539394 \ mbox{mode} \ rot 136 \ FO 131 \ Class III massive YSO RAS 18552-49-03013.03 \ mbox{mode} \ rot 126 \ Ultrawide binary Gaita.248075675666539394 \ mbox{mode} \ rot 138555-40056, \ Fe m rot 18552.49-03013.03 \ mbox{mode} \ rot 126 \ Ultrawide binary Gaita.2480756756665394 \ mbox{mode} \ rot 138555-40056, \ Fe m rot 18552.49-030130.31 \ mbox{mode} \ rot 138555-40056, \ Fe m rot 18552.49-030130.31 \ mbox{mode} \ rot 126 \ mbox{mod} \ rot 126 \ mbox$	 IFO 122 Shares a similar compact structure with JFO 123, aligned East to West IFO 123 Coincident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 K 4.80 IFO 124 Matches to proto-stellar clump 34.9340.338 J IFO 125 Cometary structure coincident with massive YSO IRAS 18527+0301 IFO 125 Connetary structure coincident with massive YSO RAS 18527+0301 IFO 125 Connetary structure coincident with massive YSO RAS 18527+0301 IFO 126 Ultrawide binary Gaia2 4280755766869539841 is the closest, IFO 126 Ultrawide binary Gaia2 42807556. [Fe1] 185522.294-030130.3 in 30" distance IFO 131 Class II/I massive YSO RAS 18555-4050130.3 in 30" distance IFO 131 Class II/I massive YSO RAS 18555-40516. [Fe1] 185522.2019-00.8638. G039.2060+00.8818 in West IFO 134 Close to YSO candidate SSTGLMC G035.2868-00.9528, proto stellar clump is coincident a 2.48 IFO 134 Close to YSO candidate SSTGLMC G035.2199+00.8638. G039.2060+00.8818 in West IFO 134 Close to YSO AGAL G039.708-01.237 and IFO 144. IFO 141 Compact component matches to massive YSO IRAS 19045+0518 IFO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 144. IFO 145 Elongated and pointing toward YSO AcJU.2037 and IFO 144. IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144. IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144. IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 145. IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144. IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 145. IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144. IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 145. IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 145. IFO 145 Elongated and	 ith IFO 123, aligned East to West andidate SSTGLMC G035.2913+00.8076 +0.338 J +0.338 J aassive YSO IRAS 18527+0301 aassive YSO IRAS 18527+0301 686953984 is the closest, VISE J185522 49+030130.3 in 30" distance +0056, [Fe II] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518 	 (0) 122 Shares a similar compact structure wi (0) 123 Coincident with flat-spectrum YSO cs (0) 124 Matches to proto-stellar clump 34.934 (0) 125 Cometary structure coincident with m (0) 125 Cometary structure coincident with m (0) 126 Ultrawide binary Gaia2 42807567266 (1268) III or more evolved YSO ALLW
IFO 123 Coincident with flat-spectrum YSO candidare SSTGLMC G035.2913+00.8076k4.80ro8/k21IFO 124 Matches to proto-stellar clump 34.93 \cdot 0.338.4ifo	 IFO 123 Coincident with flat-spectrum YSO candidate SSTGLMC G035.2913+00.8076 K 4.80 IFO 124 Matches to proto-stellar clump 34.93+0.338 / IFO 125 Cometary structure coincident with massive YSO IRAS 18527+0301 IFO 126 Ultrawide binary Gaia2.28075672668953984 is the closest, Class III or more evolved YSO ALLWISE J18555-4005, [Fe II] reported by p16 Class III or more evolved YSO ALLWISE J18555-4005, [Fe II] reported by p16 IFO 131 Class I/II massive YSO IRAS 18555-40056, [Fe II] reported by p16 IFO 134 Close to YSO cadidate SSTGLMC G039.2199+0.8638, G039.2060+00.8818 in West IFO 134 Close to YSO addiate SSTGLMC G039.2199+0.8638, G039.2060+00.8818 in West IFO 134 Close to YSO AGAL G039.708-01.237 and IFO 145 IFO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 145 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 IFO 146 Hα PN candidate, yet aligned with IFO 144, 145, H₂ knots IFO 146 Hα PN candidate, yet aligned with IFO 144, 145, H₂ knots IFO 146 Hα PN candidate, yet aligned with IFO 144, 145, H₂ knots IFO 146 IC 056 to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 IFO 151 Clonicdent with class 1 and flat-SED YSOs, [TBP2010] L673 YSO 15 & IFO 151 IFO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 IFO 153 Cuncident with class 1 II/photosphere SSTOERC G049.5851-0.3814 a 0.60 IFO 154 Class II/photosphere YSO IRAS 192361456 Connected to Class 1 YSO SSTOERC G049.1657 IFO 165 Connected to Class 1 YSO SSTOERC G049.5851-0.3814 a 0.60 a 0.60 IFO 165 Connected to Class 1 YSO SSTOERC G049.5851-0.3814 a 0.60 <li< td=""><td>andidate SSTGLMC G035.2913+00.8076 +0.338 J aassive YSO IRAS 18527+0301 686953984 is the closest, VISE J185522 49+030130.3 in 30° distance +0056, [Fe u] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518</td><td> O 123 Coincident with flat-spectrum YSO cc O 124 Matches to proto-stellar clump 34.934 O 125 Cometary structure coincident with m O 125 Cumetary structure coincident with m Class III or more evolved YSO ALLW </td></li<>	andidate SSTGLMC G035.2913+00.8076 +0.338 J aassive YSO IRAS 18527+0301 686953984 is the closest, VISE J185522 49+030130.3 in 30° distance +0056, [Fe u] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518	 O 123 Coincident with flat-spectrum YSO cc O 124 Matches to proto-stellar clump 34.934 O 125 Cometary structure coincident with m O 125 Cumetary structure coincident with m Class III or more evolved YSO ALLW
$ \begin{array}{rcrcrclim} \mbox{IFO} 124 \mbox{Matches to proto-stellar clump 34} 9.3-0.338 \mbox{Mass} 8520 \mbox{Tracture coincident with massive YSO/RAS 18527+0301 \mbox{Cometary structure coincident with massive YSO/RAS 18527+0301 \mbox{Cometary structure coincident with massive YSO/RAS 18527+0301 \mbox{Cometary structure coincident with massive YSO/RAS 18552-409-0301 \mbox{Cometary structure coincident with massive YSO/RAS 18552-409-0301 \mbox{Cometary structure coincident with massive YSO RAS 18552-409-0301 \mbox{Cometary structure coincident with massive YSO RAS 18555-40056, \mbox{Fe u1 reported by p16} \mbox{FO} 131 \mbox{Class III nor more evolved YSO ALLWISE J185522.49+0301 \mbox{Cometary stellar clump is coincident with TFO 131 \mbox{Class VII massive YSO RAS 18555-40056, \mbox{Fe u1 reported by p16} \mbox{FO} 134 \mbox{Class VII massive YSO RAS 19945-0518 \mbox{FO} 194 \mbox{Relater} with YSO Cadidate STGLMC G039.2199+00.8638. G0392060+00.8818 in West with the transmostation that the transmostation the transmostation that the transmostation the transmostation that the transmostation that the transmostation the transmostation that the transmostation the transmostation that the transmostation the transmostation that the transmostation that the transmostation that the transmostation the transmostation that the transmostation that$	 IFO 124 Matches to proto-stellar clump 34.9340.338.1 IFO 125 Cometary structure coincident with massive YSO IRAS 1852740301 IFO 126 Ultrawide binary Gaia2 428075672668953984 is the closest, Class III or more evolved YSO ALLWISE J185522.49+030130.3 in 30" distance Class III or more evolved YSO ALLWISE J185522.49+030130.3 in 30" distance IFO 131 Class I/II massive YSO IRAS 18555+0056, [Fe u] reported by p16 IFO 134 Close to YSO candidate SSTGLMC G035.2868-00.9528, proto-stellar clump is coincident a 2.48 IFO 134 Close to YSO candidate SSTGLMC G035.2199+00.8638. G039.2060+00.8818 in West IFO 134 Close to YSO candidate (2039.2199+00.8638. G039.2060+00.8818 in West IFO 134 Aligned with YSO AGAL G039.2199+00.8638. G039.2060+00.8818 in West IFO 134 Aligned with YSO AGAL G039.708-01.237 and IFO 145 IFO 145 Elongated and pointing toward YSO AGAL G039.708.10.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708.10.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708.10.237 and IFO 145 IFO 146 Hz PN candidate, yet aligned with IFO 144, 145, Hz hous IFO 146 Hz PN candidate, yet aligned with IFO 144, 145, Hz hous IFO 146 Hz PN candidate, yet aligned with IFO 145 IFO 146 Hz PN candidate, yet aligned with IFO 145, Hz hous IFO 151 Coincident with class 1 and flat-SED YSOs, [TBP2010] L673 10 and 13 IFO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 10 and 13 IFO 154 Class III/photosphere YSO SSTOERC G049.5851-0.3814 a 2.60 IFO 156 Connected to Class 1 YSO ISKE II/photosphere SSTOERC G049.5851-0.3814 a 2.60 IFO 165 Connected to Class 1 YSO ISKE II/photosphere SSTOERC G049.5851-0.3814 a 2.60 IFO 165 Connected to Class 1 YSO IRAS 192364456 A Energes from Class 1 YSO IRAS 192364456 A Energes from Class 1 YSO IRAS 1923644	 +0.338 J +0.338 J aassive YSO IRAS 18527+0301 686953984 is the closest, VISE J185522 49+030130.3 in 30" distance +0056, [Fe u] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518 	 O 124 Matches to proto-stellar clump 34.934 O 125 Cometary structure coincident with m O 126 Ultrawide binary Gaia2 42807567266 Class III or more evolved YSO ALLW
$ \begin{array}{rcl} \mbox{FO} \ 125 \ Cometary structure coincident with massive YSO IRAS 18527+0301 \ FO \ 126 \ Ultrawide binary Gaia2 428075672668953984 is the elosest, \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \ \$	 IFO 125 Cometary structure coincident with massive YSO IRAS 1852740301 IFO 126 Ultrawide binary Gaia2 4280756726686953984 is the closest. IFO 126 Ultrawide binary Gaia2 4280756726686953984 is the closest. Class III or more evolved YSO ALLWISE J185522.49+030130.3 in 30" distance IFO 131 Class I/II massive YSO IRAS 18555+0056, [Fe I] reported by p16 IFO 134 Close to YSO candidate SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West IFO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West IFO 134 Close to YSO candidate SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West IFO 144 Aligned with YSO AGAL G039.2199+00.8638. G039.2060+00.8818 in West IFO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 145 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 IFO 145 Elongated and pointing toward YSO AGAL G049.1652+0.2159 at W 51 (or foreground, see k09) IFO 154 Class III/photosphere YSO SSTOERC G049.1662+0.2159 at W 51 (or foreground, see k09) IFO 154 Class III/photosphere YSO SSTOERC G049.1662+0.2159 at W 51 (or foreground, see k09) IFO 165 Surrounding YSOs, e.g. class III/photosphere YSO R5871-0.3814 IFO 165 Surrou	aassive YSO IRAS 18527+0301 586953984 is the closest, VISE J185522.49+030130.3 in 30° distance +0056, [Fe II] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518	10 125 Cometary structure coincident with m 126 Ultrawide binary Gaia2 42807567266 Class III or more evolved YSO ALLW
$ \begin{array}{rcrcrc} \mbox{IFO} \ 126 \ Ultrawide binary Gaia2 \ 4280756726686953984 fis the closest, \\ Class III or more evolved YSO ALLWISE 118552. 49+030130.3 in 30" distance \\ \mbox{Class III or more evolved YSO ALLWISE 118552. 49+030130.3 in 30" distance \\ \mbox{Class III or more evolved YSO ALLWISE 118552. 49+030130.3 in 30" distance \\ \mbox{Class III massive YSO IRAS 18555+0056, [Fe i] reported by p16 \\ \mbox{FO 131 Class III massive YSO RAS 18555+0056, [Fe i] reported by p16 \\ \mbox{FO 134 Close to YSO candidate SSTGLMC G039.219+00.838, G0392060+00.8818 in West \\ \mbox{FO 145 Close to YSO candidate SSTGLMC G039.219+00.838, G0392060+00.8818 in West \\ \mbox{FO 141 Composent matches to massive YSO IRAS 19045-0518 \\ \mbox{FO 144 Aligned with YSO AGAL G039.708-01.237 and FPO 144 \\ \mbox{FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and FPO 144 \\ \mbox{FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and FPO 144 \\ \mbox{FO 146 Har PN candidate, yet aligned with IFO 144, 145, H2 knots \\ \mbox{FO 165 IC onicident with class I and flat-SED YSOs, [TBP2010] L673 I0 and 13 \\ \mbox{FO 165 IC onicident with class I and flat-SED YSOs, [TBP2010] L673 I0 and 13 \\ \mbox{FO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 I0 and 13 \\ \mbox{FO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 I0 and 13 \\ \mbox{FO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 I0 and 13 \\ \mbox{FO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 I0 and 13 \\ \mbox{FO 151 Coincident with class I IIP photosphere SSTOERC G049.1662-0.2159 at W 51 (or foreground, see A09) k 5.40 \ s17 \\ \mbox{FO 163 Surrounding YSO, e.e., class IIIP photosphere SSTOERC G049.5010-0.5922, \\ \mbox{FO 163 Surrounding YSO IS e.e., class IIIP photosphere SSTOERC G049.56351-0.3814 \ s5.40 \ s17 \\ \mbox{FO 163 Surrounding YSO IS e.e., class IIP photosphere SSTOERC G049.5010, 602-0.2159 \ s17 \\ \mbox{FO 163 Surrounding YSO IS e.e., class IIP photosphere SSTOERC G049.5010, 602-0.2159 \ s13 \$	 IFO 126 Ultrawide binary Gaia2 4280756726686953984 is the closest, Class III or more evolved YSO ALLWISE J185522.49+030130.3 in 30" distance IFO 131 Class I/II massive YSO IRAS 18555+0056, [Fe n] reported by p16 IFO 134 Close to YSO candidate SSTGLMC G035.2868-00.9528, proto-stellar clump is coincident a 2.48 IFO 136 Flat spectrum YSO SSTGLMC G035.2868-00.9528, proto-stellar clump is coincident a 2.48 IFO 141 Compact component matches to massive YSO IRAS 19045+0518 IFO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 145 IFO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 151 Coincident with class I and flat-SED YSOS, [TBP2010] L673 I 0 and 13 IFO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 I 0 and 13 IFO 154 Class II//photosphere YSO SSTOERC G049.060-00.5921, 0.3814 a 5.40 b 155 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 c 160 S urrounding YSOS, e.g., class III/photosphere SSTOERC G049.5851-0.3814 a 5.40 b 160 G connected to Class I YSO IRAS 192644650 c 160 S urrounding YSOS, e.g., class III/photosphere SSTOERC G049.5851-0.3814 c 160 G connected to Class I YSO IRAS 192644650 c 160 Connected to Class I YSO IRAS 19264456 c 161 A southern compact jet of massive YSO IRAS 19264456 c 17	586953984 is the closest. VISE J185522.49+030130.3 in 30" distance +0056, [Fe II] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518	O 126 Ultrawide binary Gaia2 42807567266 Class III or more evolved YSO ALLW
Class III or more evolved YSO ALLWISE J18552.2,49+030130.3 in 30" distance FO 131 Class I/II massive YSO IRAS 18555+0056, [Fe II] reported by p16 c 1.10 c 13/p16 FO 134 Close to YSO candidate SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West a 2.48 ro8 FO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West a 2.48 ro8 FO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West a 2.48 ro8 FO 144 Compact component matches to massive YSO IRAS 19045+0518 FO 144 Aligned with YSO AGAL G039.708-01.237 and FO 1445 FO 144 Aligned with YSO AGAL G039.708-01.237 and FO 1444 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and HO 1444 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and HO 1444 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and HO 1444 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and HO 1444 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and HO 1444 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and HO 1444 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and HO 1444 FO 146 H α PN candidate, yet aligned with FO 144, 145, hJz knots FO 156 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & FO 151 FO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & FO 161 FO 153 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & FO 161 FO 155 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & FO 151 FO 156 Connected to Class 1 YSO [RML2017] MC2 M105 spectral index of flat (FS) according to $k21$ FO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 FO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 FO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 FO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 FO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 FO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 FO 176 Southern diffuse	Class III or more evolved YSO ALLWISE J18552249+030130.3 in 30° distance FO 131 Class I/II massive YSO IRAS 18555+0056, [Fe u] reported by p16 c 1.10 FO 134 Close to YSO candidate SSTGLMC G035.2868-00.9528, proto-stellar clump is coincident a 2.48 FO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West c 3.60 FO 141 Compact component matches to massive YSO IRAS 19045+0518 FO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 145 FO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 144 c 7 0.60 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 a 2 0.60 FO 146 Ha PN candidate, yet aligned with IFO 144, 145, H ₂ kmost c 7 0.60 FO 146 Ha PN candidate, yet aligned with IFO 144, 145, H ₂ kmost c 7 0.60 FO 146 Ha PN candidate, yet aligned with IFO 144, 145, H ₂ kmost c 7 0.60 FO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 10 and 13 FO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 FO 154 Class III/photosphere SSTOERC G049.1662+0.2159 at W 51 (or foreground, see k09) k 5.40 FO 155 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 FO 156 Connected to Class I YSO SSTOERC G049.9010-00.5922, see cording to k21 FO 165 Surrounding YSOs, e.g., class III/photosphere SSTOERC G049.5851-0.3814 c 7 5 5.40 FO 165 Connected to Class I YSO [RML2017] MC2 M105 FO 176 A southern compact jet of massive YSO IRAS 19236+1456 c 7 5 0.00 FO 176 A southern compact jet of massive YSO IRAS 19236+1456 c 7 5 0.00 FO 175 A southern compact jet of massive YSO IRAS 19236+1456 c 7 5 0.00 FO 175 A southern compact jet of massive YSO IRAS 19236+1456 c 7 5 0.00 FO 175 A southern compact jet of massive YSO IRAS 19236+1456 c 7 5 0.00 FO 175 A southern compact jet of massive YSO IRAS 19236+1456 c 7 5 0.00 FO 175 A southern compact jet of massive YSO IRAS 19236+1456 c 7 5 0.00 FO 175 A southern compact jet of massive YSO IRAS 19236+1456 c 7 5 0.00 FO 175 A southern compact jet of massive YSO IRAS 19236+1456 c 7 5 0.00 FO 175 A southern compact jet	VISE J185522.49+030130.3 in 30" distance +0056, [Fe II] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518	Class III or more evolved YSO ALLW
IFO 131 Class I/II massive YSO IRAS 18555+0056, [Fe n] reported by p16c1.10c13/p16IFO 134 Close to YSO candidate SSTGLMC G035.2868-00.9528, proto-stellar clump is coincidenta 2.48 ro8IFO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in Westa 2.48 ro8IFO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in Westa 2.48 ro8IFO 141 Compact component matches to massive YSO IRAS 19045+0518c 3.60 $c13$ $c13$ IFO 144 Compact component matches to massive YSO IRAS 19045+0518b 0.60 $u18$ IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144b 0.60 $u18$ IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144b 0.60 $u18$ IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144b 0.60 $u18$ IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144b 0.60 $u18$ IFO 146 Ha PN candidate, yet aligned with IFO 144, 145, H ₂ knotsa 0.60 $u18$ IFO 151 Coincident with class 1 and flat-SED YSOS, [TBP2010] L673 15 & IFO 151k 0.60 $u10$ IFO 152 Close to HH 1186, 4.2 arcscc away from [TBP2010] L673 15 & IFO 151k 0.60 $u10$ IFO 153 Connected to Class ITyphotosphere SSTOERC G049.5851-0.3814a 0.60 $u10$ IFO 154 Closs III/photosphere YSO SSTOERC G049.5851-0.3814a 0.60 $u10$ IFO 155 Close to HH 1186, 4.2 arcscc away from [TBP2010] -0.5922,s <td< td=""><td> IFO 131 Class I/II massive YSO IRAS 18555+0056, [Fe n] reported by p16 IFO 134 Close to YSO candidate SSTGLMC G035.2868-00.9528, proto-stellar clump is coincident a 2.48 IFO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West a 2.48 IFO 141 Compact component matches to massive YSO IRAS 19045+0518 IFO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 145 IFO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 145 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 151 Coincident with class I and flat-SED YSOS, [TBP2010] L673 10 and 13 IFO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 IFO 154 Class III/photosphere YSO SSTOERC G049.900-00.5922, IFO 154 Class III/photosphere YSO SSTOERC G049.9010-00.5922, Spectral index of flat (FS) according to k21 IFO 165 Connected to Class I YSO ISMOS e.g. (SIRAS 19236+1456 IFO 174 Emerges from Class I YSO IRAS 19236+1456 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 </td><td>+0056, [Fe II] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 1.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518</td><td>10 101 DI VILLING WILLING TO 101 DI VILLING</td></td<>	 IFO 131 Class I/II massive YSO IRAS 18555+0056, [Fe n] reported by p16 IFO 134 Close to YSO candidate SSTGLMC G035.2868-00.9528, proto-stellar clump is coincident a 2.48 IFO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West a 2.48 IFO 141 Compact component matches to massive YSO IRAS 19045+0518 IFO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 145 IFO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 145 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 IFO 151 Coincident with class I and flat-SED YSOS, [TBP2010] L673 10 and 13 IFO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 IFO 154 Class III/photosphere YSO SSTOERC G049.900-00.5922, IFO 154 Class III/photosphere YSO SSTOERC G049.9010-00.5922, Spectral index of flat (FS) according to k21 IFO 165 Connected to Class I YSO ISMOS e.g. (SIRAS 19236+1456 IFO 174 Emerges from Class I YSO IRAS 19236+1456 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 	+0056, [Fe II] reported by p16 G035.2868-00.9528, proto-stellar clump is coi 1.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518	10 101 DI VILLING WILLING TO 101 DI VILLING
IFO 134 Close to YSO candidate SSTGLMC G035.2868-00.9528, proto-stellar clump is coincidenta2.48ro8IFO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in Westa2.48ro8IFO 141 Compact component matches to massive YSO IRAS 19045-0518c3.60c13IFO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 145b0.60u18IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144b0.60u18IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144b0.60u18IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144b0.60u18IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144b0.60u18IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144b0.60u18IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144b0.60u18IFO 146 Ha PN candidate, yet aligned with IFO 144, 145, H ₂ knotsa0.60u10IFO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 10 and 13k0.60t10IFO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151k0.60t10IFO 154 Class III/photosphere YSO SSTOERC G049.9010-00.5922,set from Class I YSO SSTOERC G049.9010-00.5922,set k09k5.40IFO 165 Connected to Class I YSO [RML2017] MC2 M105rrrrrIFO 165 Connected to Class I YSO [RML2017] MC2 M105rrr3.09 </td <td> FO 134 Close to YSO candidate SSTGLMC G035.2868-00.9528, proto-stellar clump is coincident a 2.48 FO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. G0392.2060+00.8818 in West a 2.48 FO 141 Compact component matches to massive YSO IRAS 19045+0518 FO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 145 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 FO 146 Har PN candidate, yet aligned with IFO 144, 145, H₂ knots FO 151 Coincident with class I and flat-SED YSOS, [TBP2010] L673 10 and 13 FO 151 Coincident with class I and flat-SED YSOS, [TBP2010] L673 10 and 13 FO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 10 and 13 FO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 10 and 13 FO 153 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 FO 153 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & FO 151 FO 154 Class III/photosphere YSO SSTOERC G049.1662+0.2159 at W 51 (or foreground, see k09) K 2.40 FO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922, spectral index of flat (FS) according to k21 FO 165 Connected to Class I YSO [RML2017] MC2 M105 FO 174 Emerges from Class I YSO [RML2017] MC2 M105 FO 174 Emerges from Class I YSO [RML2017] MC2 M105 FO 174 Emerges from Class I YSO [RMS 19236+1456 FO 175 A southern compact jet of massive YSO IRAS 19236+14</td> <td>G035.2868-00.9528, proto-stellar clump is coi 0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518</td> <td>U 151 CLASS I/ II MASSIVE I DU IKAD 1000</td>	 FO 134 Close to YSO candidate SSTGLMC G035.2868-00.9528, proto-stellar clump is coincident a 2.48 FO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. 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G039.2060+00.8818 in West sive YSO IRAS 19045+0518	U 151 CLASS I/ II MASSIVE I DU IKAD 1000
IFO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in Westa $-$ ro8/k21IFO 141 Compact component matches to massive YSO IRAS 19045+0518c3.60c13IFO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 145b0.60u18IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144b0.60u18IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144b0.60u18IFO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144b0.60u18IFO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 10 and 13a0.60t10IFO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151k0.60t10IFO 153 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151k0.60t10IFO 163 Surrounding YSOs, e.g., class II/photosphere SSTOERC G049.5851-0.3814a5.40s17IFO 163 Surrounding YSOs, e.g., class II/photosphere SSTOERC G049.5851-0.3814a5.40s17IFO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922,spectral index of flat (FS) according to k21c5.40s17IFO 165 Connected to Class I YSO [RML2017] MC2 M105spectral index of flat (FS) according to k21c5.40s17IFO 165 Connected to Class I YSO [RML2017] MC2 M105ffffffIFO 165 Connected to Class I YSO [RML2017] MC2 M105ffffffIFO 165 Connected to Class I YSO [RML2017] MC2 M105 <td< td=""><td>FO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West a - FO 141 Compact component matches to massive YSO IRAS 19045+0518 c 3.60 FO 141 Compact component matches to massive YSO IRAS 19045+0518 c 3.60 FO 141 Compact component matches to massive YSO IRAS 19045+0518 c 3.60 FO 141 Compact component matches to massive YSO IRAS 19045+0518 c 3.60 FO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 144 b 0.60 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 b 0.60 FO 146 Ha PN candidate, yet aligned with IFO 144, 145, H₂ knots a 0.60 FO 151 Coincident with class I and flat-SED YSOS, [TBP2010] L673 YSO 15 & IFO 151 a 0.60 FO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 a 0.60 FO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 a 0.60 IFO 153 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 a 0.60 IFO 154 Class III/photosphere YSO SSTOERC G049.1662+0.2159 at W 51 (or foreground, see k09) k 5.40 IFO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922, spectral index of flat (FS) according to k21 a 5.40</td><td>0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518</td><td>O 134 Close to YSO candidate SSTGLMC C</td></td<>	FO 136 Flat spectrum YSO SSTGLMC G039.2199+00.8638. G039.2060+00.8818 in West a - FO 141 Compact component matches to massive YSO IRAS 19045+0518 c 3.60 FO 141 Compact component matches to massive YSO IRAS 19045+0518 c 3.60 FO 141 Compact component matches to massive YSO IRAS 19045+0518 c 3.60 FO 141 Compact component matches to massive YSO IRAS 19045+0518 c 3.60 FO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 144 b 0.60 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 b 0.60 FO 146 Ha PN candidate, yet aligned with IFO 144, 145, H ₂ knots a 0.60 FO 151 Coincident with class I and flat-SED YSOS, [TBP2010] L673 YSO 15 & IFO 151 a 0.60 FO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 a 0.60 FO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 a 0.60 IFO 153 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 a 0.60 IFO 154 Class III/photosphere YSO SSTOERC G049.1662+0.2159 at W 51 (or foreground, see k09) k 5.40 IFO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922, spectral index of flat (FS) according to k21 a 5.40	0.2199+00.8638. G039.2060+00.8818 in West sive YSO IRAS 19045+0518	O 134 Close to YSO candidate SSTGLMC C
FO 141 Compact component matches to massive YSO IRAS 19045+0518 c 3560 c13 FO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 145 b 0.60 u18 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 b 0.60 u18 FO 146 H α PN candidate, yet aligned with IFO 144, 145, H $_2$ knots FO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 10 and 13 FO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 10 and 13 FO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 10 and 13 FO 154 Class III/photosphere YSO SSTOERC G049.1662+0.2159 at W 51 (or foreground, see k09) k 5.40 s17 FO 165 Surrounding YSOs, e.g., class III/photosphere SSTOERC G049.5851-0.3814 a 5.40 s17 FO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922, see k09) k 5.40 s17 FO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922, see k09) k 5.40 s17 spectral index of flat (FS) according to k21 FO 174 Emerges from Class I YSO [RML2017] MC2 M105 FO 174 Emerges from Class I YSO [RML2017] MC2 M105 FO 174 Emerges from Class I YSO [RML2017] MC2 M105 FO 175 A southern compact jet of massive YSO IRAS 19236+1456 FO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 FO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 FO 176 Southern diffuse emission from massive YSO IRAS 19236+1456	 FO 141 Compact component matches to massive YSO IRAS 19045+0518 FO 144 Aligned with YSO AGAL G039.708-01.237 and IFO 145 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 FO 145 Elongated and pointing toward YSO AGAL G039.708-01.237 and IFO 144 FO 146 Ha PN candidate, yet aligned with IFO 144, 145, H₂ knots FO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 10 and 13 FO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 10 and 13 FO 154 Class III/photosphere YSO SSTOERC G049.1662+0.2159 at W 51 (or foreground, see k09) K 0.60 FO 163 Surrounding YSOs, e.g., class III/photosphere SSTOERC G049.5851-0.3814 a 5.40 FO 163 Surrounding YSO s, e.g., class III/photosphere SSTOERC G049.5851-0.3814 a 5.40 FO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922, spectral index of flat (FS) according to k21 FO 174 Emerges from Class I YSO [RML2017] MC2 M105 FO 175 A southern compact jet of massive YSO IRAS 19236+1456 FO 175 A southern compact jet of massive YSO IRAS 19236+1456 	sive YSO IRAS 19045+0518	O 136 Flat spectrum YSO SSTGLMC G039.
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IFO 146 H α PN candidate, yet aligned with IFO 144, 145, H $_2$ knots IFO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 10 and 13 IFO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 10 and 13 IFO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 IFO 154 Class III/photosphere YSO SSTOERC G049.1662+0.2159 at W 51 (or foreground, see k09) IFO 154 Class III/photosphere YSO SSTOERC G049.5851-0.3814 IFO 165 Surrounding YSOs, e.g., class III/photosphere SSTOERC G049.5851-0.3814 IFO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922, spectral index of flat (FS) according to k21 IFO 174 Emerges from Class I YSO [RML2017] MC2 M105 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 IFO 176 Southern diffuse emission from massive YSO IRAS 19236	IFO 146 Hα PN candidate, yet aligned with IFO 144, 145, H ₂ knots IFO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 10 and 13 IFO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 IFO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 IFO 154 Class III/photosphere YSO SSTOERC G049.1662+0.2159 at W 51 (or foreground, see k09) IFO 154 Class III/photosphere YSO SSTOERC G049.1662+0.2159 at W 51 (or foreground, see k09) IFO 165 Surrounding YSOs, e.g., class III/photosphere SSTOERC G049.5851-0.3814 IFO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922, spectral index of flat (FS) according to k21 IFO 174 Emerges from Class I YSO [RML2017] MC2 M105 IFO 174 Emerges from Class I YSO [RML2017] MC2 M105 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 IFO 175 A southern compact jet of massive VSO IRAS 19236+1456	AGAL G039.708-01.237 and IFO 144	O 145 Elongated and pointing toward YSO /
IFO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 10 and 13 a 0.60 t10 IFO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 k 0.60 t10 IFO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 k 0.60 t10 IFO 154 Class II/photosphere YSO SSTOERC G049.1662+0.2159 at W 51 (or foreground, see k09) k 5.40 s17 IFO 163 Surrounding YSOs, e.g., class II/photosphere SSTOERC G049.5851-0.3814 a 5.40 s17 IFO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922, sec k09) k 5.40 s17 IFO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922, spectral index of flat (FS) according to k21 a 5.40 s17 IFO 165 Connected to Class I YSO [RML2017] MC2 M105 c 5.40 s17 IFO 174 Emerges from Class I YSO [RML2017] MC2 M105 c 3.09 c13 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 a 3.39 c13 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 a 3.39 c13	 IFO 151 Coincident with class I and flat-SED YSOs, [TBP2010] L673 10 and 13 IFO 152 Close to HH 1186, 42 arcsec away from [TBP2010] L673 YSO 15 & IFO 151 IFO 154 Class III/photosphere YSO SSTOERC G049.1662+0.2159 at W 51 (or foreground, see k09) IFO 163 Surrounding YSOs, e.g., class III/photosphere SSTOERC G049.5851-0.3814 IFO 163 Surrounding YSOs, e.g., class III/photosphere SSTOERC G049.5851-0.3814 a 5.40 IFO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922, spectral index of flat (FS) according to k21 IFO 174 Emerges from Class I YSO [RML2017] MC2 M105 IFO 174 Emerges from Class I YSO [RML2017] MC2 M105 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 a 2.00 a 2.00 a 3.09 a 3.09 a 4.00 	⁷ O 144, 145, H ₂ knots	O 146 H α PN candidate, yet aligned with IF
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IFO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922, c 5.40 sl7 spectral index of flat (FS) according to k21 IFO 174 Emerges from Class I YSO [RML2017] MC2 M105 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456 a 3.39 c13 IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 a 3.39 c13	IFO 165 Connected to Class I YSO SSTOERC G049.9010-00.5922, c 5.40 spectral index of flat (FS) according to k21 IFO 174 Emerges from Class I YSO [RML2017] MC2 M105 IFO 175 A southern compact jet of massive YSO IRAS 19236+1456	otosphere SSTOERC G049.5851-0.3814	O 163 Surrounding YSOs, e.g., class III/pho
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IFO 176 Southern diffuse emission from massive YSO IRAS 19236+1456 c13		SO IRAS 19236+1456	O 175 A southern compact jet of massive YS
		ive YSO IRAS 19236+1456	O 176 Southern diffuse emission from massi
,		foreground, 0.3814	C G049.1662+0.2159 at W 51 (or tosphere SSTOERC G049.5851-(G049.9010-00.5922, b k21 7] MC2 M105 60 IRAS 19236+1456 ve YSO IRAS 19236+1456

	Table 3 (cont d)				
IFO #	YSO / YSOc Name M	Morphology	q	Reference	
			[kpc]	YSO Counterpart	Dist
IFO 177 Di	use structure in contact with YSO AGAL G052.488-00.172	c	1.60	u18	u18
IFO 178 Cl	ss T YSO 1 (~ $10M_{\odot}$) or 2 (~ $5M_{\odot}$) in k18	а	1.60	ro8/c13/k18	u18
IFO 179 h	he proximity of Class I YSO SSTGLMC G053.1570+00.0735	а	1.60	ro8/k15	u18
IFO 180 Cl	se to Class I YSO SSTGLMC G053.1612+00.0668 and multiple YSOs	а	1.60	ro8/k15	u18
IFO 181 Su	rounding Class I YSO SSTGLMC G053.1266+00.0499	k	1.60	ro8/k15	u18
IFO 183 Co	ncident with MSXDC G053.25+00.04 MM6 & ISOGAL-P J192931.1+175954 (Class I)	а	1.60	k15	u18
IFO 184 Fla	-spectrum YSO 2MASS J19293167+1800581	k	1.60	ro8/k15	u18
IFO 185 SS	CGLMC G053.2389+00.0552 (Class I), 2MASS J19293167+1800581 (FS)	а	1.60	ro8/k15	u18
IFO 186 2N	ASS J1930021941755001 (FS)	c	ı	k15	ı
IFO 189 Mu	ltiple compact structures surrounding ES-NW of Massive YSO MSX6C G056.3694-00.6333	а	6.40	c13	c13
IFO 190 Di	use structure on South of Class II YSO SSTGLMC G058.4098+00.3279	а	2.80	ro8/k21	v13
IFO 192 Th	head of the cometary structure matches EGO G058.09-0.34, one of the low-mass EGOs	k	0.74	ro8/cy13	cy13
IFO 193 Tw	biconical structures, tails toward SE & W, Class I YSO SSTGLMC G058.4801-00.2205 at the center	q	6.15	ro8/k21	m21
IFO 194 Th	head of cometary structure corresponds to YSO candidate SSTGLMC G059.0069-00.2481	c	ı	ro8	ı
IFO 198 AI	gned with star-forming region IRAS 19410+2336 and IFO 199	а	2.20	c13	113
IFO 199 An	orphous IFO points toward star-forming region IRAS 19410+2336	а	2.20	c13	113
IFO 203 Co	npact IFO at the East of biconical outflow S87, emerging from ${\sim}20{ m M}_{\odot}$ pre-main-sequence object	а	2.20	b89	113
*Column counterpart	. Morphology categories : b - bipolar, c - cometary k - knot-like, a - amorphous. Column 4. Distance c lassification and distance.	of counterpar	t in kpc	Column 5. Refere	nces of
*Referenc	ss : a20 - Areal et al. (2020), b18 - Bailer-Jones et al. (2018), b89 - Barsony (1989), b98 - Beck et al. (1	998), b99 - I	3elikov	et al. (1999), c13 - (Cooper
et al. (2013)	cl6 - Csengeri et al. (2016), cy13 - Cyganowski et al. (2013), e03 - Egan et al. (2003), e17 - Elia et al. (20 19. Inneinente et al. (2018), 150 - Vana et al. (2000), 145 - Vana et al. (2015), 149 - Vina et al. (2018)	017), g07 - C	iuarcello	o et al. (2007), i13 -	Immer a of ol
(2102) .u. (2013), 116 -	Jio - Jayashighe et al. (2010), 807 - Naug et al. (2007), 819 - 810 et al. (2015), 810 - 810 et al. (2016), Li et al. (2016), m13 - Sánchez-Monge et al. (2013), m15 - Messineo et al. (2015), m16 - Marton et al. (2	, kz1 - Null 016), m21 -]	et al. (2 Mège et	al. (2021), m96 - M	folinari
et al. (1996)	n01 - Nielbock et al. (2001), p09 - Povich et al. (2009), p10 - Povich & Whitney (2010), p16 - Paron et	al. (2016), r(9 - Rag	an et al. (2009), r10) - Rygl
et al. (2010) (2002) +10 -	r17 - Retes-Romero et al. (2017), ro8 - Robitaille et al. (2008), s14 - Sabin et al. (2014), si - SIMBAD, Tsijali et al. (2010) +15 - Traficante et al. (2015) +20 - Tian et al. (2020) 1018 - Frunbart et al. (2018) 102	s17 - Saral e 2 - Hromhart	t al. (20 et al 70	017), t02 - Thompso 022) v13 - Venezia	n et al. ni et al
(2013), v18	Varricatt et al. (2018), v20 - Vioque et al. (2020), x19 - Xu et al. (2019), yu12 - Yuan et al. (2012)	ummhro - 7	· · · · ·		
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objects
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with
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Table 4.

			Iac	DIE 4. IFOS ASSOCIATED WITH REFOIG-RATO ODJECIS			
IFO # HH Name	d [kpc]	Exciting Source	Region	Comment	Morp.	Referenc Type Di	ist
IFO 041 IFO 042 IFO 043 IFO 043 IFO 044	2.14	N-HH	Eagle nebula	Parsec-scale HH Parsec-scale HH Parsec-scale HH Parsec-scale HH	9 9 9 9	m82 b ¹ m82 b ¹ m82 b ¹ m82 b ¹ m82 b ¹	66 66 66
IFO 118 HH 722	Y			Wrongly identified as HH 172 (Nikogossian et al. 2007)	в	c94 d	92
IFO 119 IFO 120 GGD 30 IFO 121	1.70	GGD30IR	GM 2-30	Binary. One of them is Class I YSO, $A_v \sim 25$ mag, bipolar envelope SE-NW side Binary. One of them is Class I YSO, $A_v \sim 25$ mag, bipolar envelope SE-NW side Binary. One of them is Class I YSO, $A_v \sim 25$ mag, bipolar envelope SE-NW side	ה ה ה	s07 s s07 s s07 s	87 87 87
IFO 153 HH 32 IFO 155 HH 250A IFO 195 HH 803	$^{-}0.20\pm0.0$ 0.30	3 AS353 HH 250-IRS	Aquila Rift Aquila Rift	T-Tauri binary, H α is coincident, [Fe II] detection reported (d03) Class I binary, bow-shock, H α detected. Launched 3500 yrs ago, adjacent to IFO 153 SW of parsec-scale HH (7.5 pc), H α detected. L _{bol} =580 L \odot	k a a	h74 r d97 s c04 d	06 97 92
IFO 196 HH 165 IFO 200 HH 365	2.40	1548C27 IRS1		NE counterpart of the HH 803, very faint Central structure of parsec-scale HH, both perpendicular jet and curvature are identified	99	c81 d a97 d	92 92

classification and distance .: a97 - Alten et al. (1997), b99 - Belkov et al. (1999), c04 - McGroarty et al. (2004), c81 - Craine et al. (1981), c94 - Cappellaro et al. (1994), d03 - Davis *Column 3. Distance of counterpart in kpc. Column 7. Morphology categories : b - bipolar, c - cometary, k - knot-like, a - amorphous. Column 8. References of counterpart et al. (2003), d92 -Dent & Aspin (1992), d97 -Devine et al. (1997), h74 -Herbig (1974), m82 -Meaburn & White (1982), r06 -Rice et al. (2006), s07 -Smith et al. (2007), s87 -Scoville et al. (1987), s97 -Sakamoto et al. (1997)

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IFO # UCHII, CHII Name	d [kpc]	Type	Comment	Refere Type	nce Dist
IFO 024 IRAS 18116-1646 4. IFO 025 IRAS 18116-1646 4.5	50	UCHII UCHII	Outer arc, East of cometary H II region Diffuse filament filling inside, coincident with the boundary of cometary H II region	k18 k18	u18 u18
IFO 026 IRAS 18/16-1646 4. IFO 054 M 17 UC1 1.(50 60±0.30	UCHII HCHII	Outer arc, West of cometary H II region Multiple shell-like features around hypercompact H II region	k18 s04	u18 n01
IFO 071 IRAS 18222-1317 4.0	94	UCHII	Partial shell-like + amorphous structure around UCHII	k18	114
IFO 082 IRAS 18308-0503 2. IFO 083 IRAS 18308-0503 2.5	000	UCHII UCHII	Amorphous, compact feature North of UCHII Amorphous, diffuse feature South of UCHII	96d 696	113 113
IFO 086 IRAS 18314-0720 9.3	30±0.40	UCHII	Shell-like structures, West of UCHII	k18	k03
IFO 087 IRAS 18314-0720 9. IFO 089 IRAS 18317-0757 44	30±0.40 60	UCHI	Multiple shell-like structures encircling UCHII A mombans structures enveloning the South-East of HICHII	k18 k18	k03 113
IFO 090 IRAS 18317-0757 4.0	50	UCHII	Compact knot-like structure North-East of UCHII	k18	113
IFO 097 G025.3809-00.1815 3.3 IFO 107 IRAS 18416-0420 3.3	80 20	UCHI	Amorphous structures in the central region of UCHII + South-Western jet Shell-like structure stretching from the center of UCHII to West	k18 k18	a09 118
IFO 127 IRAS 18530+0215 9.(88	UCHIN	Amorphous + knot-like feature South of UCHII	k18	113 113
IFO 132 G35.2S 2.4	40±0.50	CHN	Oval structure, slightly limb-brightened, brighter at East. Compact H II region (radius : 0.14 pc)	f11	z13
IFO 137 [CPA2006] N75 2.8	80	UCHII?	Northern shell-like structure. Single 09.5V star. UCHII precursors, Mol 91C1–5, along the structure	tre m98	z13
IFO 138 UCEC 8 2.3	80	UCHII?	Ultra-compact embedded cluster (UCEC) with a radius of 0.09 pc. Eight possible ionizing stars, most of them are early-B.	al2	113
IFO 148 IRAS 19111+1048 4.4	40	UCHII	About 19 central sources of spectral type earlier than B0.5 (v06)	k18	113
IFO 149 IRAS 19111+1048 4.4	40	UCHII	Diffuse emission North of IFO 148	k18	113
IFO 156 IRAS 19191+1538 2.	10	UCHII	Partial shell-like structure in the West, compact structure superposed in the North	k18 1-10	113
IFO 204 IRAS 19446+2505 2.7	20 L	JCHII/HII	CORNISH UCHII G061,4763+00.0892 in East, H II region G061.4758+00.0913 in West	k18 k18	113
*Column 3. Distance of count Kobulnicky (2012), b96 - Bronfr (2018), 113 - Lumsden et al. (201 (2018), v06 - Vig et al. (2006), z1	arpart in k aan et al. (3), 114 - L 3 - Zhu et	pc. Colum 1996), d2(eahy et al. al. (2013)	 m 6 : References of counterpart classification and distance : a09 - Anderson & Bania (2009), a12 - 3 - de la Fuente et al. (2020), ft1 - Floebrich & Ioannidis (2011), k03 - Kolpak et al. (2003), k18 - K. (2014), m98 - Molinari et al. (1998), n01 - Nielbock et al. (2001), s04 - Sewilo et al. (2004), u18 - L 	- Alexan Kalcheva - Urquhart	ter & et al. et al.

	Reference ype Dist	r07 u12	r07 u12	p09 n01	p09 n01	p09 n01	p09 n01	p09 n01	k97 r03	f21 113	f21 k95	f21 c09	
ole 6. IFOs Associated with HII Region or HII Region Candidates	Comment	Diffuse structure in West of H II region	Diffuse structure coincident with O stars/YSOs in the central region. UCHII G10.15-0.34 is embedded	The South-Western large-scale structure of M 17	Diffuse filaments at the South of M 17	The North-Eastern large-scale structure of M 17	Diffuse filaments at the South of M 17	? Amorphous structure located in the extremity of M 17, star Gaia2 4097840529185958272 in West	The elongated structure along the radio emission of H II region [KC97c] G015.1-00.9	Shell-like diffuse emission at the Western boundary of compact radio	Large-scale, multiple filamentary structures in W 51B	Large-scale, miscellaneous amorphous structures in W 51A	
Tab	Type	IIH	IIH	.30 HII	.30 HII	.30 HII	.30 HII	.30 HII?/YSO?	HII	IIH	ΗH	WH	
	d [kpc]	3.55 ^a	3.55^{a}	1.60 ± 0	1.60 ± 0	1.60±0	1.60 ± 0	1.60 ± 0	9 1.98	5.40	6.00	5.50	
	IFO # HII / HIIc Name	IFO 011 G10.2-0.3 (W 31)	IFO 012 G10.2-0.3 (W 31)	IFO 055 M 17 S	IFO 058 M 17 S (South)	IFO 062 M 17 N	IFO 064 M 17 S (South)	IFO 066 M 17 EB ^b	IFO 067 [KC97c] G015.1-00.	IFO 158 W 51A	IFO 159 W 51B	IFO 161 W 51A	

^aIFO distances marked with 'a' are controversial values. Column 3. Distance of counterpart in kpc. Column 6 : References of counterpart classification and distance : c09 -Clark et al. (2009), f21 - Fujita et al. (2021), k95 - Koo et al. (1995), k97 - Kuchar & Clark (1997), 113 - Lumsden et al. (2013), n01- Nielbock et al. (2001), p09 -Povich et al. (2009), r03 - Russeil (2003), r07 - Roshi et al. (2017), u12 - Urguhart et al. (2012)

^bThe region M 17 EB (extended bubble) is defined in detail by Povich et al. (2009).

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Table 7.

Ś	IFO #	PN / PNc Name	Type	q	Size	Mc	rpholog	sy		Reference	
				[kpc]	[bc]	[Fe II]	H_2	Ha I	Distance	Morphology	Typ
	IFO 005 NGC	6537	PN	2.81 ± 0.45	0.49^{a}	Bp	Bmps	Bmps	3	1,2	1
	IFO 006 NGC	6537	PN	2.81 ± 0.45	0.43^{a}	Bp	Bmps	Bmps	ю	1,2	ľ
	IFO 008 HRD	S G008.362–0.623	PNc	ı	<i>-a</i>	Bas	Bmps	I	ī	1	4
<i>•</i>	IFO 009 HRD	S G008.362–0.623	PNc	ı	-a	Bas	Bmps	ı	ı	1	4
	IFO 027 PNG	G014.5+00.4	PNc	ı	I	В	В	S	ı	1,2	S
	IFO 050 SSTC	JLMC G015.7993-00.0063	PN	ı	I	Ia	1?		ı	9	I
	IFO 073 G020	9+00.8	PNc	ı	I	A	Is	Aa	ı	1,2	٢
	IFO 081 M 1-	51	PN	1.59	0.18	Bp^b	ı	В	8	2	ľ
	IFO 095 IRAS	3 18348-0616	PNc/UCHII?	6.5	0.50	Ema	Ema	1	18	9	20,1
	IFO 112 IRAS	\$ 18458-0213	PN	4.90	0.13	Bsa	\mathbf{Bs}		6	1	1
	IFO 114 PNG	032.1+00.1	PNc/HII?	4.90	0.42	Ia	ı		6		10,
	IFO 129 IRAS	8 18551+0159	PNc	4.30 ± 0.50	0.45 ^a	Br	\mathbf{Bs}	В	11	1	4
	IFO 130 IRAS	3 18551+0159	PNc	4.30 ± 0.50	0.45 ^a	Br	\mathbf{Bs}	В	11	1	4
	IFO 157 PNG	050.4+00.7	PNc/sym?	5.30	3.08	$^{\mathrm{Bp}}$	\mathbf{Bs}	В	6	1	12,
	IFO 164 PNG	050.8+00.0	PNc	18.46 ± 2.59	0.26	I	ı		17	·	1
	IFO 173 IRAS	3 19234+1627	PN	4.70–9.50	0.31-0.6	4 Ears	Er	Emrs	13	1,2	ı
	IFO 188 IRAS	\$ 19312+1950	pAGBc/YSO?	2.42	0.46	Ams	\mathbf{As}		14	9	15,

we showed the morphological classification of H α counterpart in HASH (Hong Kong/AAO/Strasbourg/H α ; Parker et al. 2016). For the *For the morphological classification, we adopted an 'ERBIAS' classifier (Parker et al. 2006 and references therein), and when listed, The 'ERBIAS' represents the main structure of PN by its elliptical, round, bipolar, irregular, and asymmetric or quasi-stellar morphology. The substructure is further classified into the sub-classifier 'ampirs', which explains asymmetry, multiple structures, point symmetry, ring, and internal structure. When no counterpart was detected, we marked '-'. Column 3. PNc: PN candidate, sym: symbiotic star candidate, Parker et al. (2016), 3 - Navarro et al. (2012), 4 - Froebrich et al. (2015), 5 - Miszalski et al. (2008), 6 - this study (UWISH2), 7 - Boissay Lumsden et al. (2013), 10 - Ferrero et al. (2015), 11 -Zhu et al. (2013), 12 - Sabin et al. (2014), 13 - Yang et al. (2016), 14 - Vickers et al. (2015), 15 - Nakashima et al. (2016), 16 - Cordiner norphological classification of the H₂ counterpart and H α classification of IFO 129, 130, and 173, we made use of Gledhill et al. (2018). and pAGBc: post-AGB candidate. Column 4. Distance of counterpart in kpc. Column 5. Physical scale of an IFO in the adopted distance. Column 7. References of distance, $H_2/H\alpha$ morphology, counterpart type in Column 3, respectively. Ref. 1 - Gledhill et al. (2018), 2 et al. (2016), 17 - Irabor et al. (2018), 18 - Cichowolski et al. (2018), 19 - Bronfiman et al. (1996), 20 - Kanarek et al. (2015) et al. (2012): Miscellaneous Emission Nebulae (MEN) list, 8 - Phillips (2004), 9

^aPart of bipolar lobes.

^bPN with quadrupolar morphology is included in the bipolar category considering the formation mechanisms of the two could be analogous; see Manchado et al. (1996) for details.

nebula candidates
LBV1
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Table 8.

ence d	$^{4}_{4,1}$ $^{1}_{4,2}$ $^{2}_{4,2}$
Refere Type	v 1 1 2
d [kpc]	$1.55 \le 6.5 \le 6.5 \le 6.0, 2.5^{+2.4}_{-1.3}$
LBV / LBVc Name	HD 168625 26.47+0.02 26.47+0.02 26.47+0.02 [KW97] 37-17 (=LS1)
IFO #	IFO 065 IFO 102 IFO 103 IFO 162

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References of counterpart classification and distance : 1 -Clark et al. (2003) : assuming 1.8 mag kpc⁻¹, 2 - Okumura et al. (2000), 3 - Clark et al. (2009) : observational and theoretical constraints + W51 membership + parallax, 4 -Bailer-Jones et al. (2018) : Gaia DR2 parallax, 5 - Hutse-*Column 3. Distance of counterpart in kpc. Column 4. mekers et al. (1994) MANUS

IFO #	SN	R Name	d	Ftot
IFO #	G-Name	Other Name(s)	[kpc]	$[10^{-17} \text{ W m}^{-2}]$
IFO 007	G8 7-0 1	W 30	4 5	_a
IFO 015	$G_{11}^{-0.1}$	W 50	J 1 1	1000.00
IFO 013	G15.0+0.2	-	10.0	1090.00
IFO 058	G13.9+0.2	-	10.0	7.54
IFO 009	$C_{10} = 0.1$	-	5.0	21.30
IFO 070	$C_{10} = 0.1$	-	2.0	1.30
	C10.9-1.1	-	2.0	14.10
IFO 070	G18.9-1.1	-	2.0	2.33
IFO 078	G18.9-1.1	-	2.0	192.00
IFO 079	G18.9-1.1	-	2.0	1.89
IFO 080	G21.8-0.6	Kes 69	5.2	627.00
IFO 084	G21.5-0.9	-	4.6	6.59
IFO 085	G21.5-0.9	-	4.6	59.70
IFO 088	G23.3-0.3	W 41	4.2	584.00
IFO 093 ^b	G28.8+1.5	-	^c 3.4	0.39
IFO 098	G27.8+0.6	-	2.0	167.00
IFO 105	G27.4+0.0	Kes 73	8.5	274.00
IFO 106	G28.6-0.1	_	9.6	197.00
IFO 113	G31.9+0.0	3C 391	7.1	3423.55
IFO 117	G32.8-0.1	Kes 78	4.8	114.00
IFO 128	G34.7-0.4	W 44	2.8	3942.00
IFO 139	G39.2-0.3	3C 396	8.5	618.00
IFO 140	G41.5+0.4	_	4.1	1421.45
IFO 143	G41.1-0.3	3C 397	10.0	1691.00
IFO 147	G43.3-0.2	W49B	10.0	4739.00
IFO 160	G49.2-0.7	W51C	6.0	582.37

Table 9. IFOs Associated with SNRs

*Distances are from Lee et al. (2019). See Lee et al. (2019) for original references.

^aDue to significant background errors, flux was not derived and not included in the statistics of fluxes.

^bDistance from Shan et al. (2018).

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^cPartially covered in the UWIFE survey.

Unknown-type
with
Associated
IFOs
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Table 10. IFOs Associated with Unknown-type	
IFO # Comment R	Reference
IFO 001 Amorphous structure elongated from NE to SW. The bright star 15 arcsec East to the IFO	s19
is ALLWISE J180138.06–224227.8, a luminosity class of giant	
IFO 002 Compact IFO, H II region in the vicinity, H ₂ at South is slightly more extended	ı
IFO 017 Located in between late-type star [MHS2002] 196 and dark cloud SDC G12.804+0.055	ı
IFO 028 Biconical structure in the North-East of IKAS 18134-1838, however, relation is uncertain	
IFO U68 An amorphous source in the middle of a possible evolved star J18222/.66-1/1556.6 and the two cross having a bigh measure motion of South	m19
ure two stats naving a nigh proper mouon at sound IFO 096 Snatially coincident with large-scale radio shell 625.8+0.2	c18
IFO 109 IFO at the border of AllWISE J1845601.48-001716.3. Hy emission at North	
IFO 110 Amorphous IFO, possibly surrounding MIR source AllWISE J184515.66–031606.9	ı
IFO 116 Amorphous IFO, no well-known sources in the vicinity	ı
IFO 133 Cometary IFO elongated North to South, AllWISE J185905.19+004837.9 at the North	ı
IFO 135 Amorphous IFO adjoining K-band source UGPS J185923.30+010415.0 which was not detected in J- and H-bands	ı
IFO 142 Bow-shock morphology with apex on NW side, in between possible giant star at South and	s19
evolved/Class III YSO candidate at North	m16
IFO 150 Compact structure coincident with H ₂ , ALLWISE J191530.41+132737.2 showing MIR photometric anomaly is located in the vicinity	s17
IFO 166 Knot-like small-scale IFO	
IFO 167 Coincident with extended H ₂ , but with the shifted peak position	·
IFO 168 A compact, faint knot in between IFO 167 and 169	ı
IFO 169 Knot-like diffuse IFO	
IFO 170 Knot-like IFO in the middle of IFO 166-169 and IFO 172	
IFO 171 Knot-like IFO in the middle of IFO 166-169 and IFO 172	ı
IFO 172 Compact IFO, small-scale H ₂ in the vicinity	I
IFO 182 Amorphous IFO with the bright northern component coincident with MIR source GLIMPSE G052.9805-00.0394	
IFO 187 Knot-like IFO. Possible evolved star WISE J193120.72+192006.1 and	m19
dwart star Gaia DR2 1825442262813/50912 in South	819
IFO 191 Diffuse elongated IFO from East to West. MIPSGAL source MG059.1989+00.5106 at North, MIR source ALLWISE J194013.97+232657.5	·
superposed to the IFO, red giant star ZMASS J19401620+2326377 at East	
IFO 19/ Stignup elongated compact IFO, the center of NGC 0823 is located at East	1 -
IFO 201 Cometary IFO with its South-Eastern head part coincident with ALLWISE J194611.21+221534.3, unlish is a huminosity along of Annue considers	819
*Column 3 References of counternart candidates : c18 - Cichowolski et al (2018) m16 - Marton et al (2019) r10 -	Roeser et a
(2010), s17 - Solarz et al. (2017), s19 - Stassun et al. (2019)	
*Counterpart candidates around IFO 166–172: 1: YSO candidate SSTGLMC G050.3746-00.4149 2: AGB candidate SSTGLMC G050.3756-00.4	214 3: MH
2624 4: EGO G050.36-0.42 5: YSO candidate SSTGLMC G050.3666-00.3944 6: YSO candidate SSTGLMC G050.3647-00.3979 7: YSO candidate	e SSTGLM
GUDU.3587-00.4125 8: YSO candidate SST GLMC GUDU.3673-00.4089 9: YSO candidate SST GLMC GUDU.3691-00.4096 10: YSO candidate SST GLMC 00.4095 11: YSO candidate SSTGI MC G050 3741-00.4083 17: YSO AGAT G050 376-00 421 13: YSO candidate SSTGI MC G050 3762-00 4205 14: YS	SO candidat
SSTGLMC G050.3575-00.4182	

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DATA AVAILABILITY

The data underlying this article are available to download from http://gems0.kasi.re.kr/uwife/.

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Figure A.1. Example of artifacts in the $[Fe \pi]$ -H image. 1) The residual of star subtraction shown as cross stripes 2) A variable and saturated star with a diffraction pattern 3) Arc-shape ghosts near a very bright star. The diffraction pattern of a bright star is superposed. 4) Electronic cross-talk near a bright star 5) Diffraction pattern of a bright star 6) High proper-motion star. (Black: bright, white: dark.)

APPENDIX A: EXAMPLES OF ARTIFACTS

APPENDIX B: THE UWIFE [FEII]-H AND GPS 3-COLOR IMAGES OF IFOS

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Figure B.1. IFOs with **YSO** counterpart candidates in continuum-subtracted [Fe II] and three-color *KHJ* UKIDSS images in the same field of view. The format for these images is the same as that of Fig. 1. Yellow crosses denote adjacent YSOs in the field of view. Red contours are $H_2 2.12 \mu m$ emission adopted from UWISH2.



Figure B.1. – Continued





Figure B.1. – Continued



Figure B.1. – Continued



Figure B.1. – Continued





Figure B.2. Continuum-subtracted [Fe II] images of IFOs with PN counterpart candidates. The units on the UWIFE [Fe II]-H are DNs, with the darker colour denoting a higher DN. The corresponding source names for each IFO are shown. The yellow cross marks the central position of the counterpart. The images of IFO 73, 095, 112, 129, 157, and 164 are smoothed with a two-pixel Gaussian. In all images, North is at the top and east is on the left side.





Figure B.3. IFOs with counterpart candidates unknown in continuum-subtracted [Fe II] images as in Fig. 1. IFO 95 is a PN-IFO inside the field of view. The format for these images is the same as that of Fig. 1.



Figure B.3. – Continued