Kent Academic Repository

Paladini, R., Mottram, J.C., Veneziani, M., Traficante, A., Schisano, E., Giardino, G., Falgarone, E., Urquhart, J.S., Harrison, D. L., Joncas, G. and others (2021) ***The Planck Submillimeter Properties of Galactic High-mass Star-forming Regions: Dust Temperatures, Luminosities, Masses, and Star Formation Efficiency.*** **Astrophysical Journal, 911 (1). ISSN 0004-637X.**

Downloaded from

https://kar.kent.ac.uk/86670/ The University of Kent's Academic Repository KAR

The version of record is available from

https://doi.org/10.3847/1538-4357/abe711

This document version

Author's Accepted Manuscript

DOI for this version

Licence for this version

UNSPECIFIED

Additional information

Versions of research works

Versions of Record

If this version is the version of record, it is the same as the published version available on the publisher's web site. Cite as the published version.

Author Accepted Manuscripts

If this document is identified as the Author Accepted Manuscript it is the version after peer review but before type setting, copy editing or publisher branding. Cite as Surname, Initial. (Year) 'Title of article'. To be published in ***Title of Journal*** , Volume and issue numbers [peer-reviewed accepted version]. Available at: DOI or URL (Accessed: date).

Enquiries

If you have questions about this document contact ResearchSupport@kent.ac.uk. Please include the URL of the record in KAR. If you believe that your, or a third party's rights have been compromised through this document please see our Take Down policy (available from https://www.kent.ac.uk/guides/kar-the-kent-academic-repository#policies).

DRAFT VERSION FEBRUARY 16, 2021
Preprint typeset using LaTeX style emulateapj v. 12/16/11

THE *PLANCK* SUBMILLIMETER PROPERTIES OF GALACTIC HIGH-MASS STAR FORMING REGIONS: DUST TEMPERATURES, LUMINOSITIES, MASSES AND STAR FORMATION EFFICIENCY

R. PALADINI[1,*], J. C. MOTTRAM[2], M. VENEZIANI[3], A. TRAFICANTE[4], E. SCHISANO[4], G. GIARDINO[5], E. FALGARONE[6], J. S. URQUHART[7] D. L. HARRISON[8,9], G. JONCAS[10], G. UMANA[11], S. MOLINARI[4],

Draft version February 16, 2021

ABSTRACT

Massive star formation occurs in the interior of giant molecular clouds (GMC) and proceeds through many stages. In this work, we focus on massive young stellar objects (MYSOs) and Ultra-Compact HII regions (UCH II), where the former are enshrouded in dense envelopes of dust and gas, which the latter have begun dispersing. By selecting a complete sample of MYSOs and UCH II from the Red MSX Source (RMS) survey data base, we combine *Planck* and IRAS data and build their Spectral Energy Distributions (SEDs). With these, we estimate the physical properties (dust temperatures, mass, luminosity) of the sample. Because the RMS database provides unique solar distances, it also allows investigating the *instantaneous* Star Formation Efficiency (SFE) as a function of Galactocentric radius. We find that the SFE increase between 2 and 4.5 kpc, where it reaches a peak, likely in correspondence of the accumulation of molecular material at the end of the Galactic bar. It then stays approximately constant up to 9 kpc, after which it linearly declines, in agreement with predictions from extragalactic studies. This behavior suggests the presence of a significant amount of undetected molecular gas at $R_G > 8$ kpc. Finally we present diagnostic colors that can be used to identify sites of massive star formation.

Subject headings: – – dust

1. INTRODUCTION

The initial stages of a massive star can be traced back to a giant molecular cloud (GMC), i.e. a cloud with a mass ranging from 10^5 $M_\odot$ to 10^6 $M_\odot$ (e.g Dame et al. 2001; Miville-Deschenes et al. 2017) and a linear size up to hundreds of pc. Within GMCs, very dense molecular cores (n $> 10^5$ cm-3, e.g. Giannetti et al. 2013) collapse, and evolve first into massive young stellar objects (MYSOs), and later, when the OB star begins to ionize the surrounding material, into Ultra-Compact H II regions (UCH II).

In this framework, characterizing clumps hosting massive star formation, an intermediate stage between GMCs and cores, is of primary importance: massive star formation is known to take place in cold ($T_d < 25$ K), massive ($M > 100$ $M_\odot$), luminous ($L > 10^3$ $L_\odot$) environments, so constraining dust temperatures, luminosities and masses of the clumps allows one to assess whether the clump under investigation is able to effectively form massive stars. In general, determining accurate masses and luminosities for the clumps is also crucial to assess their evolutionary stage on a mass-luminosity plot (e.g. Molinari et al. 2008).

In this work we are going to use the combined *Planck* High Frequency Instrument (HFI) (Tauber et al. 2010, Planck Collaboration 2011a, Lamarre et al. 2010; Planck HFI Core Team 2011a) and IRAS (Neugebauer et al. 1984) data to investigate a complete sample of clumps harboring MYSOs and UCH II regions. This study has two primary objectives: 1) to derive the properties (dust temperature, luminosity, mass, surface density) of these clumps and therefore fully characterize this evolutionary stage; 2) to make use of the estimated luminosities and masses, to compute an *instantaneous* Star Formation Efficiency (SFE) for the first 0.1 Myr (the approximate age of UCH II regions). For 2), we aim at addressing the question: At this fixed evolutionary stage, is there variation in the amount of gas relative to the star-formation (modulo the spread caused by the variation in age ?) And, what could cause variations in the efficiency that gas is converted to stars: galactic dynamics (sheer, bars), star formation mechanisms (cloud-cloud collisions), or hidden gas (e.g. undetected molecular material) ?

The combination of the spectral bands from *Planck* and IRAS is extremely effective in probing the Spectral Energy Distribution (SED) of sources such as MYSOs and UCH II. The coldest dust component present in these objects typically thermalizes at temperatures of the order of 15 - 20 K, and this makes the 850 μm *Planck* band a very accurate tool in estimating the total mass. The accuracy of the *Planck* mass measurements is also due to the fact that this experiment operated from space, which means that no filtering for atmospheric noise is applied at any stage of data processing, hence emission on all angular scales is preserved. This is not the case for ground experiments such as APEX/LABOCA, JCMT/SCUBA and CSO/Bolocam, where considerable flux

* Corresponding author: paladini@ipac.caltech.edu
[1] Infrared Processing Center, California Institute of Technology, 1200, E. California Blvd., Pasadena, CA 91125, USA
[2] Max Planck Institute for Astronomy, Konigstuhl 17, 69117 Heidelberg, Germany
[3] Science & Technology corporation, Olof Palmestraat 14, 2616 LR Delft, The Netherlands
[4] IAPS - INAF, Via Fosso Del Cavaliere 100, I-00133, Roma, Italy
[5] ESA, Science Operations Department, ESTEC, 2200AG Noordwijk, The Netherlands
[6] LERMA/LRA, Observatoire de Paris, PSL Research University, CNRS, Sorbonne Universites, UPMC, Universite Paris 06, Ecole Normale Superiore, 75005 Paris, France
[7] Centre for Astrophysics and Planetary Science, University of Kent, Canterbury, CT2 7NH, UK
[8] Institute of Astronomy, Madingley Road, Cambridge, CB3 0HA, UK
[9] Kavli Institute of Cosmology Cambridge, Madingley Road, Cambridge, CB3 OHA, UK
[10] Universite' Laval, Pavillon Alexandre-Vachon 1045, Avenue de la Medecine, Canada
[11] INAF - Osservatorio Astrofisico di Catania, Via S. Sofia 78, 95123 Catania, Italy

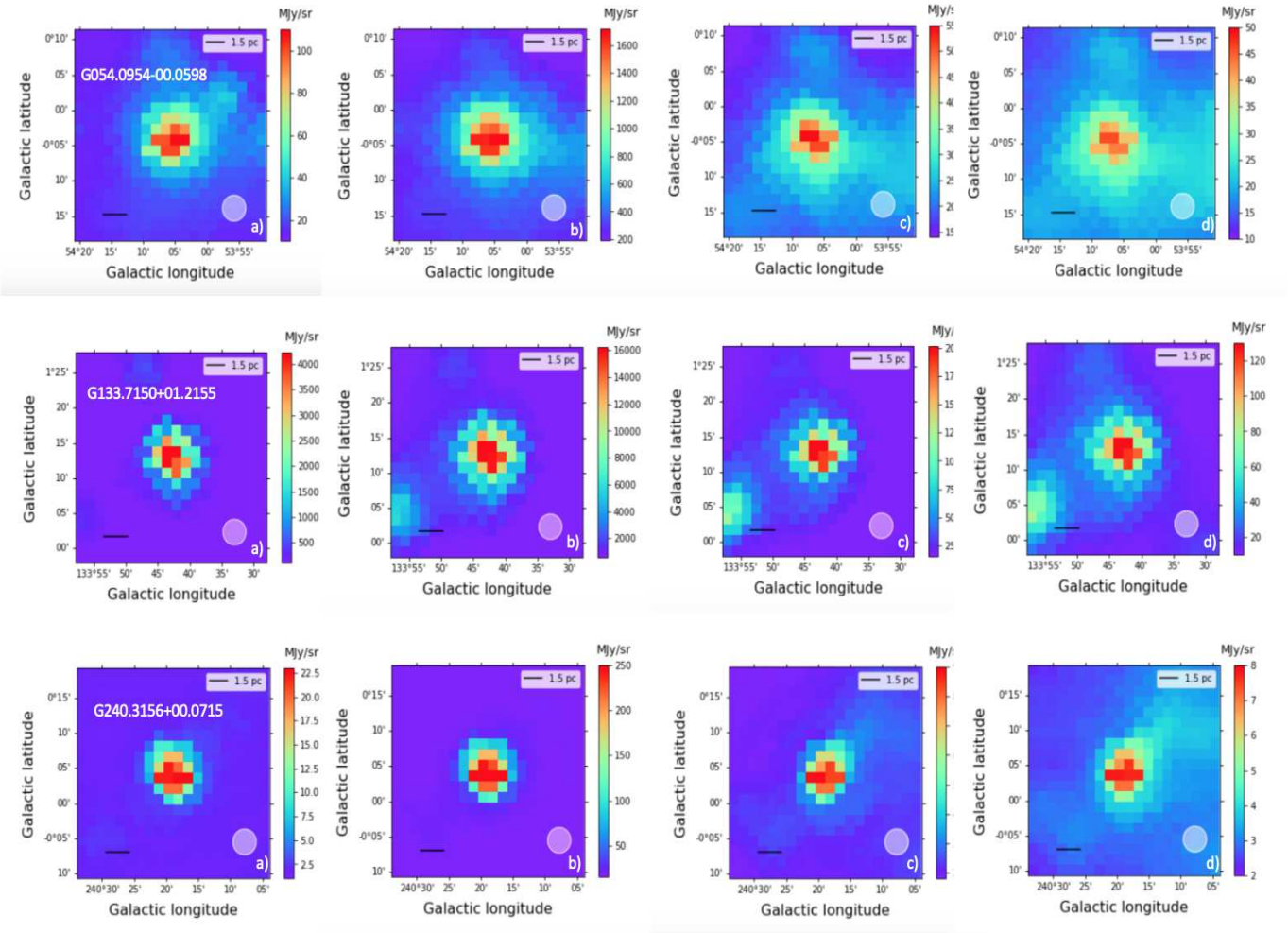


FIG. 1.— Examples of MYSOs and UCH II regions identified in the *Planck* data at the RMS position. From left to right: a) IRIS 25 μm; b) IRIS 100 μm; c) *Planck* HFI 857 GHz (350 μm); d) *Planck* HFI 353 GHz (850 μm). All the data are convolved to the common *Planck* HFI 353 GHz angular resolution (4.86′).

loss occurs, with effects on mass determination. An attempt to combine *Planck* and APEX/LABOCA data was carried out for the APEX Telescope Large Area Survey of the Galaxy (ATLASGAL, Schuller et al. 2009), a large and sensitive sub-millimeter survey of the inner Galactic plane. This work (Csengeri et al. 2016) though was limited to the investigation of the large-scale structure of cold dust.

Noteworthy, *Planck* and IRAS both mapped the whole sky, including both the inner and outer Galactic Plane. The outer Galactic disk ($R_G > 8.5$ kpc) has characteristics that sets it apart from the inner Galaxy. For instance, for $R_G > 13$ kpc, both the H_2 surface density (Scoville & Sanders 1987; Digel et al. 1996; Heyer et al. 1998) and the stellar disk (Robin et al. 1992; Ruphy et al. 1996) appear to exhibit a sharp decline. In addition, the metallicity, at these large Galactocentric distances, is only half solar ($\sim 0.5\ Z_{sol}$, Yong et al. 2005). These observational facts suggest a dramatic change in star formation activity with respect to the inner Galaxy.

We emphasize that the *Planck* and IRAS beams (of the order of 5 arcmins) are sensitive to structures of ~ 1.5 pc (at a distance of 1.5 kpc), which is the typical size of Galactic clumps. Although clumps will contain multiple cores, we expect MYSOs or UCH II regions to be the dominant component within each individual clump.

The paper is organized as follows. Section 2 describes the *Planck* and the IRAS data sets. Section 3 provides details on the samples selection and on the identification of the sources in the *Planck* data. Section 4 discusses the photometric measurements, the SEDs, as well as the estimate of the clump dust temperatures, luminosities, masses and surface densities. Section 5 presents the Star Formation Efficiency in terms of clump luminosity-to-mass ratio and its variations across the Galaxy. A summary is provided in the Conclusions (Section 6). Finally the Appendix contains an estimate of the colors that can be used as diagnostics to identify regions of massive star formation.

2. PLANCK DATA

Planck[13] (Tauber et al. 2010; Planck Collaboration 2011a) was the third-generation space mission to measure the anisotropy of the Cosmic Microwave Background (CMB). It observed the sky in nine frequency bands covering 30 - 857

[13] *Planck* (http://www.esa.int/Planck) is a project of the European Space Agency (ESA) with instruments provided by two scientific consortia funded by ESA member states (in particular the lead countries France and Italy), with contributions from NASA (USA) and telescope reflectors provided by a collaboration between ESA and a scientific consortium led and funded by Denmark.

GHz (i.e., 10,000 μm - 350 μm), with high sensitivity and angular resolution from 31′ to 5′. The Low Frequency Instrument (LFI, Mandolesi et al. 2010) covered the 30, 44 and 70 GHz bands with amplifiers cooled to 20 K. The High Frequency Instrument (HFI, Lamarre et al. 2010; Planck HFI Core Team 2011a) covered the 100, 143, 217, 353, 545 and 857 GHz bands with bolometers cooled to 0.1 K. Polarisation was measured in all but the two highest frequency bands (Leahy et al. 2010; Rosset et al. 2010). *Planck*'s sensitivity, angular resolution, and frequency coverage made it a powerful instrument for Galactic and extragalactic astrophysics as well as cosmology.

In this paper we use the *Planck* PR2 full channel, full mission temperature maps at nominal frequencies 353 (850 μm), 535 (560 μm) and 857 GHz (350 μm). We downloaded the maps from the IRSA archive[14]. The 353 GHz map are in T_{CMB} units, while the 535 and 857 GHz maps are in MJy/sr. To perform the conversion from CMB thermodynamic units to Rayleigh-Jeans brightness temperature units we make use of the conversion factor given in Table 3 of Planck Collaboration.X. (2016).

We adopt the effective beam sizes provided in the Explanatory Supplement[15]. These are: 4.86′, 4.84′ and 4.63′ at, respectively, 353, 545 and 857 GHz.

2.1. *Ancillary data*

In order to construct SEDs of the clumps, we complement the *Planck* data with IRAS (Neugebauer et al. 1984) data at 100, 60 and 25μm. In particular, we use the IRIS (Improved Reprocessing of the IRAS Survey, Miville-Deschenes & Lagasche 2005) maps, which benefit from the COBE-DIRBE (Hauser et al. 1998) calibration and zero point, as well from a better zodiacal light subtraction and destriping. We do not include the IRIS 12μm data since this band is significantly contributed to by emission from Polycyclic Aromatic Hydrocarbons (PAHs), whose modeling and characterization is beyond the scope of this paper. The Full-Width-Half-Maximum (FWHM) of the IRIS beams are, respectively, 3.8', 4.0' and 4.3', at 25, 60 and 100 μm.

Both the *Planck* and IRIS data are reprojected into the HEALPIX (Hierarchical Equal Area isoLatitude Pixelization, Gorski et al., 2005) format with Nside = 2048. All maps are smoothed to the *Planck* 353 GHz angular resolution (4.86′).

3. SAMPLE SELECTION

Our reference data base for the source selection is the RMS[16] survey (Lumsden et al. 2013), i.e. a campaign of follow-up observations designed to identify possible contaminants (e.g. evolved stars and planetary nebulae, PNe) of a sample of $\sim$ 2000 candidate MYSOs assembled by Lumsden et al. (2002) using MSX color selection criteria. Examples of the follow-up observations which have been undertaken by the RMS team are high-resolution (1") radio continuum observations (Urquhart et al. 2007a, 2009b) which have allowed the descrimination of UCH II regions from PNe, and near-IR spectroscopic measurements which have led to the identification of evolved stars (Clarke et al. 2006). This effort has made it possible to single out some 1500 (1420) MYSOs and UCH II regions with uniquely constrained distances (Urquhart et al. 2013) . Radial velocities have been obtained from ^{13}CO J = 1 - 0 and J = 2 -1 observations (Urquhart et al. 2007b, 2008b), and these have been coupled with the Reid et al. (2009) rotation curve to derive kinematic distances.

See Figure 1 for examples of MYSOs and UCH II regions from the RMS database in the IRAS/IRIS and *Planck* data.

Mottram et al. (2011) investigate the completeness of the RMS sample. To this end, they compute the volume of the Galaxy probed by the RMS survey at each luminosity L and assume that the RMS MYSOs and HII regions are distributed in a layer with a hole in the center, according to the model by Robin et al. (2003) for the thin disk Galactic stellar population. Following this method, they obtain that the survey is 100% complete for luminosities greater than $\sim 1\times10^4$ $L_\odot$. This selection leads to 731 sources.

The longitude and latitude distributions of the complete sample are shown in Figure 2. From the longitude distribution, it is evident that there is a higher source concentration towards the intersection of the line of sight with the Sagittarius and Scutum-Crux arms (first Galactic quadrant, $l \sim 30°$) and with the Norma and Scutum-Crux arms (fourth Galactic quadrant, $l \sim 330°$). Out of 731 sources, 209 are MYSOs candidate, 509 are UCH II regions, 13 are thought to be transition objects, i.e. sources that are older than MYSOs but not old enought to be classidied as UCH II regions. In addition, 151 ($\sim$ 20 %) are located at Galactocentric distances greater than 8.5 kpc, and 108 are part of complexes of multiple sources.

For each source in the RMS data, we create a 30′×30′ image centered on the source at each *Planck* frequency (353, 545, 857 GHz). Each image is visually inspected to assure that the emission in the considered region is not dominated by background emission. From this visual inspection we notice that strong sources of emission centered on the location of the RMS source are, with a few exceptions ($\sim$ 5%), present in the *Planck* HFI bands.

To estimate the angular size of these structures, we follow two complementary approaches. First we assume that all the sources are unresolved with respect to the *Planck* 353 GHz convolution beam (4.86′), second we perform a 2-dimensional gaussian fits and, for each clump, we derive a measure of the major and minor axis at 353 GHz (850 μm). In this case, not all the sources turn out to be resolved. Out of 731, only for 238 (32.5 %) we can determine the major and minor axis, with an average size of 5.6' $\pm$ 1.5'. In the following sections, we will only show the results for the unresolved case, while we will annotate in parenthesis the results for the fitted size case.

4. SEDS

To compute the flux in each band, we use the HEALpix aperture photometry code developed for Planck Collaboration (2011c). As input aperture radius, we use half the convolution beam FWHM (4.86') if the sources are unresolved, otherwise we use half of the estimated FWHM from the 2-d gaussian fit. In case the sources are quasi-spherically symmetric, i.e. their aspect ratio is $\sim$ 1, the major and minor axis are averaged together and half this average is used as the source aperture, $aper_s$. Alternatively, half of the major axis is used as the aperture.

After converting the maps in units of Jy/pixel, the pixels within an aperture equal to $aper_s$ are summed together. An estimate of the background is subtracted using a median estimator of the pixels within radii [$aper_s$, 2×$aper_s$]. Uncertainties are obtained by summing in quadrature the r.m.s of the values in the background annulus to the absolute calibration

[14] https://irsa.ipac.caltech.edu/data/Planck/release_2/
[15] https://wiki.cosmos.esa.int/planck−legacy−archive/index.php/Effective_Beams
[16] http://www.ast.leeds.ac.uk/RMS/

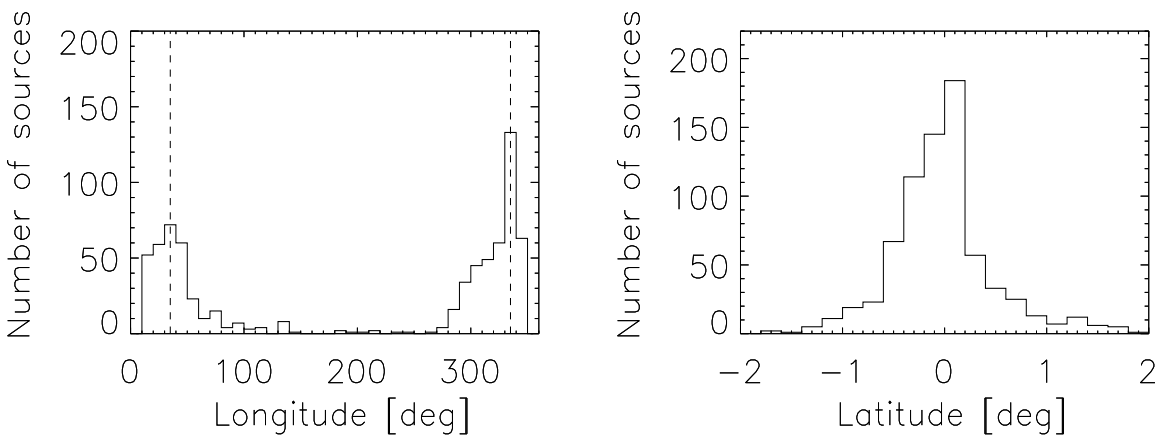


FIG. 2.— Longitude (left panel) and latitude (right panel) distributions of the sample of candidate MYSOs and UCH II regions above the 10^4 $L_\odot$ completeness limit. The dashed lines on the longitude plot denote, respectively, the Sagittarius and Scutum-Crux arms (at $l \sim 30°$) and the Norma and Scutum-Crux arms (at $l \sim 330°$).

uncertainties for each map.

For the sources in double (90) or triple systems (18), in order to avoid double counting, we have subdivided the total flux in the Planck beam by two or three, depending on the case.

Since the pioneering works of Chini et al. (1986a, 1986b, 1986c, 1987), it is known that, in order to explain observations of HII regions above and below $\sim 100\mu$m, one has to invoke the existence of a 2-temperature component dust distribution: a warm, low density population of dust grains situated in the proximity of the central source, and a colder dust population in the periphery of the cloud. The preliminary finding by Chini et al. was subsequently confirmed by data at increasingly higher spatial resolution and larger spectral coverage (e.g. Povich et al. 2007), as well as by sophisticated radiative transfer modeling (e.g. Zhang & Tan 2011). The SEDs obtained from the combination of the *Planck* and IRIS data show a behaviour similar to the one just described (Figure 3, blue points), as the IRIS 60μm and 25μm data points cannot be represented by the same grey-body as the measurements at longer wavelengths. Taking this fact into account, we fit our SEDs with the functional form:

$$S_\lambda = A_1 \left(\frac{\lambda}{\lambda_0}\right)^{-\beta} B_\lambda(T_c) + A_2 \left(\frac{\lambda}{\lambda_0}\right)^{-\beta} B_\lambda(T_w) \qquad (1)$$

where T_c and T_w are, respectively, the temperatures of the cold and warm components, and λ_0 is set to 100 μm.

To estimate the parameters A_1, A_2, T_c and T_w, we use a χ^2 goodness-of-fit method. For the spectral emissivity index, β, we test three different values: 1.8, 2.0 and 2.2. The best χ^2 are consistely given by fits performed with β = 1.8 (see Table 1), therefore this is the value that we adopt for both grey-bodies in the final runs.

Note that we do not perform the photometric SED fitting by using a more sophisticated modeling (e.g. Robitaille et al. 2006, 2007) for several reasons: at the angular resolution we are working, we are not sensitive to the parameters probed by these models (e.g. disk inner/outer radius, cavity opening angle, etc.); moreover we are probing the extent of clumps, while those frameworks were designed for describing the behavior of individual cores; finally UCH II represent a far too advanced evolutionary stage not accounted for by those models.

TABLE 1
SPECTRAL EMISSIVITY INDEX, β, AND CORRESPONDING AVERAGE SED FITTING χ^2.

β	χ^2
1.8	7.3
2.0	15.9
2.2	31.2

An illustration of the result of the fits is provided in Figure 3 (red solid line). Occasionally, in one (or more) HFI frequency band the best-fit model slightly over-/under-predicts the measured data points. This is a consequence of the adopted value for β. We have tested this hypothesis by making trial fits with the spectral emissivity index as a free parameter. Although in some cases the fit χ^2 improves, we prefer to keep β fixed and equal to the canonical value of 1.8, as the analysis of variations of β and of its degeneracy with respect to dust temperature (e.g. Schnee et al. 2007; Juvela et al. 2013, 2018) is beyond the scope of this paper.

4.1. *Dust temperatures distribution*

From the fitting procedure, we derive estimates for the cold (T_c) and warm (T_w) dust temperatures components, of $\overline{T_c}$ = 21.2±2.9 K (21.2±3.1 K for the fitted size case) and $\overline{T_w}$ = 54.0±4.6 K (53.8±4.4 K for the fitted size case), i.e. comparable to those found for more evolved HII regions (e.g. Povich et al., 2007; Paladini et al. 2012). For the uncertainties, rather than the modeling errors, we adopt the standard deviations of the sample, given that the former underestimate the true errors. König et al. (2017) investigate a small sample of MYSOs (36 sources) and UCH II regions (25 sources), and for these they derive dust temperatures from the modelling of the combined MSX (Egan et al. 2003), WISE (Wright et al. 2010), Hi-GAL and ATLASGAL data. They obtain mean temperatures of 28.1±3.6 K and 31.7±4.0 K for MYSOs and UCH II regions, respectively, which lie in between our average cold and warm components. This work was extended to the whole ATLASGAL sample by Urquhart et al. (2018) who find, for MYSOs and HII regions, values of dust temperature between $\sim$ 15 and 40 K.

We subdivide the sample into 6 Galactocentric bins of equal

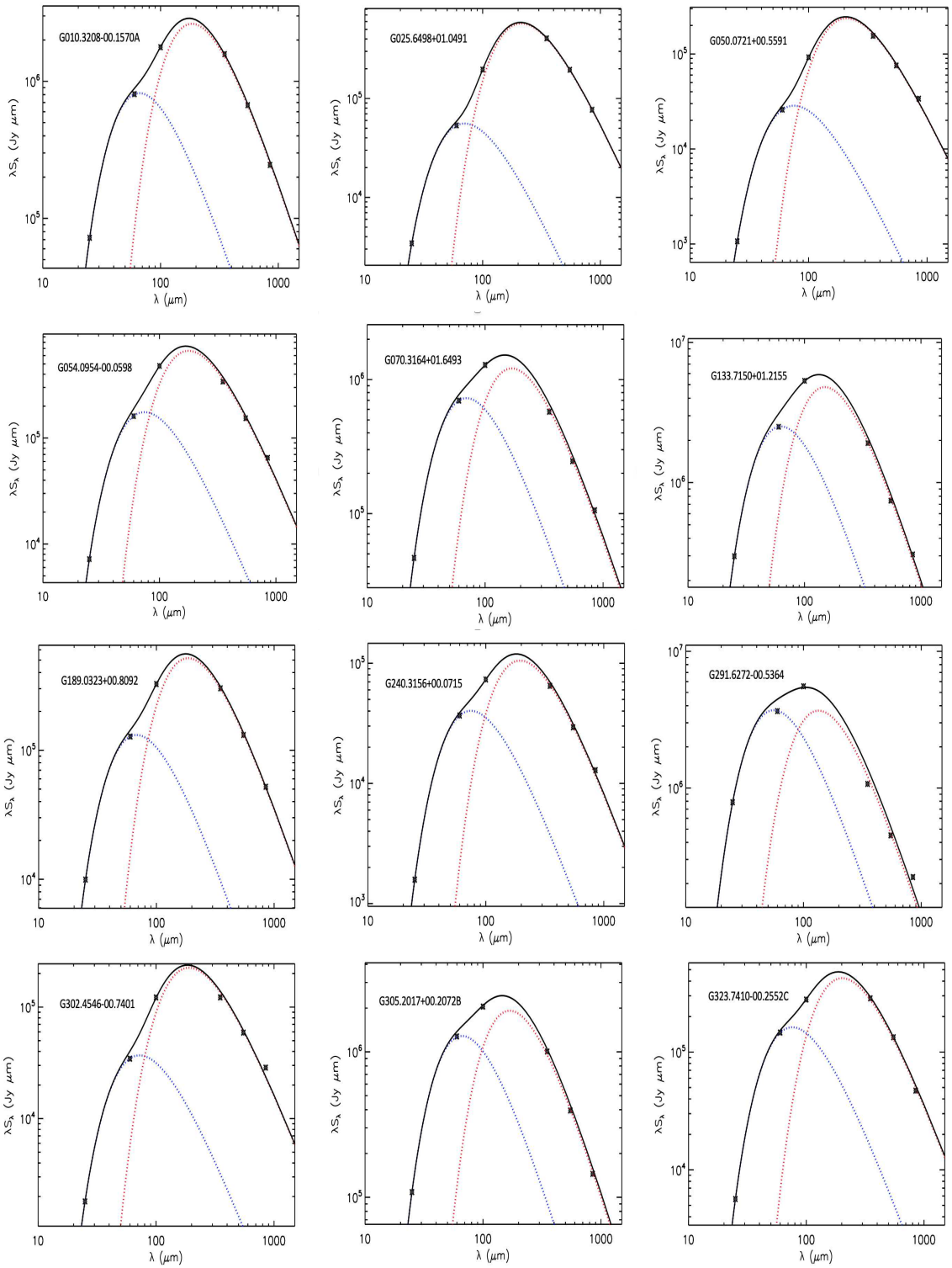


FIG. 3.— Examples SED fits to the *Planck* and IRIS clumps fluxes. The best-fit 2-temperature component model is shown (black solid line). Black diamonds denote the *Planck* 350, 500, 850 μm and IRIS 100, 60 and 25μm data points. The cold (blue-dashed line) and warm (red-dashed line) temperature components are also shown.

width (i.e. 2 kpc) and for each bin we compute the average T_c and T_w (see Table 2). The distributions of T_c and T_w vs. R_G are plotted in Figure 4 (top panel), and indicate a decreasing temperature for the cold component towards the outer Galaxy, in particular for $R_G > 10$ kpc, accompanied by an opposite trend for the warm component.

An important caveat to keep in mind is that, by binning according to distance from the Galactic center, we assume that the Galaxy is radially symmetric and, in doing so, we ignore the presence of spiral arms. To take this effect into account and explore even further dust temperature variations across the Galaxy, we split the sample according to longitude. We denote *inner-Galaxy* sources those located either in the first ($0° < l < 90°$) or fourth ($270° < l < 360°$) Galactic quadrants, and *outer-Galaxy* sources those either in the second or third Galactic quadrants ($90° < l < 270°$). We have 699 candidate MYSOs/UCH II regions in the *inner* sample and 32 in the *outer* one. Then, for each dust temperature component (T_c or T_w), we compare the histogram distribution for the *inner/outer Galaxy* sub-samples (Figure 4, middle panels). In this case the average temperatures are: $\overline{T_{c,i}}$ = 21.2±2.9 K (21.2±3.1 K for the fitted size case), $\overline{T_{c,o}}$ = 20.8±2.4 K (20.6±2.7 K); $\overline{T_{w,i}}$ = 53.7±4.4 K (53.5±4.3 K for the fitted size case), $\overline{T_{w,o}}$ = 59.5±4.6 K (58.8±4.6 K for the fitted size case), where $T_{c,i}$/$T_{c,o}$ and $T_{w,i}$/$T_{w,o}$ are the cold/warm dust temperatures for the *inner/outer Galaxy* sub-samples, respectively. The histogram distributions highlight a trend similar to what is found in Figure 4 (upper panels).

We interpret this result in the context of the large scale properties of the Galactic plane. A dust temperature gradient is known to characterize the diffuse emission along the Galactic Plane, with temperatures ranging from 14 - 15 K in the outer Plane, to $\simeq$ 19 K for the inner one (Planck Collaboration 2011m). The temperature enhancement towards the center of the Galaxy is typically explained as due to the presence of a high concentration of star forming regions, especially in correspondence of the molecular ring ($R_G \sim 5$ kpc). From our analysis, it appears that the small scale behaviour reflects the situation on larger scales, with clumps harboring massive stars at large Galactic radii being at lower temperatures with respect to their counterparts closer to the Galactic center. Since $T_D \propto X_{ISRF}^{1/(4+\beta)}$ – with T_D denoting the dust temperature, X_{ISRF} the intensity of the Interstellar Radiation Field (ISRF) and β the dust emissivity index –, if we assume a Mathis et al. (1983) radiation field, i.e. scaling with the inverse of R_G, we obtain that the clumps cold dust component appears to be a local measure of the global radiation field. On the contrary, the clumps warm dust component depends on the stellar radiation field, since this dust is located in proximity of the young massive stars. Remarkably, an O5 star will heat up dust up to 30K out to a radius of 0.75 pc from the star (Whitney et al. 2005). The trend that we find appears to mimic the well-known invertionally proportional increase of electron temperature, T_e, with galactocentric radius in H_{II} regions (Shaver et al. 1983, Paladini et al. 2004), which is consequence of the metallicity Galactic gradient, and of the fact that metals such as oxygen are coolants. Interestingly, though, we do not see a closer correlation when we split the sample in MYSOs and UCH II regions.

A dependance of dust temperature on Galactocentric distance is also reported by Urquhart et al. (2018). In their analysis, these authors do not distinguish between cold and warm temperature components, however they find increasing temperatures at larger galactocenctric distances.

Finally, we have analyzed potential differences between the warm/cold dust temperatures of candidate MYSOs and UCH II regions (Figure 4, bottom panel). We do not find any indication that these two populations are characterized by different temperatures. We find: $\overline{T_{c,MYSO}}$ = 20.5±3.0 K (same for the fitted size case), $\overline{T_{w,MYSO}}$ = 54.5±5.0 K (54.3±4.9 K for the fitted size case); $\overline{T_{c,UCH_{II}}}$ = 21.5±2.8 K (21.5±3.0 K for the fitted size case), $\overline{T_{w,UCH_{II}}}$ = 53.8±4.4 K (53.6±4.2 K for the fitted size case), where $T_{c,MYSO}$/$T_{c,UCH_{II}}$ and $T_{w,MYSO}$/$T_{w,UCH_{II}}$ are the cold/warm dust temperatures for the MYSOs/UCH II regions.

4.2. *Luminosities, masses and surface densities*

Having fitted the source SEDs, we can compute their luminosities. To this end, we use the relation:

$$L_{IR} = 4\pi D^2 \int_{\lambda_{min}}^{\lambda_{max}} S_\lambda d\lambda \tag{2}$$

where S_λ is the best-fit 2-temperature component model discussed in Section 4, and D is the solar distance from the RMS data base. The integration is performed between 25μm and 850μm, i.e. the shortest/longest IRIS/Planck wavelength considered in our analysis. The result is shown in Figure 5. We obtain $\overline{L_{IR}} = 8.3 \times 10^{5+1.4\times10^6}_{-6.5\times10^4}$ $L_\odot$ ($\overline{L} = 1.1 \times 10^{6+1.7\times10^6}_{-6.8\times10^4}$ $L_\odot$ for the fitted-sizes case). We find that sources in the inner Galaxy are on average more luminous than in the outer Galaxy, with $\overline{L_{IR,inner}}$= $8.9 \times 10^{5+1.5\times10^6}_{-6.9\times10^4}$ $L_\odot$ ($\overline{L_{IR,inner}}$ = $1.2 \times 10^{6+1.8\times10^6}_{-7.6\times10^4}$ $L_\odot$ for the fitted-sizes case) and $\overline{L_{IR,outer}}$ = $3.0 \times 10^{5\ +6\times10^5}_{-2.8\times10^4}$ $L_\odot$ ($\overline{L_{IR,outer}}$ = $3.6 \times 10^{5+6.2\times10^5}_{-3.1\times10^4}$ $L_\odot$ for the fitted-sizes case). We also look into differences between the average luminosity of clumps associated with MYSOs and UCH II regions and the latter turn out to be more luminous, with $\overline{L_{IR,UCHII}}$ = $9.6 \times 10^{5+1.6\times10^6}_{-7.7\times10^4}$ $L_\odot$ ($\overline{L_{IR,UCHII}}$ = $1.3 \times 10^{6+1.9\times10^6}_{-8.2\times10^4}$ $L_\odot$ for the fitted-sizes case) and $\overline{L_{IR,MYSO}}$ = $5.3 \times 10^{5+7.8\times10^5}_{-4.3\times10^4}$ $L_\odot$ ($\overline{L_{IR,MYSO}}$ = $6.7 \times 10^{5+8.7\times10^5}_{-4.4\times10^4}$ $L_\odot$ for the fitted-size case), which is consistent with their more advanced evolutionary stage. This is also similar to what has been reported by Urquhart et al. (2014) who analyzed a sample of $\sim$ 800 ATLASGAL clumps associated with $\sim$ 1,000 RMS sources. These authors compute their luminosities by using the model SED fitter developed by Robitaille et al. (2007) and by combining flux measurements from 2MASS (Skrutskie et al. 2006), UKIDSS (Lucas et al. 2008) or Vista-VVV (Minniti et al. 2010), MSX, WISE, Hi-Gal and ATLASGAL. They obtain that the luminosity distributions of the MYSO and UCH II regions subsamples are significantly different, with the two distributions peaking at $\sim 1\times10^4$ $L_\odot$ and $\sim 4\times10^4$ $L_\odot$,respectively. This difference in the luminosity function was first discussed in Mottram et al. (2011b). Note that the lower luminosities reported by Urquhart et al. (2014) with respect to our values, are likely due to the fact that the ATLASGAL clumps are smaller in size, i.e. on average $\sim$ 1.2 pc (see below for a comparison with our average size).

For the calculation of clump masses, assuming thermal equilibrium, we have:

$$M_{IR} = \frac{S_{850\mu m} D^2}{\kappa_{850\mu m} B_\nu(T_d)} \tag{3}$$

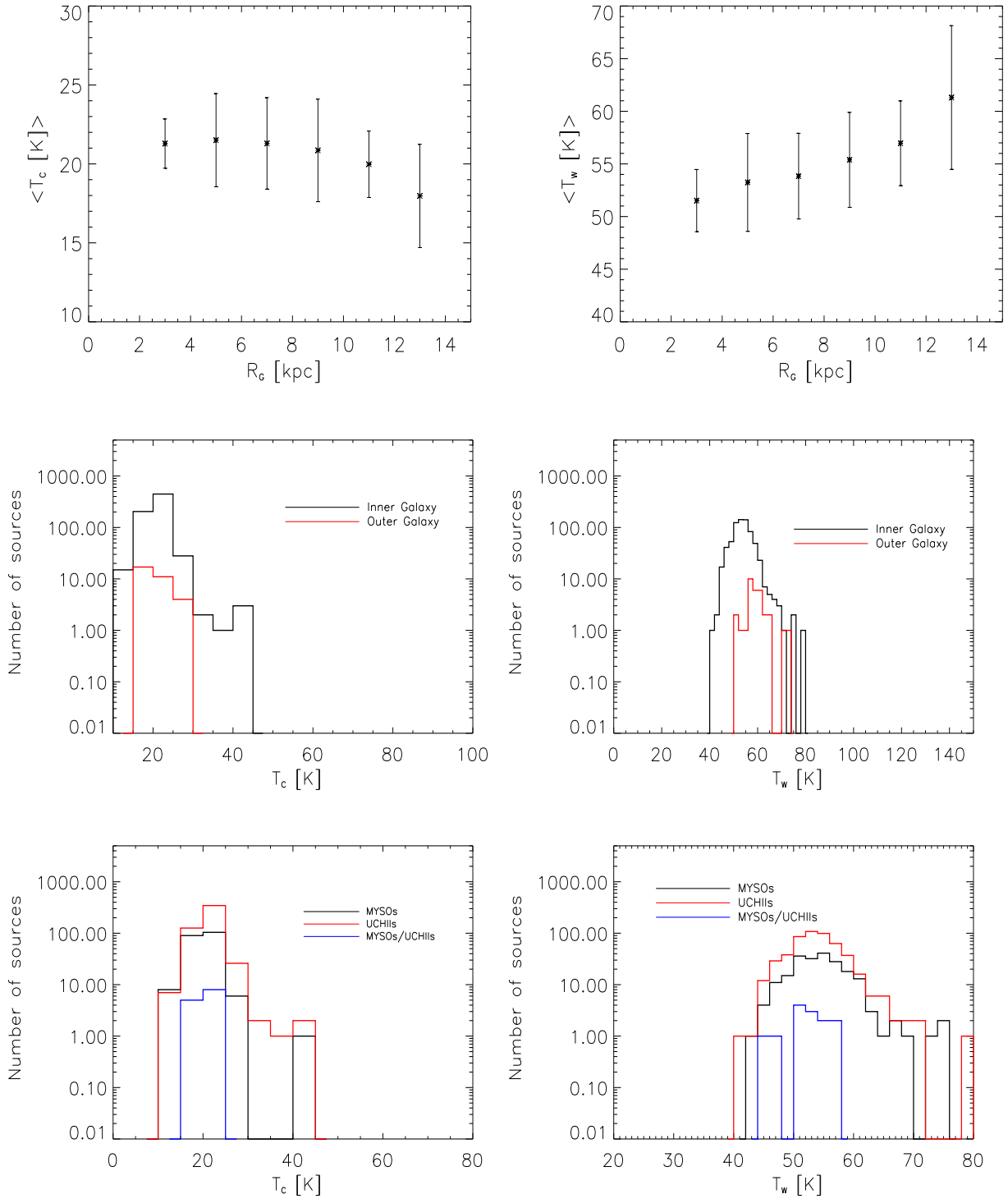


FIG. 4.— Top: Average cold (T_c) and warm (T_w) dust temperature components as a function of Galactocentric radius. R_G values are sampled in 2-kpc bins. Plotted error bars are the standard deviations in each bin. Middle: cold and warm dust temperature distributions for sources in the inner (first and fourth) and outer (second and third) Galactic quadrants. Bottom: cold and warm dust temperature distributions for MYSOs and UCH II regions.

TABLE 2
AVERAGE COLD (T_c) AND WARM (T_w) DUST TEMPERATURES PER GALACTOCENTRIC BIN FOR THE TWO CASES OF UNRESOLVED SOURCES AND FITTED DIAMETERS.

bin center (kpc)	n. sources	T_c (unresolved) (K)	T_w (unresolved) (K)	T_c (fitted sizes) (k)	T_w (fitted sizes) (K)
3	57	21.3±1.6	51.5±2.9	21.2±1.5	51.5±2.8
5	274	21.5±2.9	53.2±4.6	21.6±3.2	53.1±4.5
7	224	21.3±2.9	53.8±4.1	21.3±3.0	53.7±3.9
9	126	20.8±3.2	55.4±4.5	20.7±3.3	55.1±4.2
11	39	19.9±2.1	56.9±4.0	19.9±2.4	56.5±4.1
> 13	11	17.9±3.2	61.3±6.8	18.1±3.4	60.7±7.3

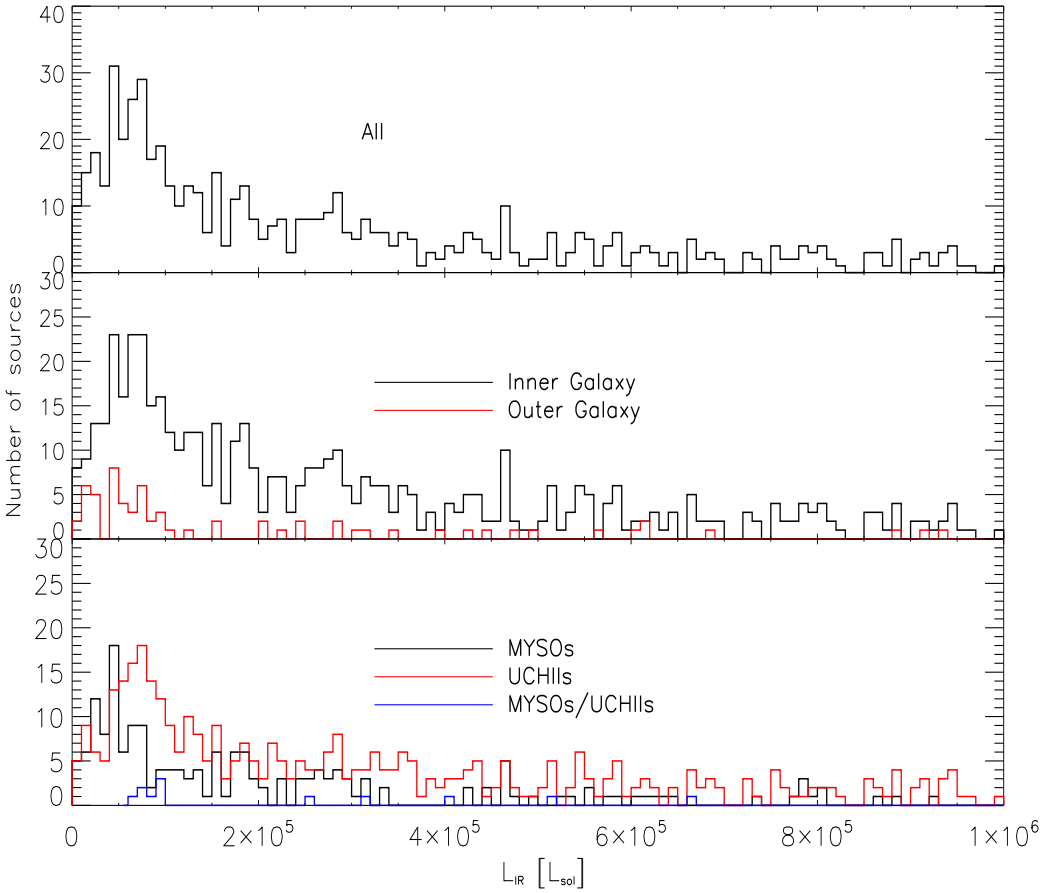


FIG. 5.— Top panel: Overall distribution of luminosities for the complete sample. Middle panel: inner and outer Galaxy luminosity distribution. Bottom panel: MYSOs, UCH II regions and MYSOs/UCH II luminosity distribution.

where the dust opacity, assuming a gas-to-dust ratio of 100, $\kappa_{850\mu m} = \kappa_{353\ GHz} = 0.012$ cm^2 g-1 (Preibisch et al. 1993), and $T_d = T_c$. Given that we only consider the cold dust component and ignore dust depletion effects in the source proximity, our computed mass will be an overestimate of the true values. Also, we emphasize that changing the value of the adopted dust opacity with affect the mass estimates. For instance, Ossenkopf & Henning (1994) give $\kappa_{700\mu m} = 0.0257$ cm^2 g-1, causing masses to be a factor of two lower. Finally, the dust-to-gas ratio, rather than being a fixed value across the Galaxy, varies with Galactocentric distance and this will introduce further uncertanties in mass determinations.

Figure 6 illustrates the overall mass distribution (top panel), the inner and outer Galaxy mass distribution (middle panel), and the mass distribution for MYSOs, UCH II regions and MYSOs/UCH II. The resulting average mass is $\overline{M_{IR}} = 7.7 \times 10^{4+1.3\times10^4}_{-0.9\times10^3}$ M$_\odot$ ($\overline{M_{IR}} = 1.1 \times 10^{4+1.8\times10^4}_{-1.0\times10^3}$ M$_\odot$ for the fitted-sizes case). The sources in the inner Galaxy are not only more luminous but also more massive compared to those in the outer Galaxy, in fact $\overline{M_{IR,inner}} = 8.2 \times 10^{3\ +1.4\times10^4}_{-1.15\times10^3}$ M$_\odot$ ($\overline{M_{IR,inner}} = 1.2 \times 10^{4+1.8\times10^4}_{-1.2\times10^3}$ M$_\odot$ for the fitted-sizes case) and $\overline{M_{IR,outer}} = 2.2 \times 10^{3+4.5\times10^3}_{-0.4\times10^3}$ M$_\odot$ ($\overline{M_{IR,outer}} = 3.3 \times 10^{3+5.3\times10^3}_{-0.6\times10^3}$ M$_\odot$ for the fitted-sizes cases). Finally, clumps associated with UCH II regions are more massive than clumps associated to MYSOs, with $\overline{M_{IR,UCHII}} = 8.5 \times 10^{3+1.4\times10^4}_{-1.1\times10^3}$ M$_\odot$ ($\overline{M_{IR,UCHII}} = 1.2 \times 10^{4+1.9\times10^4}_{-1.1\times10^3}$ M$_\odot$ for the fitted-sizes case) and $\overline{M_{IR,MYSO}} = 5.5 \times 10^{3+9.9\times10^3}_{-0.7\times10^3}$ M$_\odot$ ($\overline{M_{IR,MYSO}} = 6.8 \times 10^{3+1.2\times10^4}_{-0.9\times10^3}$ M$_\odot$ for the fitted-sizes case), again in agreement with the former being older objects, which have undergone completely the accretion process. A similar result is reported in Urquhart et al. (2014) based on the analysis of the sample of ATLASGAL clumps (see above). They find that the MYSO and UCH II regions associated clump distributions have a median value of $\sim$ 1000 M$_\odot$ and 2800 M$_\odot$, respectively. Note, once again, that the lower ATLASGAL values can be ascribed to the smaller clump sizes.

In the overall sample, we have identified a sub-sample of 15 very massive objects, i.e. with a mass $> 4\times10^4$ M$_\odot$. These are 3 MYSOs and 12 UCH II regions, which are concentrated towards $l \sim 43°$ and $l \sim 338°$, at a Galactocentric distance of, respectively, 7 kpc and 5 kpc.

The derived masses allow us to estimate the sources surface densities, $\Sigma = M_{IR} / (\pi\ r^2)$, where r is the linear radius. We obtain a mean surface density equal to $0.25^{+0.17}_{-0.003}$ g cm^{-2} (Figure 7). This mean value is greater than the characteristic surface densities of GMCs, $\Sigma \sim 0.035$ g cm^{-2} (Solomon et al. 1987), indicating that the clumps in our sample are gravitationally bound (Bertoldi & McKee 1992; Williams et al. 2000). At the same time this is lower than the average surface density (i.e. $\Sigma >= 1$ g cm^{-2}) found by Plume et al. (1997) for a sample of regions of massive star formation observed in the $J = 5 \rightarrow 4$ and $2 \rightarrow 1$ transitions of CS and C^{34}S. Importantly, Plume's sources have a mean linear radius and (virial) mass of 0.5 pc and 3800 M$_\odot$, respectively. Our sample is characterized by generally larger and more massive clumps, with a mean linear radius of 4.7$\pm$3.7 pc, and they do not represent the dusty counterparts of the molecular dense clumps studied by Plume et al. (1997): while Plume et al.'s observations trace the densest material directly involved in the production of massive stars, the *Planck* measurements are sensitive to the total mass involved, including less dense material. More recent studies (Butler & Tan 2012; Traficante et al. 2018, 2020) have revised down the massive star formation threshold value and today it is generally accepted that this is to be found in the range 0.1 - 0.35 g cm^{-2}, in agreement with the fact that our clumps are forming massive stars.

5. VARIATION OF THE STAR FORMATION EFFICIENCY WITH GALACTOCENTRIC RADIUS

It is well established that the outer Galaxy presents a very different environment with respect to the inner Galaxy, being the former characterized by low total gas densities, low metallicities and gas abundances, as well as by higher shear. Because of the Kennicutt (1998) law, the low densities makes the gas gravitationally stable, hence not apt to star formation. Despite these unfavorable environmental conditions, sporadic star formation is observed towards the outer Galaxy, suggesting that other mechanisms, other than gravity, may trigger the star formation process. In particular, Elmegreen & Hunter (2006) propose that some degree of turbulence persists in the outer disk, allowing the formation of clouds and compensating for the lack of gravitational instabilities.

A measure of the global *instantaneous* SFE can be obtained from the ratio of the IR luminosity, $L_{\rm IR}$, to clump mass, $M_{\rm IR}$. Both these quantities have the same square dependence on distance (D^2), therefore their ratio is distance independent. This is very important given that solar distances, especially when derived from kinematic measurements, can be affected by large uncertainties. Using the same binning in Galactocentric radii described in Section 4.1, we then investigate how the luminosity-to-mass ratio, $L_{\rm IR}/M_{\rm IR}$, varies from the inner to the outer Galaxy. For this purpose, we first analyze the $L_{\rm IR}/M_{\rm IR}$ distributions per Galactocentric bin (Figure 8). Then, we average the $L_{\rm IR}/M_{\rm IR}$ values in each bin, and look at the distribution of the mean values as a function of Galactocentric radius (Figure 9 and Table 3). It is apparent from Figure 8 that there is a significant dispersion of the $L_{\rm IR}$ to $M_{\rm IR}$ ratios in each bin, indicating that star formation does not simply scale with distance from the Galactic center, but instead is a complex process which depends on many parameters. If the determing factors in defining the star formation activity at a given location in the Galaxy were quantities which linearly scale with the Galactocentric radius, such as gas column density and metallicity, we would expect a much tighter correlation. Despite the large scatter in every bin, a global trend emerges from both Figure 8 and Figure 9: the mean luminosity-to-mass ratio decreases towards the outer Galaxy, with a difference between the most three inner bins (2 kpc $<$ R$_G$ $<$ 8 kpc) and the most three outer ones (8 kpc $<$ R$_G$ $<$ 14 kpc) of the order of 53%.

Similarly, if one considers the $L_{\rm IR}/M_{\rm IR}$ distributions for the *inner/outer Galaxy* sources, using the same convention for defining the two sub-samples described in Section 4.1, we obtain: $\overline{(L_{IR,i}/M_{IR,i})} = 119.1^{+177.6}_{-34.2}$ L$_\odot$/M$_\odot$ ($\overline{(L_{IR,i}/M_{IR,i})} = 118.9^{+169.4}_{-31.9}$ L$_\odot$/M$_\odot$ for the fitted-sizes case) and $\overline{(L_{IR,o}/M_{IR,o})} = 109.8^{+1150.8}_{-43.1}$ L$_\odot$/M$_\odot$ ($\overline{(L_{IR,o}/M_{IR,o})} = 94.9^{+135.3}_{-33.7}$ L$_\odot$/M$_\odot$ for the fitted-sizes case), where $L_{\rm IR,i}/M_{\rm IR,i}$ and $L_{\rm IR,o}/M_{\rm IR,o}$ are the IR luminosity to clump mass ratio for the *inner* and *outer Galaxy* samples, respectively (Figure 10).

If we look closely at Figure 8 and Figure 9, we notice that the SFE, instead of uniformly decreasing with R$_G$, first increases at small radii (R$_G$ $<$ 4 kpc) and then peaks at R$_G$ $\sim$ 5

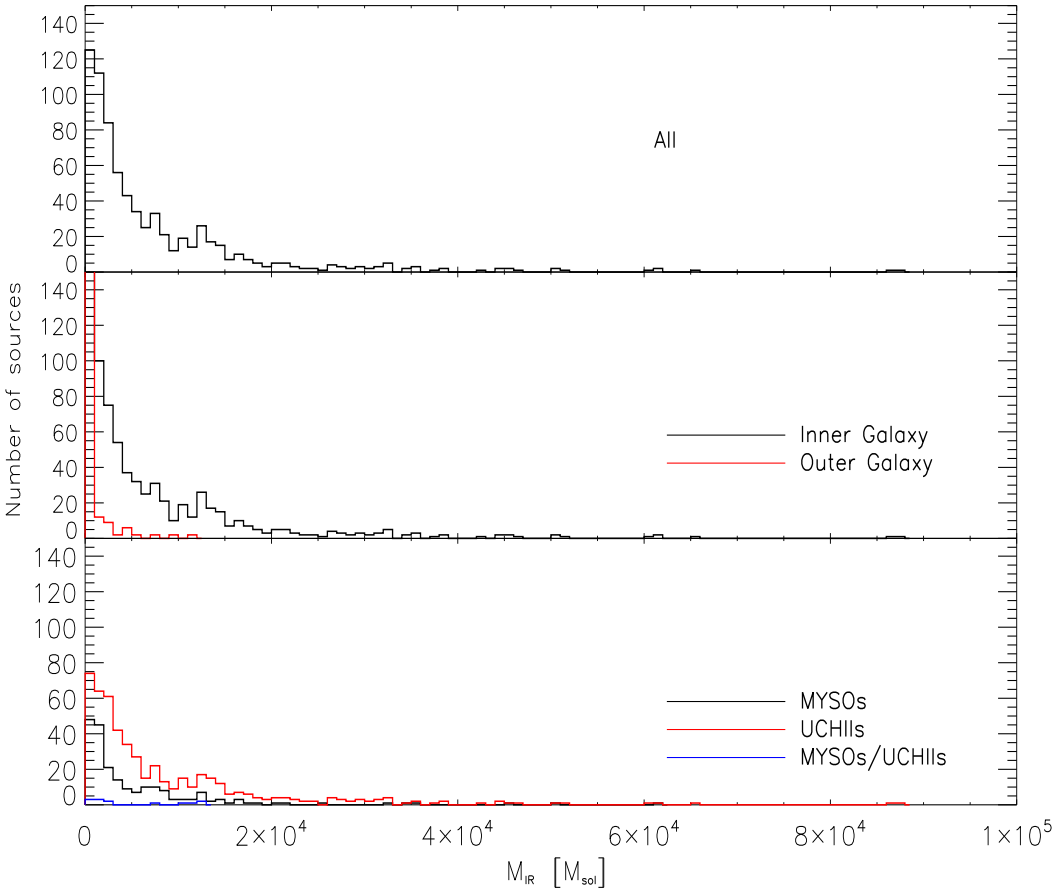


FIG. 6.— Top panel: Overall distribution of mass for the complete sample. Middle panel: inner and outer Galaxy mass distribution. Bottom panel: MYSOs, UCH II regions and MYSOs/UCH II mass distribution.

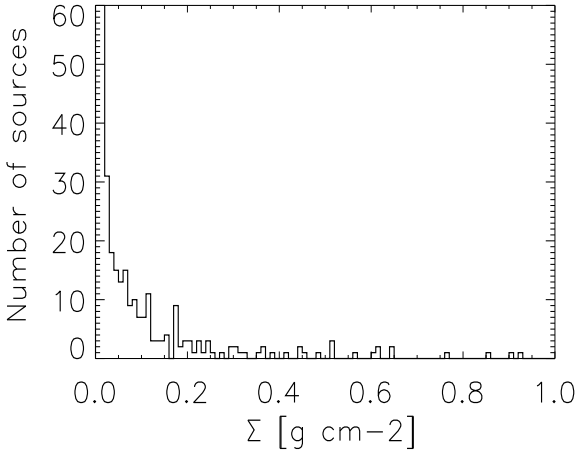


FIG. 7.— Surface density distribution for the clumps in the complete sample.

kpc, after which it linearly decreases. This behavior can overall be explained by considering the molecular fraction radial profile as discussed, for instance, in Koda, Scoville & Heyer (2016). These authors show that, from a quantitative point of view, the fraction of molecular gas decreases from 100% near the center of the Galaxy, dropping to 50% at 6 kpc, and decreasing to 10-20% at $R \sim 8.5$ kpc. The prominent bump at ~ 4.5 kpc is due to the bar structure of our Galaxy, with a half-length of 4.4±0.5 kpc (Benjamin et al. 2005), and this is associated to bright CO at the end of the bar, as observed in many external barred spiral systems (e.g. Sheth et al. 2002).

The correlation between SFE and molecular gas is a long-standing problem discussed in the literature, and even more so is the role played by spiral arms. Theoretical models (e.g. Dobbs et al. 2006) predict that spiral arms play a pivotal role in the formation of GMCs and enhance the SFE. However, so far we were lacking evidence of this. Previous efforts carried out by Moore et al. (2012) in the first Galactic quadrant found that the $L_{\rm IR}/M_{\rm IR}$ ratio was relatively flat for the inner 5 kpc. Two significant peaks were detected at ~ 6 kpc and ~ 8 kpc, respectively, but these were attributed to the presence of of the star formation complexes W51 snd W49. The work of Moore et al. (2012) was later extended by Eden et al. (2013; 2015) who reported no significant variation of the SFE between different arms or the inter-arm regions. More recently, Urquhart et al. (2018, 2020), in analyzing the ATLASGAL sample in the first and fourth quadrant, found that the SFE as described by the L_{bol}/M_{clump} ratio is relatively flat between 2 and 9 kpc when evaluated over kiloparsec scales, although local enhancements can be detected on smaller scales, in agreement with Moore et al. (2012) and Eden et al. (2013; 2015). This body of work led all these authors to conclude that: "*the spiral arms are principally collecting material together via orbit crowding but there is no evidence that they are play-*

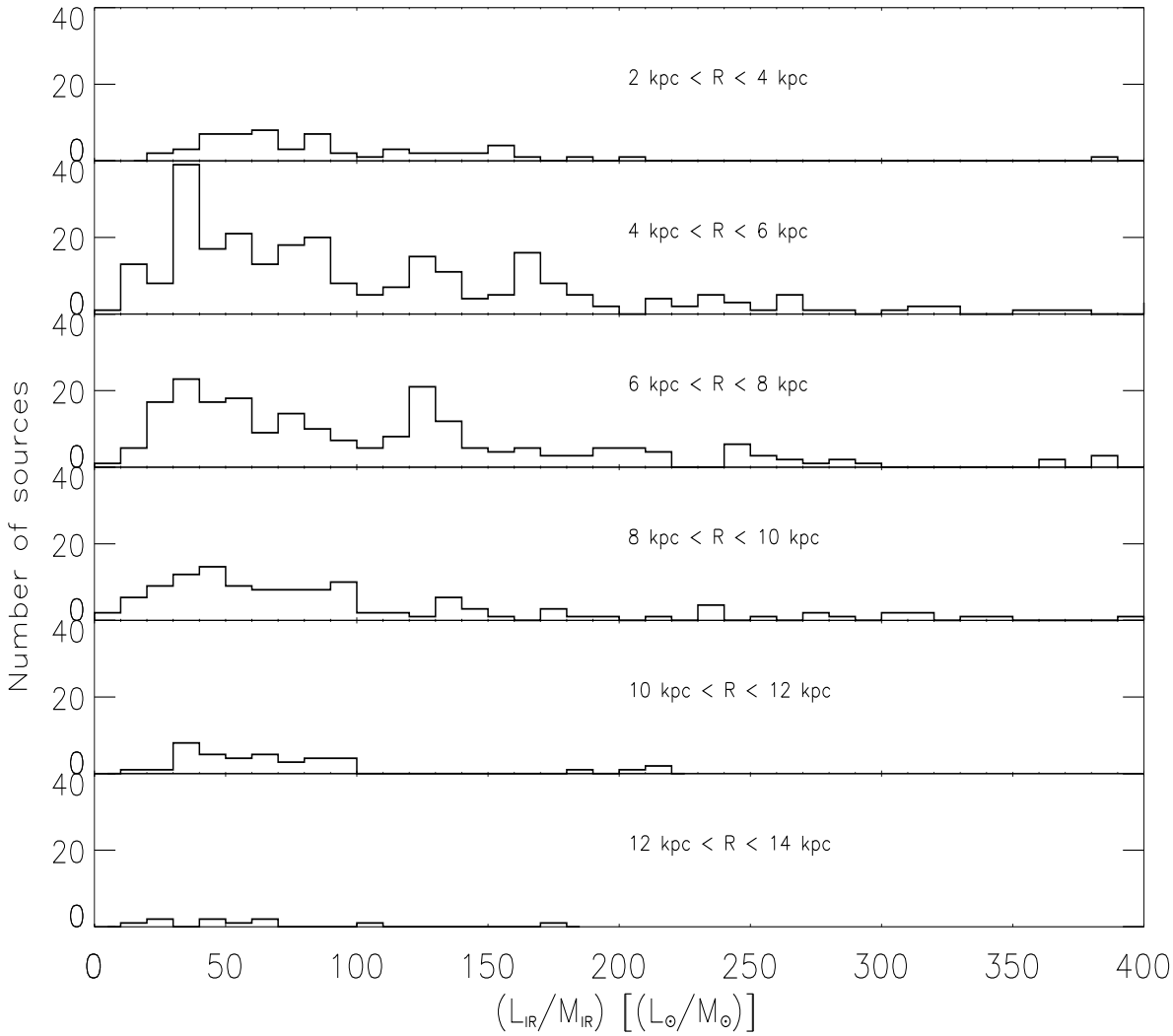


FIG. 8.— IR luminosity ($L_{\rm IR}$) to dust mass ($M_{\rm IR}$) ratio, $L_{\rm IR}/M_{\rm IR}$, as a function of Galactocentric radius, R_G.

ing a role in enhancing the star formation within molecular clouds."

We believe that we are detecting correlations of the SFE with Galactocentric distance, and with the overall distribution of molecular material, because, for the first time, we are combining three ingredients: 1) a complete sample of objects; 2) the exploration of the Galactic plane as a whole; 3) the use of data not affected by spatial filtering.

Figure 9 also shows that the peak at $R_G \sim 5$ kpc is not followed by a steady decrease in SFE. Instead, the IR luminosity, $L_{\rm IR}$, to clump mass, $M_{\rm IR}$, ratio remains almost constant up to ~ 9 kpc. A best-fit to $L_{\rm IR}/M_{\rm IR}$ as a function of Galactocentric radius gives:

$$L_{IR}/M_{IR} = \begin{cases} 28.3 + 21.1 \times R_G & R_G < 5 \text{ kpc} \\ 152.8 - 4.2 \times R_G & 5 \text{ kpc} < R_G < 9 \text{ kpc} \\ 232.6 - 13.5 \times R_G & R_G > 9 \text{ kpc} \end{cases} \quad (4)$$

Importantly, the best-fit solution is obtained by folding in the spread in the data points in each bin. We notice that the two last bins contains only few data points, i.e. 39 and 11, respectively. However, since the sample is complete, these bins are significant and their weight in the fit is the same as the other bins.

Leroy et al. (2008), based on a study of 23 external spiral and dwarf galaxies, find that in spiral galaxies the SFE appears approximately constant up to the optical radius (r_{25}), after which it declines exponentially. Taking into account that for our Galaxy the optical radius is estimated to be at ~ 13 - 19 kpc (Ruffle et al 2007; Sale et al. 2010), the Leroy et al.'s relation for the Milky Way can be re-written as:

$$SFE \propto L_{IR}/M_{IR} = \begin{cases} const & R_G < 5.6 - 8.2 \text{ kpc} \\ \sim e^{-R_G/(3.25\ -\ 4.75 \text{ kpc})} & R_G > 5.6 - 8.2 \text{ kpc} \end{cases} \quad (5)$$

Figure 9 illustrates the linear fit to the data (equation 4) and the Leroy et al.'s exponential solutions for two values of the optical radius (equation 5). Leroy et al. point out that the ISM is equal parts of HI and H2 at $R_G = 0.4 \times r_{25} \pm 0.18$. Therefore their result can be interpreted as a steady decrease of the SFE beyond the HI-to-H2 transition radius. However, as we see from Figure 9, in the case of our Galaxy the SFE re-

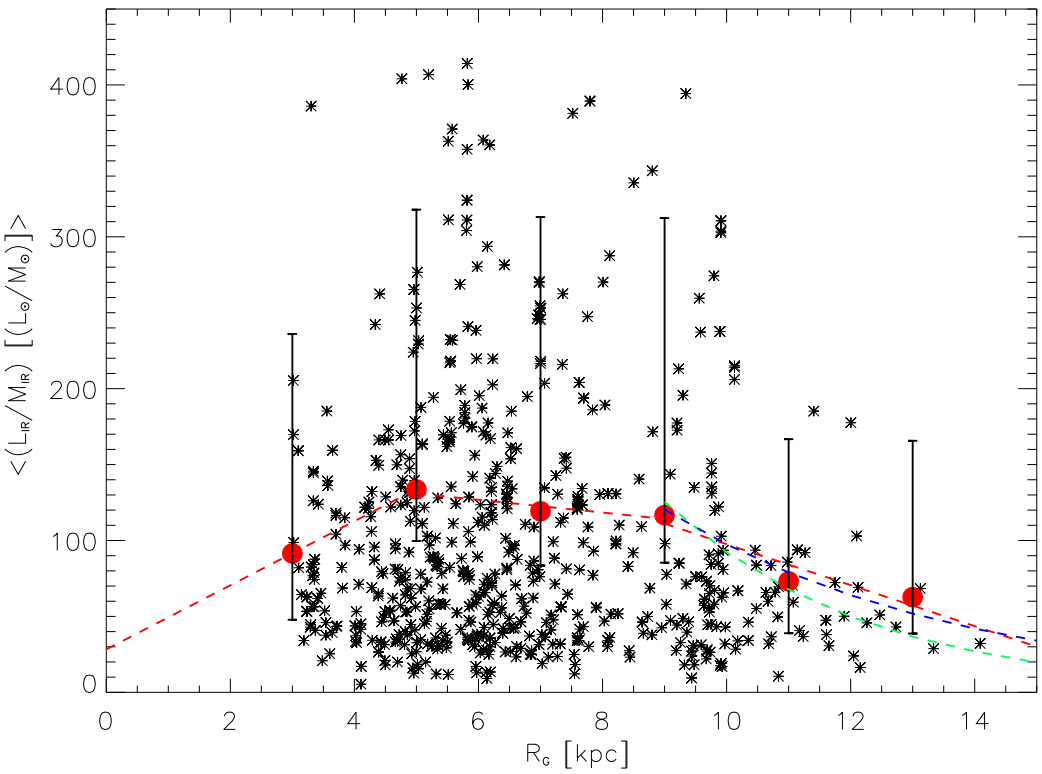


FIG. 9.— IR luminosity ($L_{\rm IR}$) to dust mass ($M_{\rm IR}$) ratio, $L_{\rm IR}/M_{\rm IR}$, as a function of Galactocentric radius, R_G. The error bars are the 15.9% and 84.1% percentiles. The best-fit to the data is shown (red-dashed line). Also shown is the exponential decay from Leroy et al. for the case of an optical radius of 13 kpc (green-dashed line) and 19 kpc (blue-dashed line). The stars represent the individual data points.

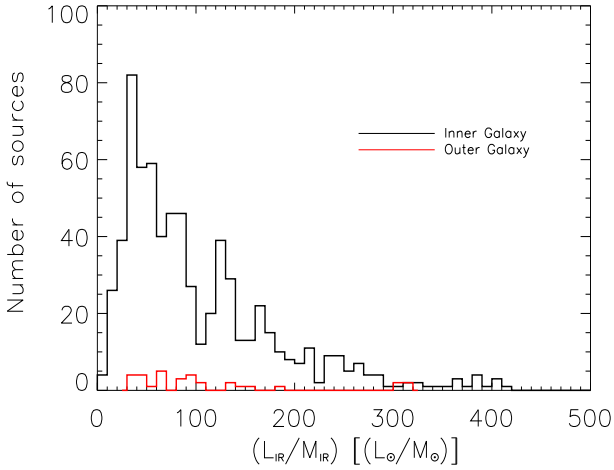


FIG. 10.— Luminosity-to-mass ratio in the inner (first and fourth) and outer (second and third) Galactic quadrants.

mains fairly constant up to $R_G \sim$ 9- 10 kpc, suggesting either the presence of a significant, previously undetected, molecular material at $R_G > 8$ kpc and/or that the radial dependence of the SFE may not be simply a function of the HI-to-H2 transition radius as previously thought.

From the observational point of view, the former finds at least partial confirmation in the hydrogen recombinations lines (H86α to H93α) and radio continuum (9 GHz) survey carried out by Anderson & Bania (2011) with the Green Bank Telescope (GBT). These observations have allowed the discovery, in the first Galactic quadrant, of a new population of 34 HII regions located in the Outer Arm (Bania et al. 2010), at $R_G > 8$ - 9 kpc. Previously, there were only 7 HII regions known in this area.

Finally, we have investigated variations of the SFE as a function of Galactocentric radius for the two populations of MYSOs and UCH II regions, separately. A behavior similar to the parent population is found for each of these evolutionary stages (see Table 3).

6. CONCLUSIONS

Planck whole-sky sub-millimeter maps provide an unprecedented opportunity to carry out an unbiased study of the environmental conditions in which massive star formation takes place. Using *Planck* HFI upper frequency bands (350, 500, 850 μm) complemented by IRAS/IRIS 100, 60 and 25μm bands, we have estimated dust temperatures, luminosities, masses and surface densities of clumps associated to a complete sample of candidate MYSOs and UCH II regions selected from the RMS survey data base.

Exploting the *Planck* and IRAS/IRIS full-coverage of the Galactic Plane, we have searched for variations of the clump dust temperatures (warm and cold component) with respect to distance from the Galactic center. We find that the distribution of the cold dust temperature component of the clumps generally reflects the large scale variation of dust temperature for the ISM along the Galactic Plane, that is: massive star complexes at smaller Galactocentric radii have higher tem-

TABLE 3
AVERAGE L_{IR}/M_{IR} PER GALACTOCENTRIC BIN.

bin center (kpc)	n. sources	Log(L_{IR}/M_{IR}) (unresolved) Log($L_\odot/M_\odot$)	Log(L_{IR}/M_{IR}) (fitted sizes) Log($L_\odot/M_\odot$)	Log(L_{IR}/M_{IR}) (UCH II) Log($L_\odot/M_\odot$)	Log(L_{IR}/M_{IR}) () Log($L_\odot/M_\odot$)
3	57	$91.5^{+144.5}_{-43.7}$	$83.2^{+126.2}_{-44.0}$	$81.4^{+176.8}_{-39.3}$	$83.6^{+126.2}_{-44.6}$
5	274	$133.6^{+184.2}_{-33.9}$	$141.8^{+183.5}_{-33.1}$	$104.6^{+163.2}_{-28.6}$	$157.9^{+203.5}_{-33.2}$
7	224	$119.2^{+193.8}_{-35.8}$	$115.6^{+174.1}_{-31.9}$	$93.2^{+152.7}_{-28.9}$	$124.1^{+180.9}_{-34.2}$
9	126	$116.7^{+195.6}_{-31.3}$	$107.3^{+208.9}_{-30.7}$	$88.5^{+171.6}_{-23.2}$	$113.9^{+223.9}_{-31.9}$
11	39	$73.1^{+93.6}_{-34.2}$	$65.5^{+84.8}_{-30.7}$	$52.3^{+77.7}_{-28.3}$	$72.9^{+162.3}_{-33.7}$
>13	11	$62.7^{+102.9}_{-23.9}$	$57.5^{+107.2}_{-23.9}$	$62.8^{+145.8}_{-16.3}$	$52.2^{+107.2}_{-23.9}$

peratures than complexes at larger radii. An opposite trend is found for the warm dust component that appears to increase with distance from the center of the Galaxy, suggesting that the more embedded dust is subjet to the local conditions rather than to the ISRF.

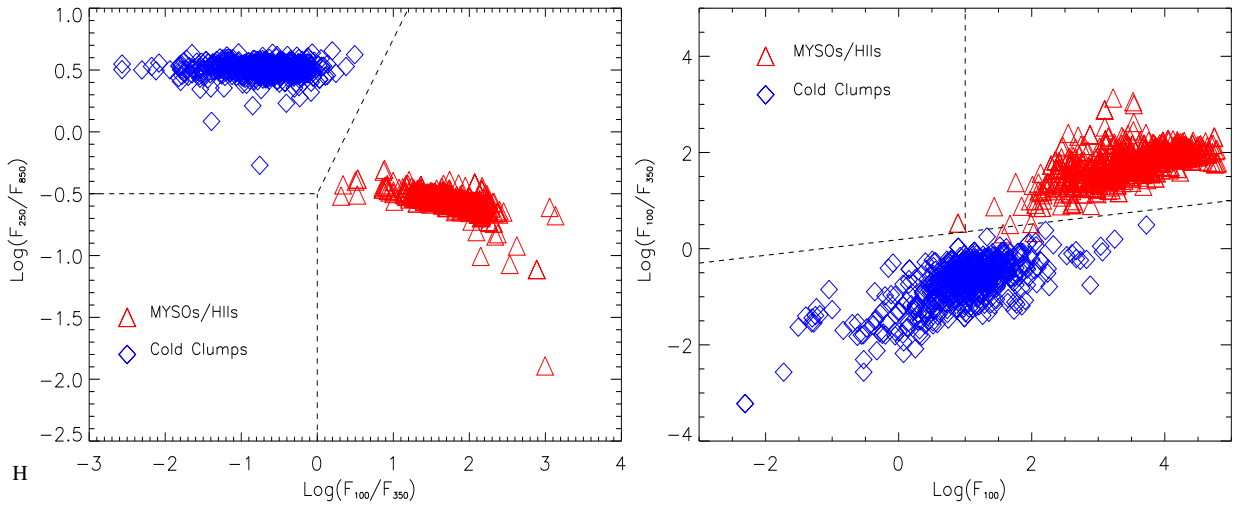
We have also explored how the luminosity-to-mass ratio (L/M) associated with the clumps varies from the inner to the outer Galaxy. Our results are consistent with a SFE decreasing from the inner to the outer Galaxy. However, the cut-off radius that identifies the transition between a fairly constant and a declining SFE is at a larger radius ($R_G \sim 9$ kpc) than expected. This may have implications for the amount of molecular gas still undetected at $R_G > 8$ kpc.

Finally, we present in the Appendix color-color and color-magnitude plots of MYSOs and UCH II regions compared to cold clumps from the PGCC. These colors will be useful for identification of sites of massive star formation in future large-scale surveys.

This paper is based and made use of information from the Red MSX Source survey database at www.ast.leeds.ac.uk/RMS which was constructed with support from the Science and Technology Facilities Council of the UK.

REFERENCES

Anderson, L. D., Bania, T. M., Balser, D. S., Rood, R. T., 2011, ApJS, 194, 32
Bania, T. M., Anderson, L. D., Balser, D. S. & Rood, R. T., 2010, ApJ, 718, 106
Benjamin, R. A., Churchwell, E., Babler, B. L., et al., 2005, ApJ, 630, 149
Bertoldi, F. & McKee, C. F., 1992, ApJ, 395, 140
Bonnel, I. A., Bate, M. R., Zinnecker, H., 1998, MNRAS, 298, 93
Bonnel, I. A., Bate, M. R., Clarke, C. & Pringle, J. E., 2001, MNRAS, 323, 785
Bonnel, I. A., Clarke, C., Bate, M. R. & Pringle, J. E., 2001, MNRAS, 324, 573
Butler, M. J. & Tan, J. C., 2012, ApJ, 754, 5
Chini, R., Kreysa, E., Kruegel, E., Mezger, P. G., 1986, A&A, 166, 8
Chini, R., Kreysa, E., Mezger, P. G., et al., 1986a, A&A, 154, 8
Chini, R., Kreysa, E., Mezger, P. G., et al., 1986b, A&A, 157,1
Clarke, A. J., Lumsden, S. L., Oudmaijer, R. D., et al., 2006, A&A, 457, 183
Csengeri, T. Weiss, A., Wyrowski, F., et al., 2016, A&A, 585, 104
Dame, T. M., Hartmann, D., Thaddeus, P., 2001, ApJ, 547, 792
Dame, T. M. & Thaddeus, P., 2011, ApJ, 734, 24
Digel, S. W., Grenier, I. A., Heithausen, A., et al., 1996, ApJ, 463, 609
Dobbs, C. L., Bonnell, I. A., Pringle, J. E., 2006, MNRAS, 371, 1663
Eden, D. J., Moore, T. J. T., Morgan, L. K., et al., 2013, MNRAS, 431, 1587
Eden, D. J., Moore, T. J. T., Urquhart, J. S., et al., 2015, MNRAS, 452, 289
Egan, M. P., Price, S. D., Kraemer, K. E., 2003, BAAS, 35, 1301
Elmegreen, B. G. & Hunter, D. A., 2006, ApJ, 636, 712
Giannetti, A., Brand, J., Sanchez-Monge, A., et al., 2013, A&A, 556, 16
Gorski, K. M., Hivon, E., banday, A. J., et al., 2005, ApJ, 622, 759
Hauser, M. G., Arendt, R. G., Kelsall, T., et al., 1998, ApJ, 508, 25
eyer, M. H., Brunt, C., Snell, R. L., et al., 1998, ApJS, 115, 241
Juvela, M. J., et al., 2013, A&A, 556, 63
Juvela, M. J., et al., 2018, A&A, 612, 71
Kennicutt, R. C., Edgard, B. Kevin & Hodge, P. W., 1989, ApJ, 337, 761
Koda, J., Scoville, N. Heyer, M., 2016, ApJ, 823, 76
Krumholz, M. R. & McKee, C. F., 2008, Nature, 451, 1082
Lamarre, J., Puget, J.-P., Ade, P. A. R., 2010, A&A, 520, 9
Larson, R. B., 2005, MNRAS, 359, 211
Leahy, J. P., Bersanelli, M., D'Arcangelo, O., et al., 2010, A&A, 520, 8
Leroy, A., Walter, F., Brinks, E., et al., 2008, AJ, 136, 2782
Lucas, P. W., et al., 2008, MNRAS, 391, 136
Lumsden, S. L., Hoare, M. G., Oudmaijer, R. D. & Richards, D., 2002, MNRAS, 336, 621
Lumsden, S. L., Hoare, M. G., Urquhart, J. S., et al., 2013, ApJS, 208, 11
Mandolesi, N., Bersanelli, M., Butler, R. C., et al., 2010, A&A, 520, 3
Mathis, J. S., Mezger, P. G., Panagia, N., 1983, A&A, 128, 212
McClure-Griffiths, N. M., Dickey, J. M., Gaensler, B. M., et al., 2005, ApJS, 158, 178
McKee, C. F. & Tan, J. C., 2003, ApJ, 585, 850
Minniti, D., et al., 2010, New Astron, 15, 433
Miville-Deschenes, M. A. & Lagasche, G., 2005, ApJSS, 157, 302
Miville-Deschenes, M. A., et al., 2017, ApJ, 834, 57
Molinari, S., Pezzuto, S., Cesaroni, R., et al., 2008, A&A, 481, 345
Molinari, S., Swinyard, B., Bally, J., et al., 2010, A&A, 518, 100
Moore, T. J. T., Urquhart, J. S., Morgan, L. K., et al., 2012, MNRAS, 426, 701
Mottram, J. C., Hoare, M. G., Urquhart, J. S., et al., 2011, A&A, 525, 149
Neugebauer, G, Habing, H. J., van Duinen, R., et al., 1984, ApJ, 278, 1
Ossenkopf, V & Henning, T, 1994, A&A, 291, 943
Padoan, P. & Nordlund, A., 2002, ApJ, 576, 870
Paladini, R., Davies, R. D., De Zotti, G., 2004, MNRAS, 347, 237
Paladini, R., De Zotti, G., Noriega-Crespo, A., Carey, S. J., 2009, ApJ, 702, 1036
Pilbratt, G. L., Rieding, J. R., Passvogel, T., et al., 2010, A&A, 518, 1
Planck Collaboration, 2011a, A&A, 536, 1
Planck Collaboration, 2011m, A&A, 536, 17
Planck HFI Core Team, 2011a, 536, 4
Planck HFI Core Team, 2011b, 536, 6
Planck Collaboration. I., 2016, 594, 1
Plume, R., Jaffe, D. T., Evans, N. J., et al., 1997, ApJ, 476, 730
Reid, M. J. et al., 2009, ApJ, 700, 137
Robin, A. C., Creze, M., Mohan, V., 1992, ApJ, 400, 25
Robin, A. C., Reyle, C., Derriere, S., Picaud, S., 2003, A&A, 409, 523
Robitaille, T. P., Whitney, B. A., Indebetouw, R., et al., 2006, ApJS, 167, 256
Robitaille, T. P., Whitney, B. A., Indebetouw, R. & Wood, K., 2007, ApJS, 169, 328
Rosset, C., Tristram, M., Ponthieu, N., et al., 2010, A&A, 520, 13
Ruffle, P. M. E., Millar, T. J., Roberts, H., et al., 2007, ApJ, 671, 1766
Ruphy, S., Robin, A. C., Epchtein, N., et al., 1996, A&A, 313, 21
Sale, S. E., Drew, J. E., Knigge, C., et al. 2010, MNRAS, 402, 71330
Scoville, N. Z. & Good, J. C., 1989, ApJ, 339, 149
Sheth, K., Vogel, S. N., Regan, M. W., et al., 2002, AJ, 124, 2581
Schnee, S., Kauffmann, J., Goodman, A., Bertoldi, F., 2007, ApJ, 657, 838
Schuller, F., Menten, K. M., Contreras, Y., et al., 2009, A&A, 504, 415
Shu, F. H., Adams, F. C., Lizano, S., 1987, ARAA, 25, 23
Skrutskie, M. F., Cutri, R. M., Stiening, R., et al., 2006 AJ, 131, 1163
Solomon, P. M., Rikvolo, A. R., Barrett, J. & Yahil, A., 1987, ApJ, 319, 730
Tauber, J. A., Mandolesi, N., Puget, J.-P., et al., 2010, A&A, 520, 1
Thilker, D. A., Bianchi, L., Meurer, G., et al., 2007, ApJS, 173, 538
Traficante, A., et al., 2018, MNRAS, 473, 4975
Traficante, A., et al., 2020, MNRAS, 491, 4310
Urquhart, J. S., Busfield, A. L., Hoare, M. G., et al., 2007a, A&A, 461, 11
Urquhart, J. S., Busfield, A. L., Hoare, M. G., et al., 2007b, A&A, 474, 891
Urquhart, J. S., Hoare, M. G., Lumsden, S. L., et al., 2008b, *Massive Star Formation: Observations Confront Theory*, ed. H. Beuther, H. Linz & T. Henning, ASP Conf. Ser., 387, 381



H

FIG. 11.— Color-color (left) and color-magnitude (right) plots for MYSOs and UCH II regions from the RMS database (red triangles) and cold clumps from the PGCC (blue triangles).

Urquhart, J. S., Hoare, M. G., Purcell, C. R., et al., 2009, A&A, 501, 539
Urquhart, J. S., Moore, T. J. T., Hoare, M. G., et al., 2011, MNRAS, 410, 1237
Urquhart, J. S., Hoare, M. G., Lumsden, S. L., et al., 2012, MNRAS, 420, 1656
Urquhart, J. S., Thompson, M. A., Moore, T. J. T., et al., 2013, MNRAS, 435, 400
Urquhart, J. S., Moore, T. J. T., Csengeri, T., et al., 2014, MNRAS, 443, 1555
Urquhart, J. S., Konig, C., Giannetti, A., et al., 2018, MNRAS, 473, 1059
Urquhart, J. S., Figura, C, Cross, J. R., et al. 2020, astro-ph/
Whitney, B. A., Robitaille, T. P., Indebentouw, R., et al., 2005, *Massive Star Birth: A Crossroads of Astrophysics*, Proceedings IAU Symposium N0. 227, 2005, ed. Cesaroni, R., Felli, M., Churchwell, E. & Walmsley, C. M.
Williams, J. P., Blitz, L. & McKee, C. F., 2000, *Protostars & Planets IV*, ed. V. mannings, A. P. Boss, & S. S. Russell, Tucson, Univ. Arizona Press, 97
Wright, E. L., Eisenhardt, P. R. M, Mainzer, A. K., et al., 2010, AJ, 140, 1868
Yong, D., Carney, B. W., Teixera de Almeida, M. L., 2005, AJ, 130, 597
Zhang, Y. & Tan, J. C., 2011, ApJ, 733, 55

APPENDIX

The photometric measurements described in Section 4, as well as allowing us to derive the physical properties of the sources presented above, also make it possible to generate color-color and color-magnitude plots, the purpose of which is to serve as diagnostic tools for the identification of MYSOs and UCH II regions in blind Galactic surveys. To this end, we have performed a random selection from the *Planck* Catalog of Galactic Cold Clumps (PGCC, Planck Collaboration XXVIII 2016) and compared the colors/magnitudes of the randomly extracted cold clumps to those of the complete sample of MYSOs and UCH II regions from the RMS database. The PGCC contains 13,188 Galactic cold sources (median temperature between 13 and 14.5 K) spread across the whole sky.

Figure 11 shows that the MYSOs and UCH II regions (red triangles) and cold clumps (blue triangles) occupy very distint regions in the color-color and color-magnitude plots. The MYSOs and UCH II regions are found for:

$$-0.5 < \mathrm{Log}(F_{100}/F_{350}) < 4., \; -2.5 < \mathrm{Log}(F_{250}/F_{850}) < 0.$$

The two populations are characterized by different dust temperatures that cause this segregation: as we have discussed in Section 4.1, there are two temperature dust components to the the MYSOs and UCH II regions SED, one warm, of the order of 30 - 50 K, and one cold, around 20 K. On the contrary, cold clumps present only one dust temperature component, i.e. the cold one. This implies that the SED of MYSOs and UCH II regions declines rapidly between 250 and 850 μm and between 100 and 350 μm, while in these wavelength ranges, the spectrum of a cold clump decreases less rapidly ($250 < \lambda < 750$ μm) or even still rises ($100 < \lambda < 350$ μm).